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# ChiantiPy Documentation

*Release 0.9.5*

**Ken Dere**

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Welcome to the ChiantiPy documentation. ChiantiPy is a pure Python package for performing calculations of astrophysical spectra using the [CHIANTI atomic database](#).

The latest version of ChiantiPy is 0.9.5 and is compatible with CHIANTI database version 9.0. It is not compatible with previous versions

ChiantiPy v0.9.5 is released under the OSI approved ISC license. From [Wikipedia](#): The ISC license is a permissive free software license written by the Internet Software Consortium (ISC). It is functionally equivalent to the simplified BSD and MIT/Expat licenses, ...

[CHIANTI](#) consists of a database of atomic data that can be used to interpret spectral lines and continua emitted from high-temperature, optically-thin astrophysical sources.

CHIANTI is developed and maintained by scientists at George Mason University (USA), the University of Michigan (USA), and the University of Cambridge (UK). The first version of CHIANTI was released in 1997 and version 9.0 in 2019.



## GETTING STARTED WITH CHIANTIPY

### 1.1 Prerequisites

- [CHIANTI](#), the atomic database for astrophysical spectroscopy (Version 8)
- [Python3](#) (3.6)
- [Numpy](#) (currently developed with 1.14)
- [Scipy](#) (currently developed with 1.00)
- [Matplotlib](#) (currently developed with 2.2.2)
- Matplotlib requires a GUI library

#### [PyQt5](#)

Once one of these is installed, it must be set as the backend in your matplotlibrc file, e.g., backend: Qt5Agg

- [ipyparallel](#) (required for multiprocessing with [ipynb](#))
- (not really a prerequisite but **extremely** useful) [IPython](#) version 6 and [Jupyter](#)

### 1.2 Install the CHIANTI database

The gzipped *data* tar ball can be downloaded from the [CHIANTI website](#)

- put the file in a convenient directory, cd to the directory and untar the file
- ChiantiPy uses the environment variable *XUVTOP* to find the database. Set XUVTOP to the name of the directory where the CHIANTI data tarball was placed. For example

```
setenv XUVTOP /data1/xuv/directory.where.the.tarball.was.placed
```

or on Windows: To set the environment variable, go to Control Panel -> System -> Advanced System Properties -> Environment Variables.

Some sites have the CHIANTI database maintained as part of a SolarSoft distribution. In that case, simply set XUVTOP to the directory where it resides, usually something like \$SSW/packages/chianti/dbase

## 1.3 Install the Prerequisites

On **Linux** systems this can usually be done with your package manager.

On **Windows**, **Linux** and **Mac** systems, it is possible to use

- the [Anaconda](#) distribution from from Continuum, or,
- the [Canopy](#) distribution from Enthought.

On **Windows**, it is also possible to use:

- [WinPython](#)

All of these packages are free, at least for noncommercial use (I believe) and have a considerable amount of documentation. **You should check the version of IPython that is provided.**

## 1.4 Install the ChiantiPy package

In order to be compatible with the latest version (8.0) of the CHIANTI atomic database, it is necessary to install the latest version (0.6.0) of ChiantiPy

```
pip install ChiantiPy
```

or

```
pip3 install ChiantiPy
```

I have not tried this with ChiantiPy, myself.

The ChiantiPy package can be downloaded from the [ChiantiPy](#) project page at Sourceforge, untar it, cd to the directory where it was unpacked, and then, as root

```
python setup.py install
```

If you do not have root privileges, simply put the ChiantiPy directory in your PYTHONPATH

```
python setup.py install --prefix=somewhere_in_my_PYTHONPATH
```

or on a Mac, with the Anaconda package

```
python setup.py install --prefix=/Users/your_user_name/anaconda/
```

**Thanks to Peter Young (GMU) for providing the instructions for installation on Mac and Windows**

## 1.5 Note - ChiantiPy interactions with various GUI backends

First, Matplotlib requires a GUI backend and can be specified by the user in the matplotlibrc file. Matplotlib expects to find this file in \$HOME/.config/matplotlib, although it might require that you copy it to that directory.

ChiantiPy also uses a set of GUI dialog widget set. Selections can also be made via the command line. The user choice is specified in the \$HOME/.chianti/chiantirc file if one has been copied there. Otherwise, the default GUI is to use the command line. A default chiantirc file is included with the distribution.

In order for the ChiantiPy dialog widget to be used, a backend for them must be initiated. If you choose the same backend for matplotlib (PyQt5 is best) as for the ChiantiPy widgets, then running `%matplotlib` in an IPython or Jupyter



session will do the trick. In an interactive Python session, invoking a matplotlib command first should also do the trick.

```
matplotlib inline
```

or

```
matplotlib qt
```

if you are using Qt5

If you choose to use a GUI backend other than that used for matplotlib, then in an IPython or a Jupyter command the following magic commands are also available to start the backend:

```
%gui qt
```

ChiantiPy has mostly been tested with the Qt5 backend for Matplotlib and using the ChiantiPy Qt5 widgets.



## QUICK START

This short tutorial will demonstrate some of the capabilities of ChiantiPy and the CHIANTI database. It assumes that you know what the CHIANTI database provides and why you want to use it. It is useful to begin by exploring the properties of the **ion class**, as much of ChiantiPy is based on it. An ion such as Fe XIV is specified by the string 'fe\_14', in the usual CHIANTI notation.

Perhaps the easiest way is to use a jupyter-notebook or a jupyter3-notebook to load the quick start notebook file QuickStart.ipynb in the directory ipython\_notebooks. Then, just run each cell step by step. If you are not familiar with notebooks, then you can cut and paste the following code into a Python/IPython session.

N.B.: in the time some of the plots and data were produced, there have been some changes to ChiantiPy and CHIANTI. It is possible that you might find difference (hopefully small).

Bring up a Python session (using `> Python -i`), or better yet, an IPython session

```
import ChiantiPy.core as ch
import numpy as np
import matplotlib.pyplot as plt
```

What we will really be interested in are various properties of the Fe XIV emissivities as a function of temperature and density. So, let's define a numpy array of temperatures

```
t = 10.** (5.8 + 0.05*np.arange(21.))
```

In ChiantiPy, temperatures are currently given in degrees Kelvin and densities as the number electron density per cubic cm. Then, construct fe14 as would be typically done

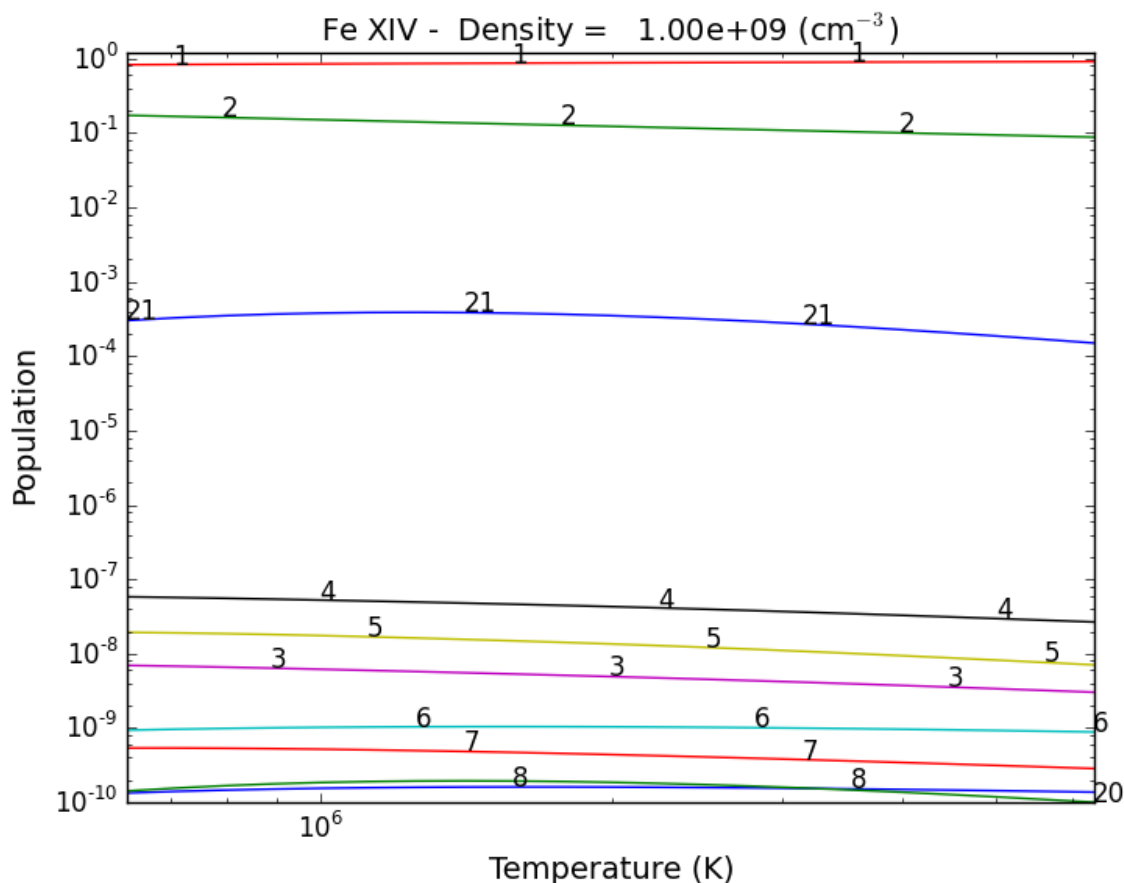
```
fe14 = ch.ion('fe_14', temperature=t, eDensity=1.e+9, em=1.e+27)
```

note that eDensity is the new keyword for electron density

### 2.1 Level populations

```
fe14.popPlot()
```

produces a matplotlib plot window where the population of the top 10 (the default) levels are plotted as a function of temperature.



If the level populations had not already been calculated, `popPlot()` would have invoked the `populate()` method which calculates the level populations and stores them in the Population dictionary, with keys = ['protonDensity', 'population', 'temperature', 'density'].

## 2.2 A ChiantiPy Convention

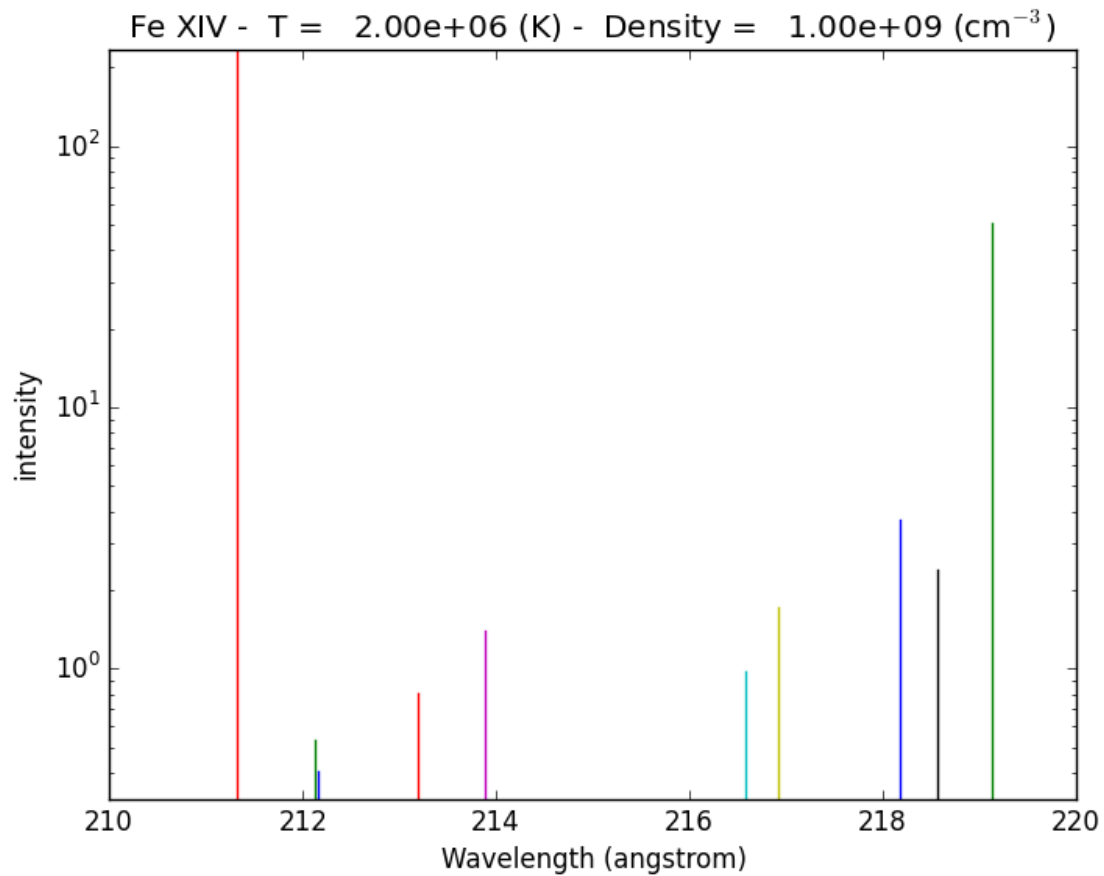
Classes and function of ChiantiPy start with lower case letters. Data/attributes that are attached to the instantiation of a class will start with a capital letter. For example,

```
fe14.populate() creates fe14.Population containing the level population information
fe14.intensity() created fe14.Intensity contain the line intensities information
fe14.spectrum() creates fe14.Spectrum contain the line and continuum spectrum,
↪ information
```

## 2.3 Spectral Line Intensities

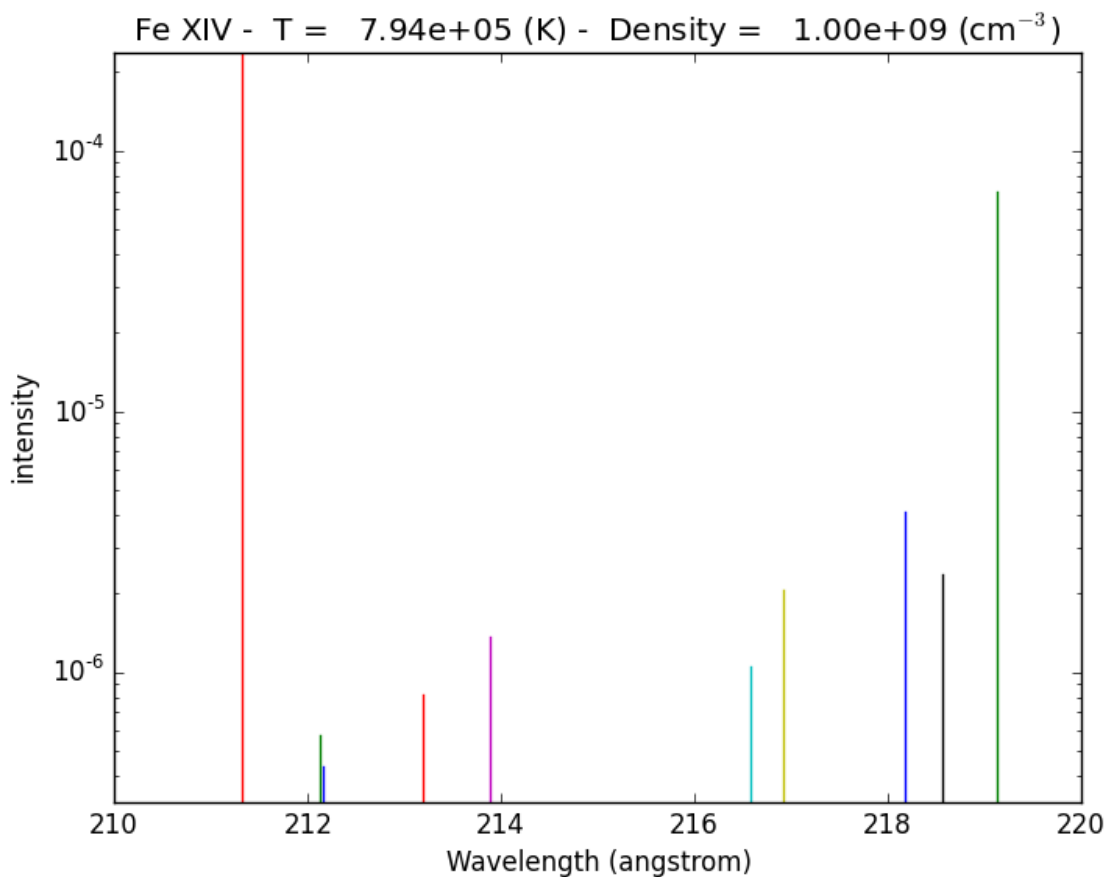
```
fe14.intensityPlot(wvlRange=[210., 220.], linLog='log')
```

will plot the intensities for the top (default = 10) lines in the specified wavelength range. If the **Intensity** attribute has not yet been calculated, it will calculate it. Since there are 21 temperature involved, a single temperature is selected ( $21/2 = 10$ ). Otherwise,



```
fe14.intensityPlot(index=2, wvlRange=[210., 220.], linLog = 'log')
```

plots the intensities for a temperature =  $t[2] = 7.9e+5$ , in this case. And, by specifying `relative = 1`, the emissivities will be plotted relative to the strongest line.



```
fe14.intensityList(wvlRange=[200,220], index=10)
```

gives the following terminal output:

```
using index = 10 specifying temperature = 2.00e+06, eDensity = 1.00e+09 em =
↳ 1.00e+27

-----

Ion  lvl1  lvl2                lower - upper                Wvl (A)
↳ Intensity  A value Obs
fe_14    1    11                3s2.3p 2P0.5 - 3s2.3d 2D1.5                211.3172
↳ 2.336e+02    3.81e+10 Y
fe_14    4    27                3s.3p2 4P1.5 - 3s.3p(3P).3d 4P1.5                212.1255
↳ 5.355e-01    2.21e+10 Y
fe_14    4    28                3s.3p2 4P1.5 - 3s.3p(3P).3d 4D2.5                212.1682
↳ 4.039e-01    1.15e+10 Y
fe_14    3    24                3s.3p2 4P0.5 - 3s.3p(3P).3d 4D0.5                213.1955
↳ 8.073e-01    4.26e+10 Y
fe_14    3    23                3s.3p2 4P0.5 - 3s.3p(3P).3d 4D1.5                213.8822
↳ 1.393e+00    2.97e+10 Y
fe_14    5    28                3s.3p2 4P2.5 - 3s.3p(3P).3d 4D2.5                216.5786
↳ 9.736e-01    2.83e+10 Y
fe_14    5    25                3s.3p2 4P2.5 - 3s.3p(3P).3d 4D3.5                216.9173
↳ 1.730e+00    4.29e+10 Y
```

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```

fe_14      7      32      3s.3p2 2D2.5 - 3s.3p(3P).3d 2F3.5      218.1767
↪ 3.734e+00      1.70e+10 Y
fe_14      4      22      3s.3p2 4P1.5 - 3s.3p(3P).3d 4P2.5      218.5725
↪ 2.391e+00      2.65e+10 Y
fe_14      2      12      3s2.3p 2P1.5 - 3s2.3d 2D2.5      219.1305
↪ 5.077e+01      4.27e+10 Y
-----

```

optionally, an output file could also be created by setting the keyword outFile to the name of the desired name

```
fe14.intensityList(wvlRange=[210.,220.], relative=1, index=11)
```

give the following terminal/notebook output

```

using index = 11 specifying temperature = 2.24e+06, eDensity = 1.00e+09 em =
↪ 1.00e+27
-----

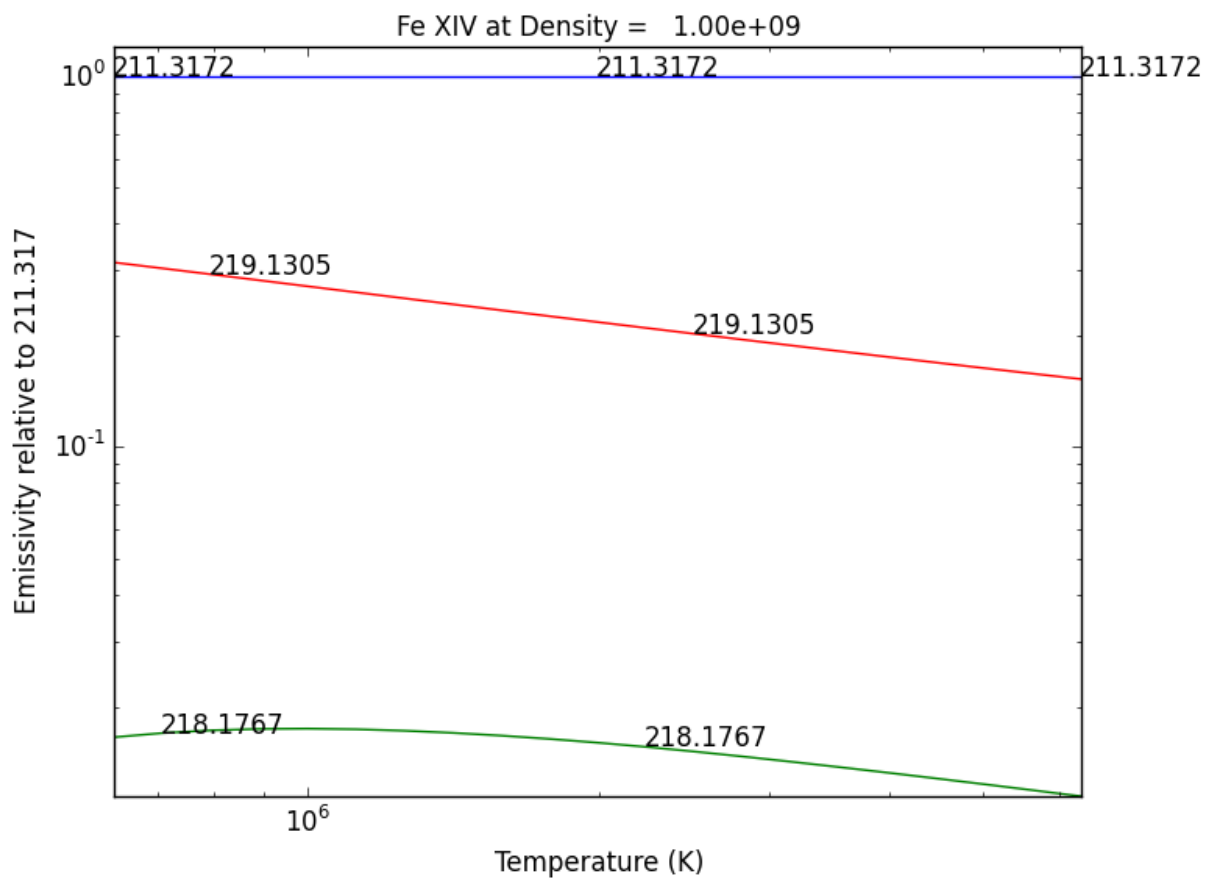
Ion  lvl1  lvl2      lower - upper      Wvl (A)
↪ Intensity      A value Obs
fe_14      1      11      3s2.3p 2P0.5 - 3s2.3d 2D1.5      211.3172
↪ 1.000e+00      3.81e+10 Y
fe_14      4      27      3s.3p2 4P1.5 - 3s.3p(3P).3d 4P1.5      212.1255
↪ 2.267e-03      2.21e+10 Y
fe_14      4      28      3s.3p2 4P1.5 - 3s.3p(3P).3d 4D2.5      212.1682
↪ 1.694e-03      1.15e+10 Y
fe_14      3      24      3s.3p2 4P0.5 - 3s.3p(3P).3d 4D0.5      213.1955
↪ 3.390e-03      4.26e+10 Y
fe_14      3      23      3s.3p2 4P0.5 - 3s.3p(3P).3d 4D1.5      213.8822
↪ 5.891e-03      2.97e+10 Y
fe_14      5      28      3s.3p2 4P2.5 - 3s.3p(3P).3d 4D2.5      216.5786
↪ 4.083e-03      2.83e+10 Y
fe_14      5      25      3s.3p2 4P2.5 - 3s.3p(3P).3d 4D3.5      216.9173
↪ 7.085e-03      4.29e+10 Y
fe_14      7      32      3s.3p2 2D2.5 - 3s.3p(3P).3d 2F3.5      218.1767
↪ 1.557e-02      1.70e+10 Y
fe_14      4      22      3s.3p2 4P1.5 - 3s.3p(3P).3d 4P2.5      218.5725
↪ 1.009e-02      2.65e+10 Y
fe_14      2      12      3s2.3p 2P1.5 - 3s2.3d 2D2.5      219.1305
↪ 2.096e-01      4.27e+10 Y
-----

```

## 2.4 G(n,T) function

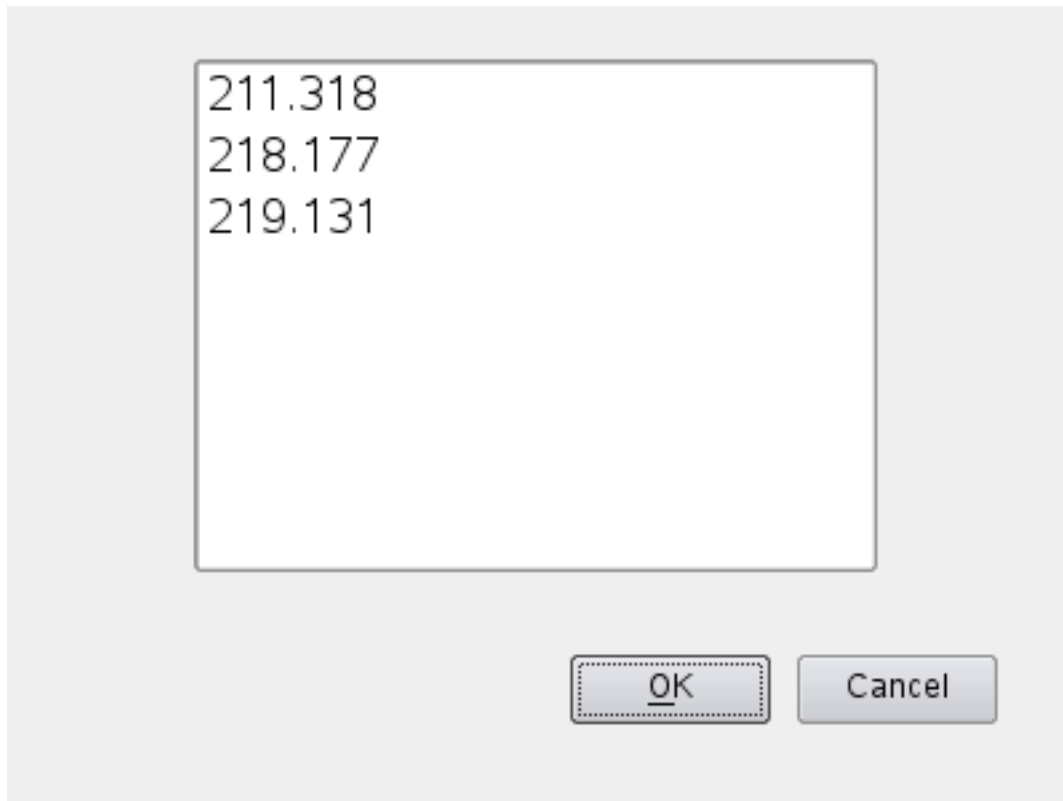
```
fe14.gofnt(wvlRange=[210., 220.], top=3)
```

brings up a matplotlib plot window which shows the emissivities of the top (strongest) 3 lines in the wavelength region from 210 to 220 Angstroms.

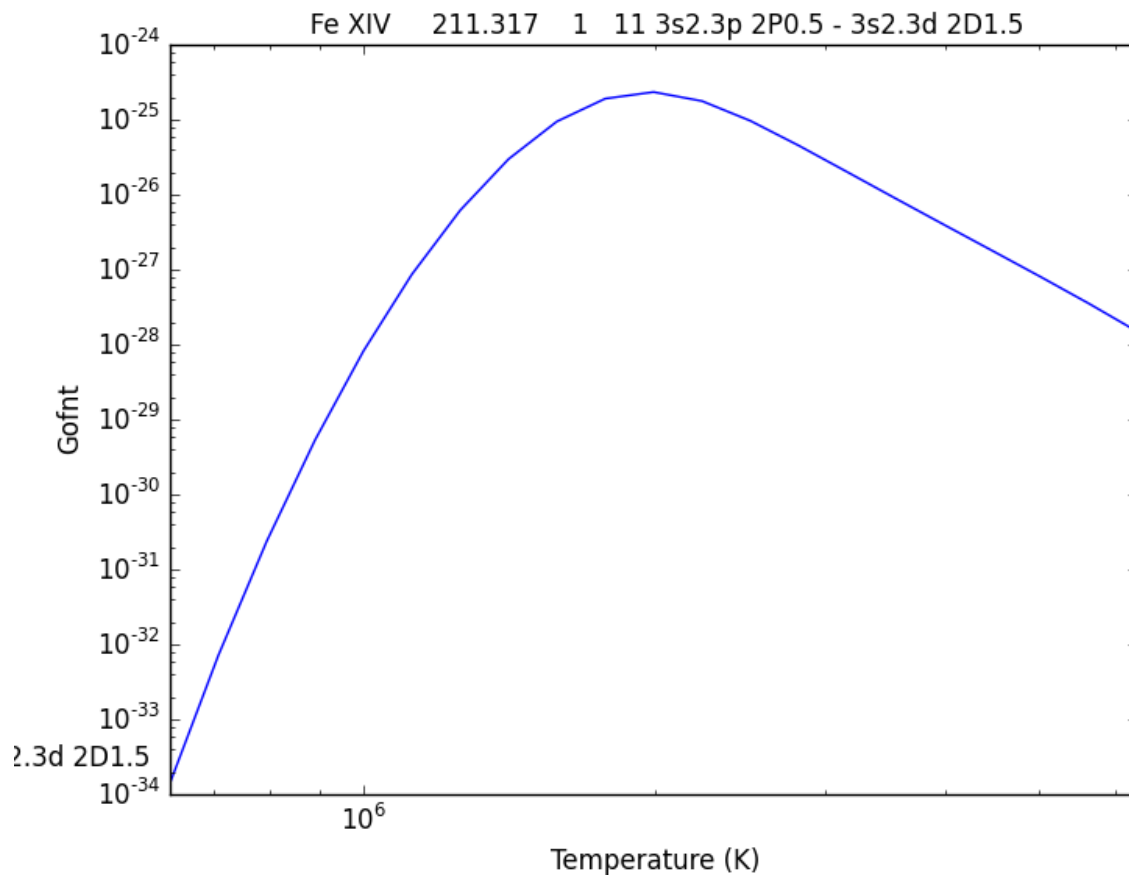


quickly followed by a dialog where the line(s) of interest can be specified





and finally a plot of the  $G(n,T)$  function for the specified lines(s).



The  $G(n,T)$  calculation is stored in the Gofnt dictionary, with keys = ['gofnt', 'temperature', 'density']

while there is a fairly straightforward way to get a  $G(T)$  function, it is not very practical to use for more than a handful of lines. For if the fe\_14 line at 211.3172 is in a list of lines to be analyzed, a more practical way is the following

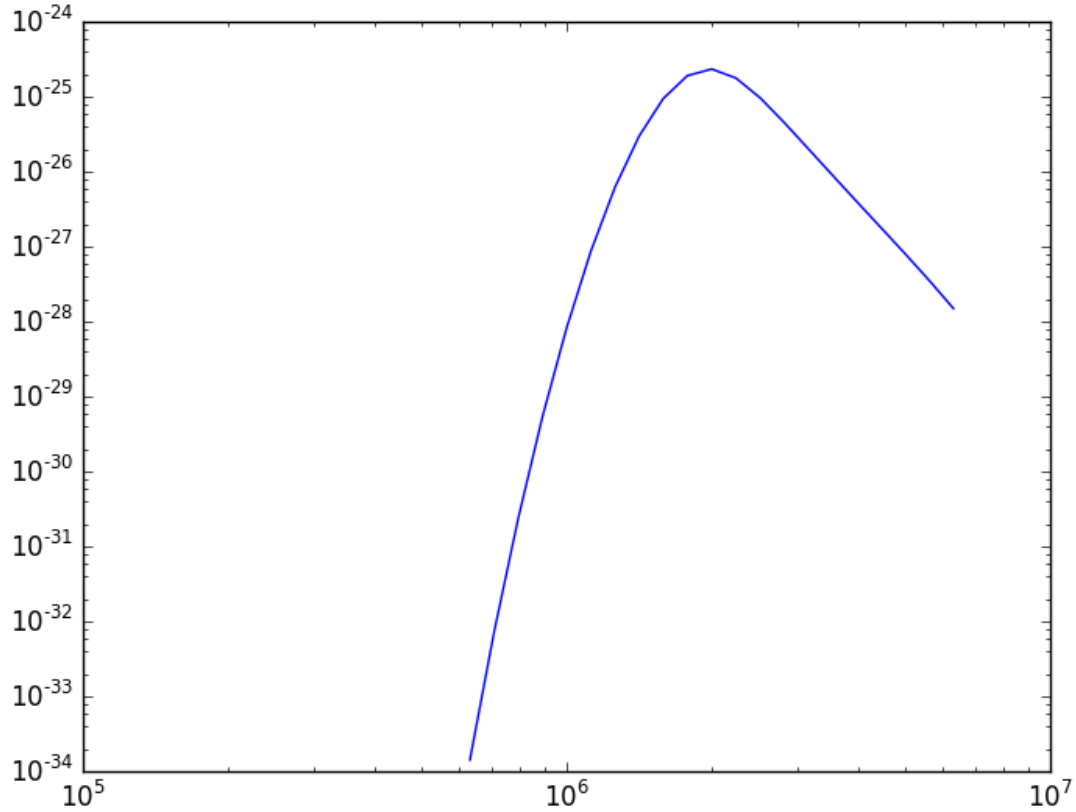
```
fe14.intensity()
dist = np.abs(np.asarray(fe14.Intensity['wvl']) - 211.3172)
idx = np.argmin(dist)
print(' wvl = %10.3f'%(fe14.Intensity['wvl'][idx]))
```

prints

wvl = 211.317

```
plt.loglog(temp, fe14.Intensity['intensity'][:,idx])
```

once the axes are properly scaled, this produces the same values as `fe14.Gofnt['gofnt']`

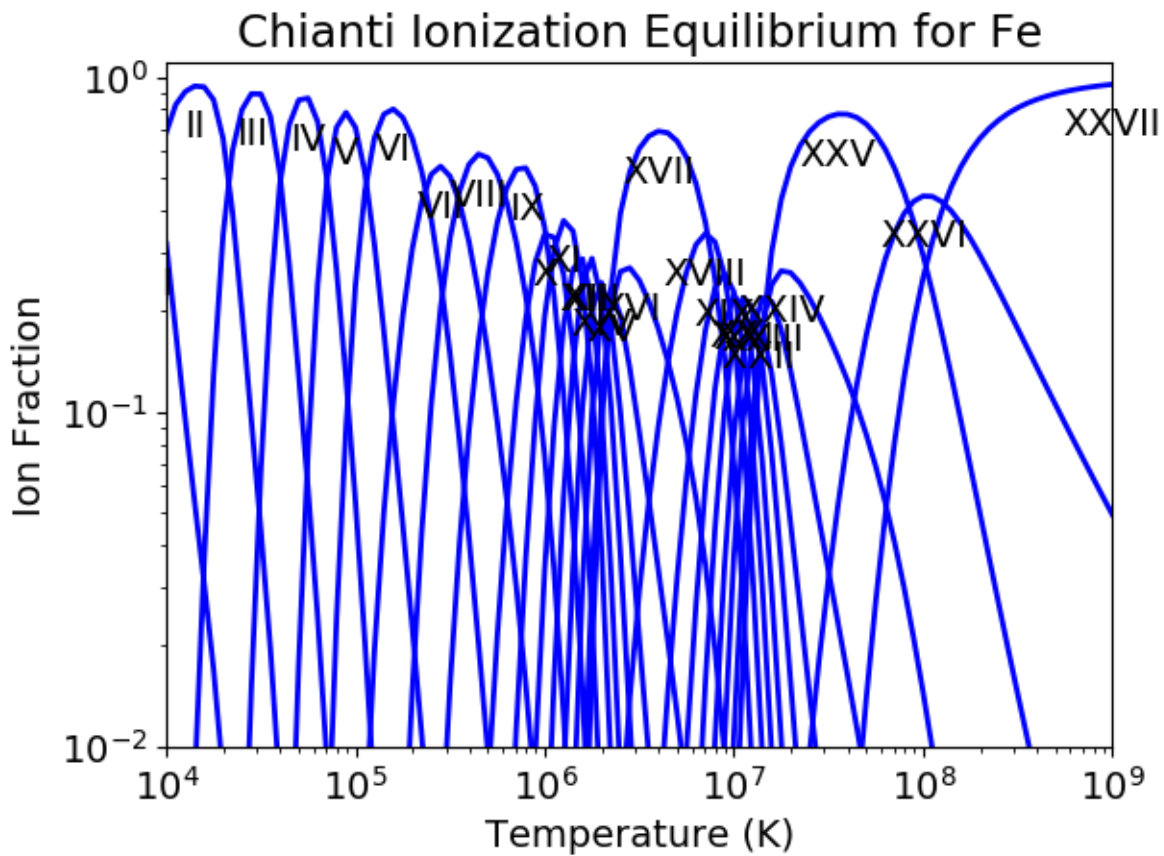


## 2.5 Ionization Equilibrium

For the Fe XIV example, the temperature was chosen to center around  $2 \times 10^6$ . It was not immediately apparent why this was done but in most of the following examples it is necessary to pick an appropriate temperature. This can be done with the **ioneq** class. To look at the ionization equilibrium for the iron ions ( $Z = 26$ , or 'fe')

```
fe = ch.ioneq(26)
fe.load()
fe.plot()
plt.tight_layout()
```

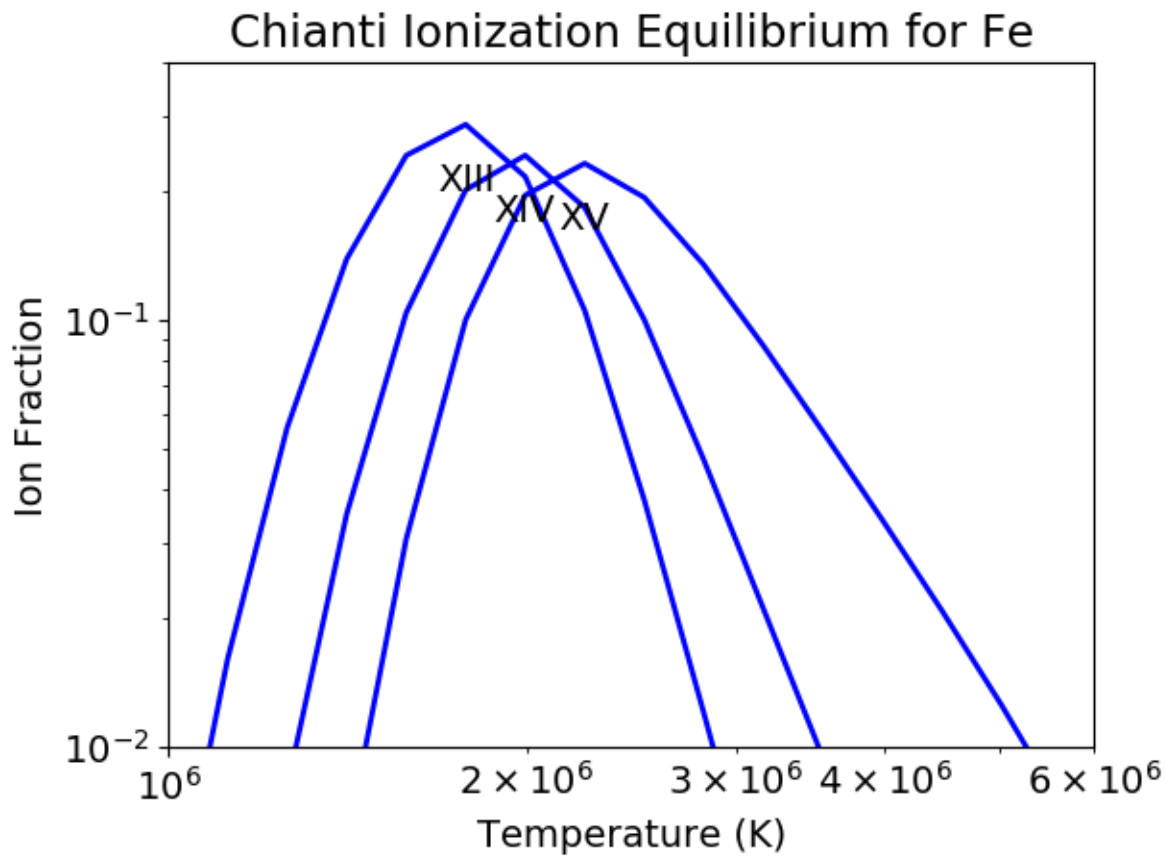
brings up a plot showing the ionization equilibrium for all of the stages of iron as a function of temperature



This is pretty crowded and we are only interested in Fe XIV (fe\_14), so

```
plt.figure()
fe.plot(stages=[13,14,15],tRange=[1.e+6, 6.e+6], yr = [1.e-2, 0.4])
plt.tight_layout()
```

produces a plot of the ionization equilibria of Fe XIII, XIV and XV over a limited temperature range (tRange) and vertical range (yr)

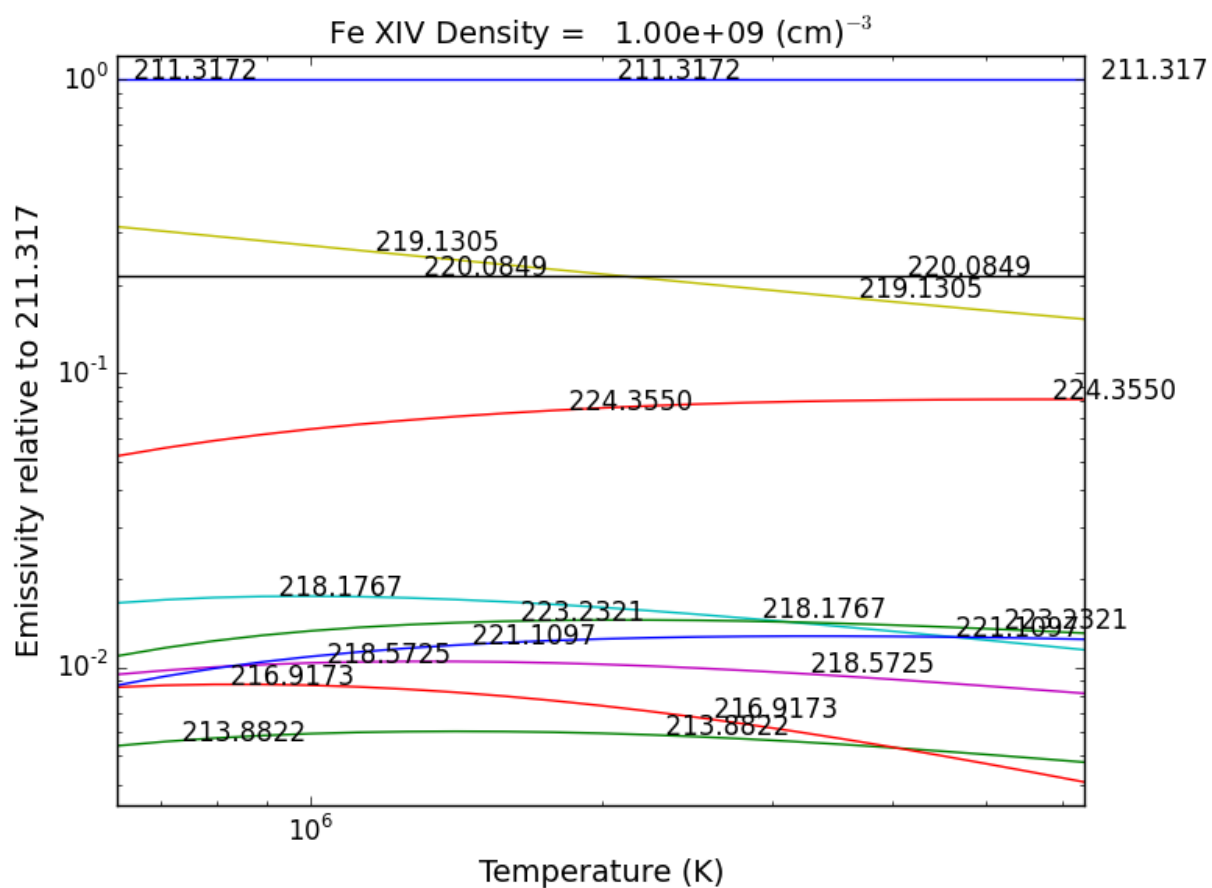


from this it is clear that Fe XIV (fe\_14) is formed at temperatures near  $2 \times 10^6$  K

## 2.6 Intensity Ratios

```
fe14.intensityRatio(wvlRange=[210., 225.])
```

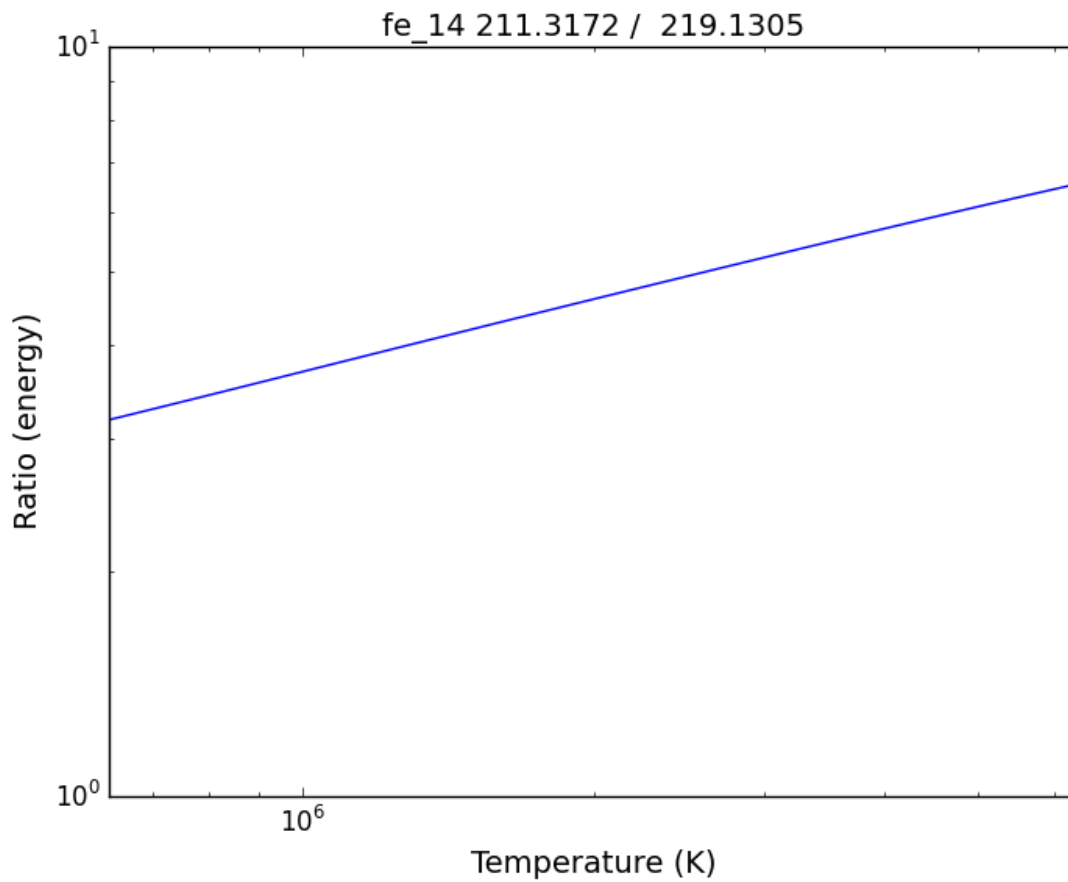
this brings up a plot showing the relative emissivities on the Fe XIV lines



following by a dialog where you can selector the numerator(s) and denominator(s) of the desired intensity ratio

| Numerator | Denominator |
|-----------|-------------|
| 211.318   | 211.318     |
| 213.883   | 213.883     |
| 216.918   | 216.918     |
| 218.177   | 218.177     |
| 218.573   | 218.573     |
| 219.131   | 219.131     |
| 220.085   | 220.085     |
| 221.11    | 221.11      |
| 223.233   | 223.233     |
| 224.355   | 224.355     |

so the specified ratio is then plotted

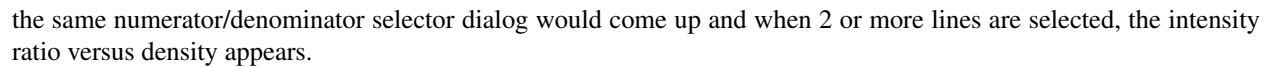


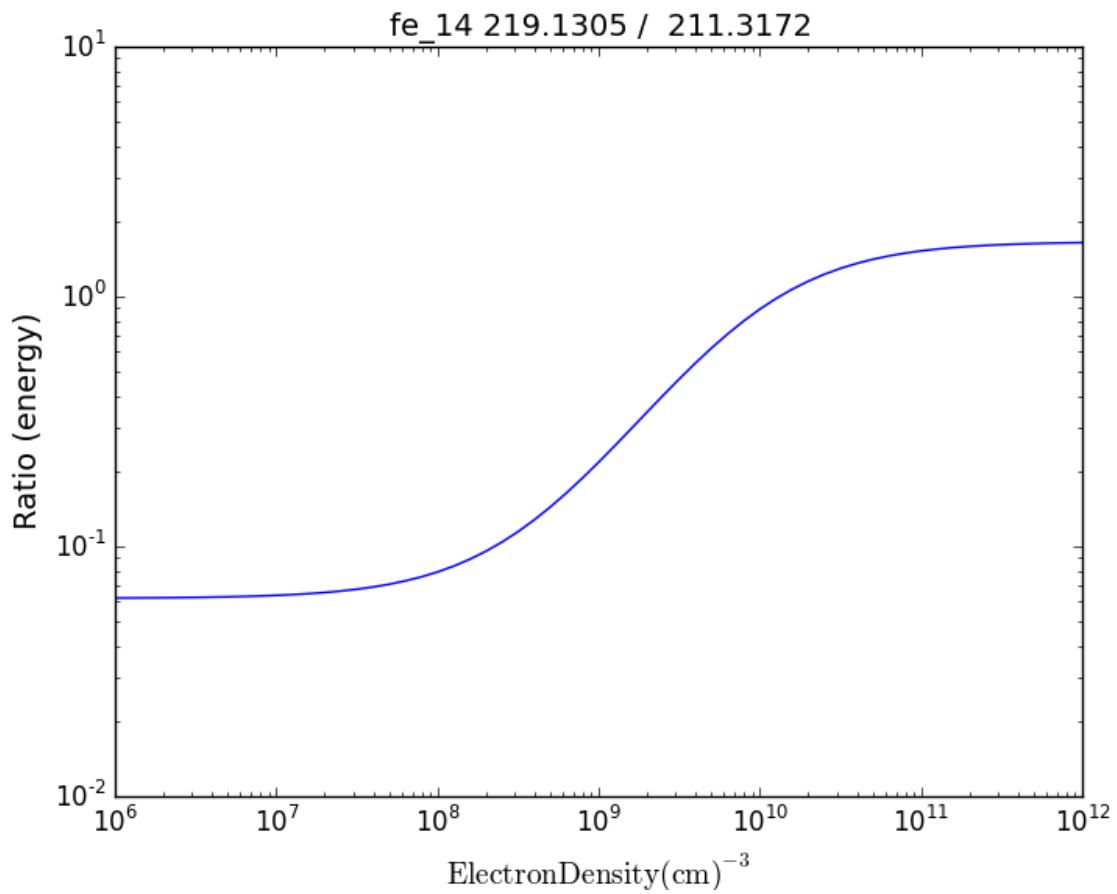
if previously, we had done

```
dens = 10.**(6. + 0.1*arange(61))
fe14 = ch.ion('fe_14', 2.e+6, dens)
fe14.intensityRatio(wvlRange=[210., 225.]
```

then the plot of relative intensities vs density would appear

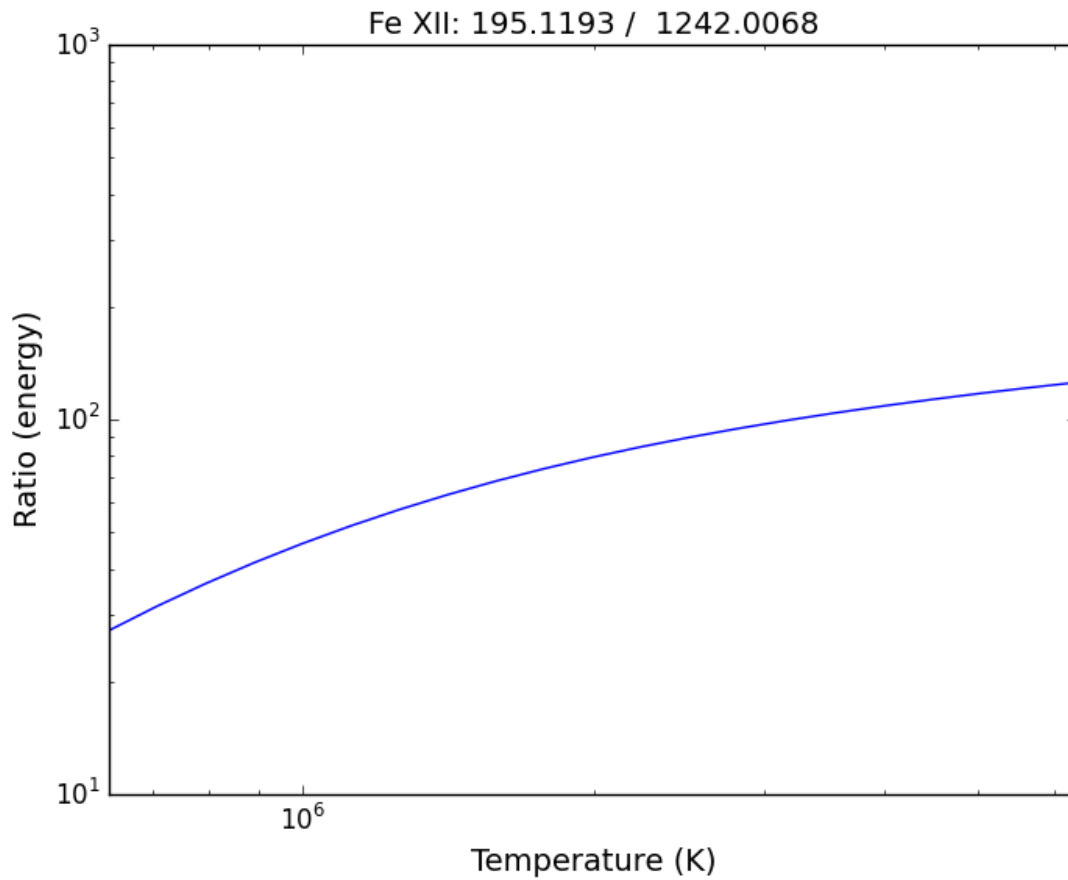






to obtain ratios of lines widely separated in wavelength, the `wvlRanges` keyword can be used:

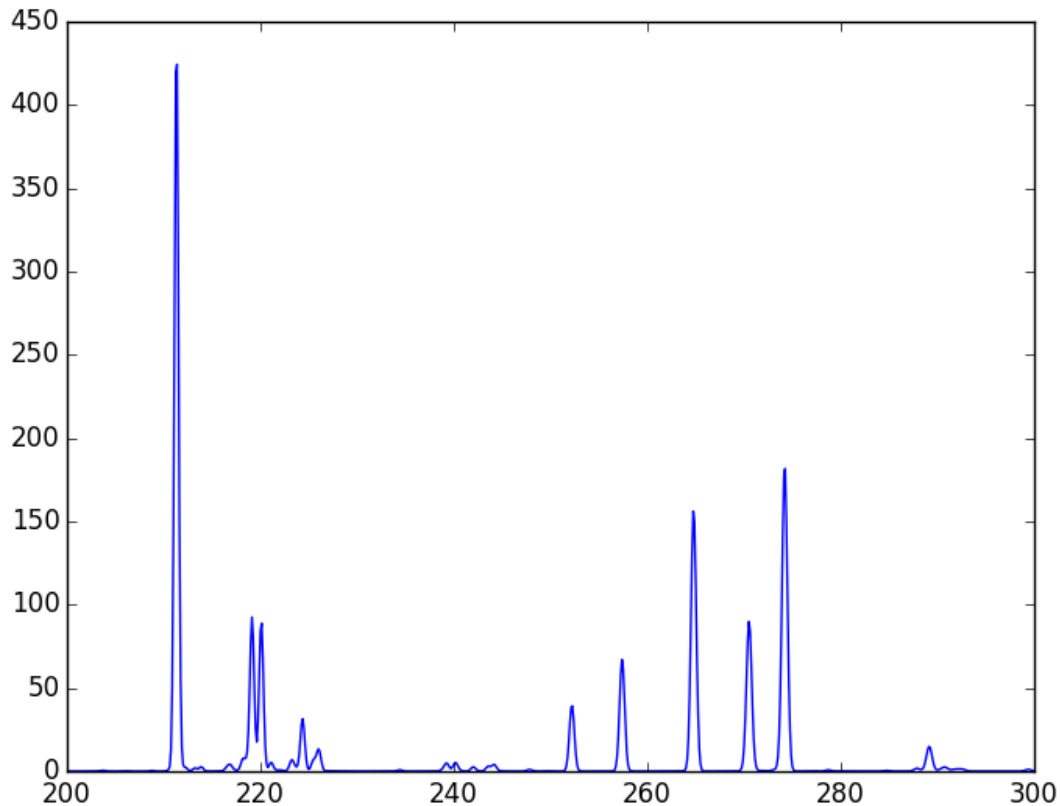
```
fe12 = ch.ion('fe_12', temperature=t, eDensity=1.e+9
fe12.intensityRatio(wvlRanges=[[190.,200.],[1240.,1250.]])
```



## 2.7 Spectra of a single ion

```
fe14 = ch.ion('fe_14', temperature = 2.e+6, density = 1.e+9)
wvl = wvl=200. + 0.125*arange(801)
fe14.spectrum(wvl, em=1.e+27)
plot(wvl, fe14.Spectrum['intensity'])
```

this will calculate the spectrum of fe\_14 over the specified wavelength range and filter it with the default filter which is a gaussian (filters.gaussianR) with a 'resolving power' of 1000 which gives a gaussian width of wvl/1000.



other filters available in `chianti.tools.filters` include a boxcar filter and a gaussian filter where the width can be specified directly

```
if hasattr(fe14, 'Em') :
    print(' Emission Measure = %12.2e'%(fe14.Em))
else:
    print(' the value for the emission measure is unspecified')
```

Emission Measure = 1.00e+27

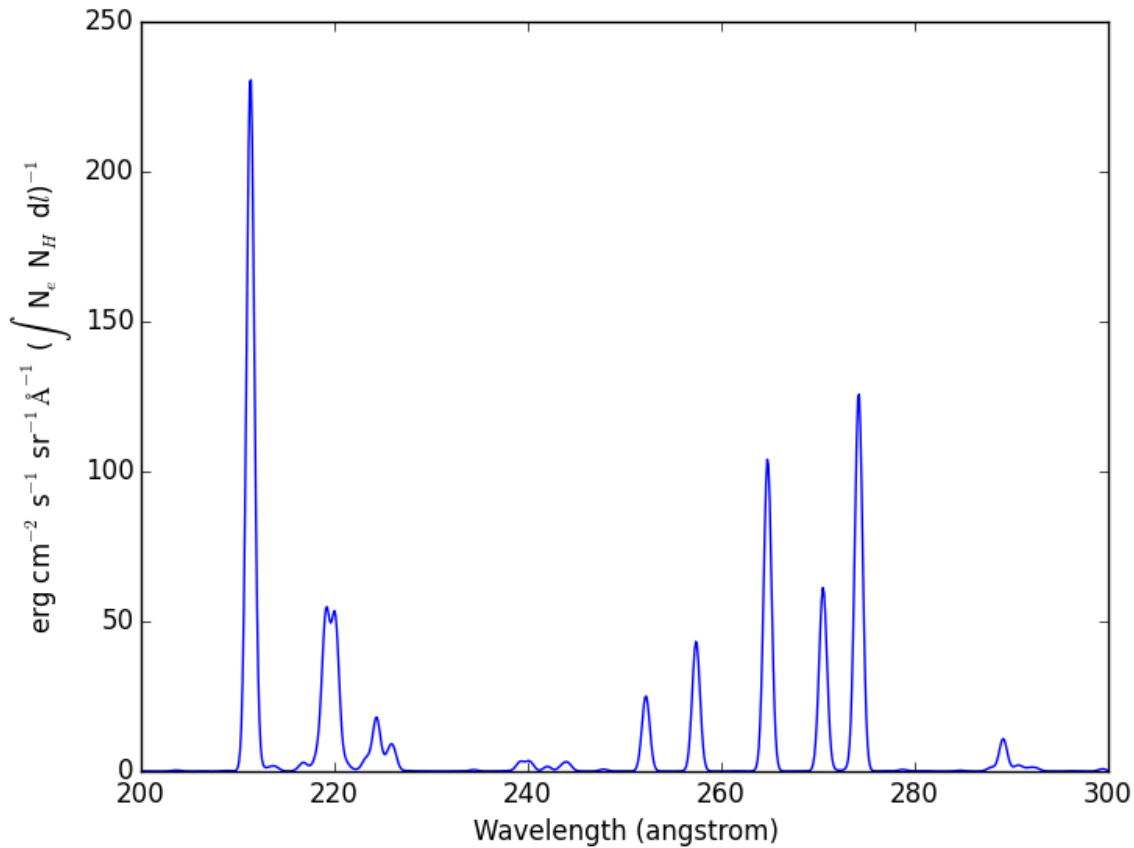
```
import chianti.tools.filters as chfilters
fe14.spectrum(wvl, filter=(chfilters.gaussian, .04))
```

calculates the spectrum of fe\_14 for a gaussian filter with a width of 0.04 Angstroms. The current value of the spectrum is kept in `fe14.Spectrum` with the following keys:

```
for akey in sorted(fe14.Spectrum.keys()):
    print(' %10s'%(akey))
```

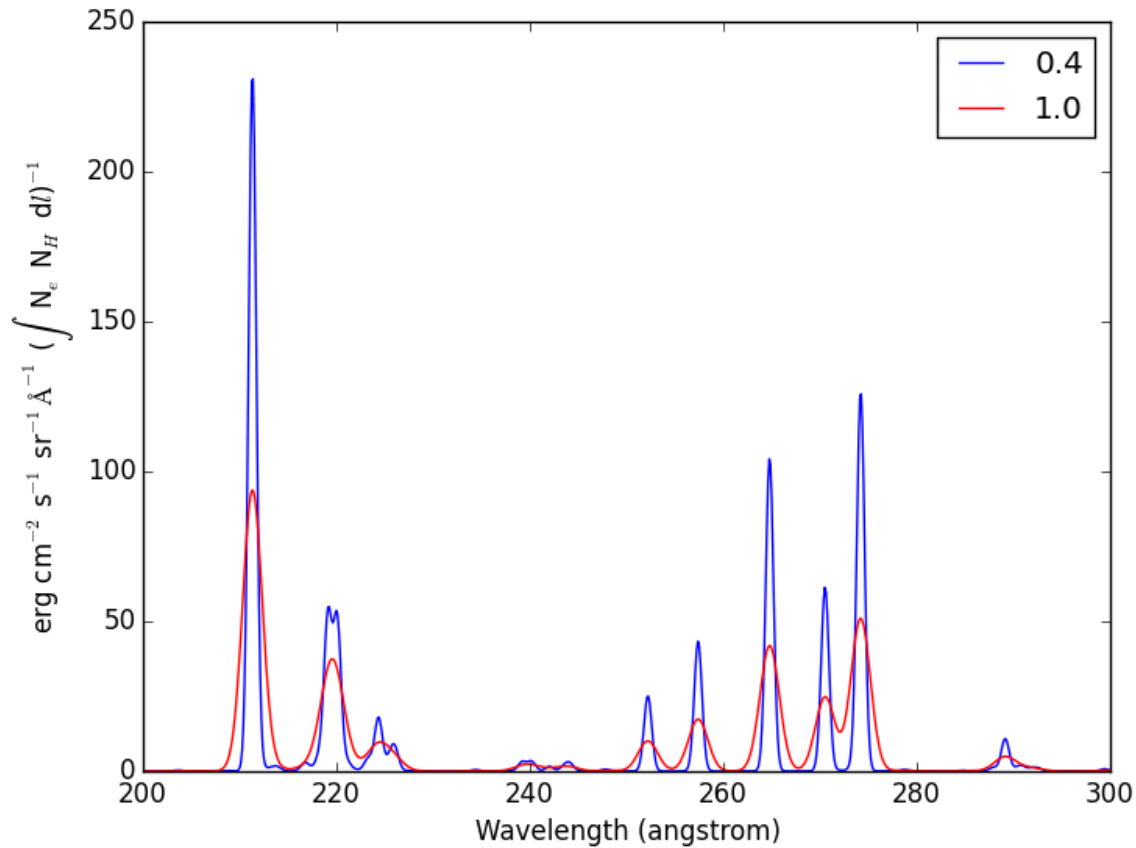
allLines em filter filterWidth intensity wvl xlabel ylabel

```
plot(wvl, fe14.Spectrum['intensity'])
plt.xlabel(fe14.Spectrum['xlabel'])
plt.ylabel(fe14.Spectrum['ylabel'])
```



New in **ChiantiPy 0.6**, the *label* keyword has been added to the `ion.spectrum` method, and also to the other various spectral classes. This allows several spectral calculations for different filters to be saved and compared

```
temp = 10.**(5.8 + 0.1*np.arange(11.))
dens = 1.e+9
fe14 = ch.ion('fe_14', temp, dens)
emeas = np.ones(11, 'float64')*1.e+27
wvl = 200. + 0.125*np.arange(801)
fe14.spectrum(wvl, filter=(chfilters.gaussian,.4), label='.4', em=emeas, label='0.4')
fe14.spectrum(wvl, filter=(chfilters.gaussian,1.), label='1.', label='1.0')
plt.plot(wvl, fe14.Spectrum['.4']['intensity'][5])
plt.plot(wvl, fe14.Spectrum['1.']['intensity'][5], '-r')
plt.xlabel(fe14.Spectrum['.4']['xlabel'])
plt.ylabel(fe14.Spectrum['.4']['ylabel'])
plt.legend(loc='upper right')
```

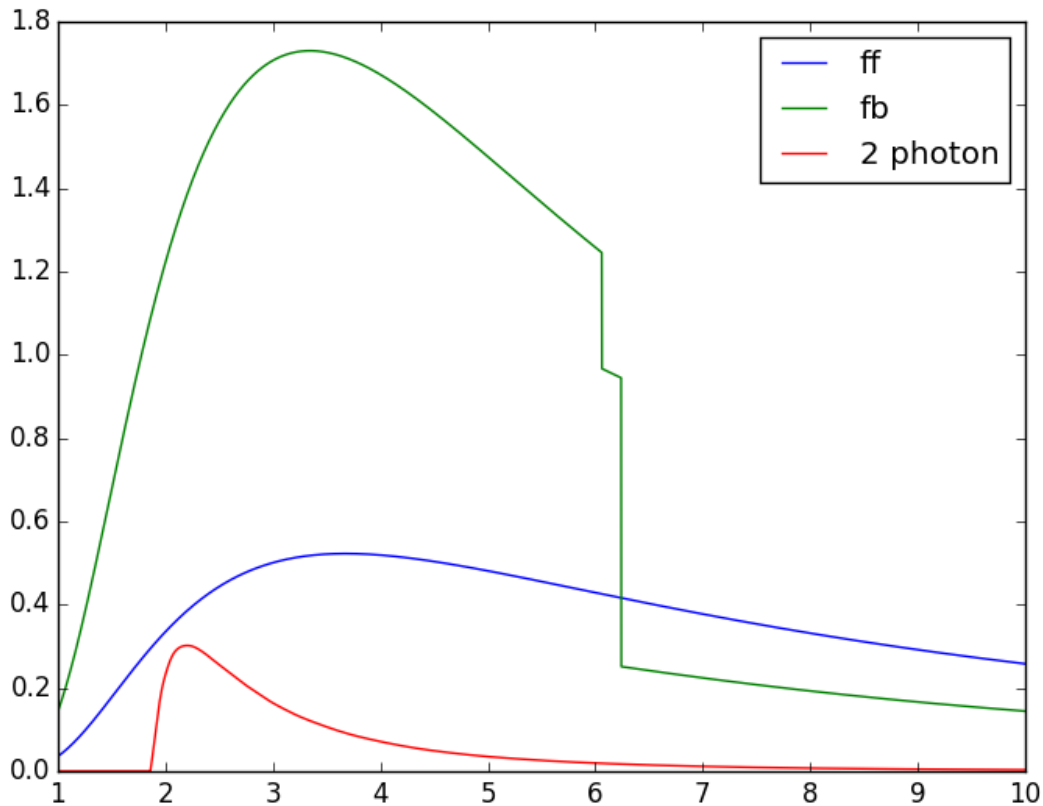


## 2.8 Free-free and free-bound continuum

The module continuum provides the ability to calculate the free-free and free-bound spectrum for a large number of individual ions. The two-photon continuum is produced only by the hydrogen-like and helium-like ions

```
temperature = 2.e+7
em = 1.e+27
c = ch.continuum('fe_25', temperature = temperature, em = em)
wvl = 1. + 0.002*arange(4501)
c.freeFree(wvl)
plot(wvl, c.FreeFree['intensity'])
c.freeBound(wvl)
plot(wvl, c.FreeBound['intensity'])
fe25=ch.ion('fe_25',2.e+7,1.e+9,em=1.e+27)
fe25.twoPhoton(wvl)
plt.plot(wvl,fe25.TwoPhoton['intensity'],label='2 photon')
plt.legend(loc='upper right')
```

produces



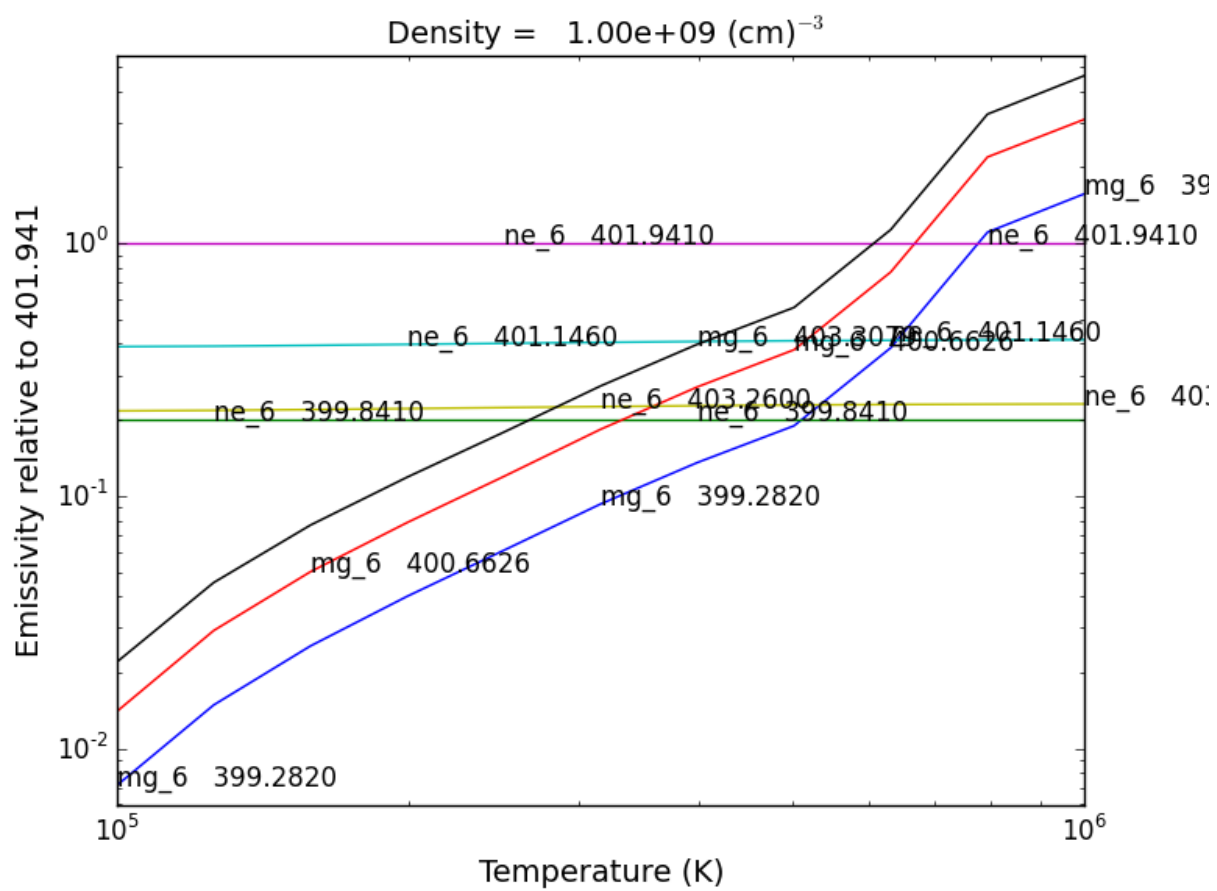
In the continuum calculations, the specified ion, Fe XXV in this case, is the target ion for the free-free calculation. For the free-bound calculation, specified ion is also the target ion. In this case, the radiative recombination spectrum of Fe XXV recombining to form Fe XXIV is returned.

## 2.9 The multi-ion class Bunch

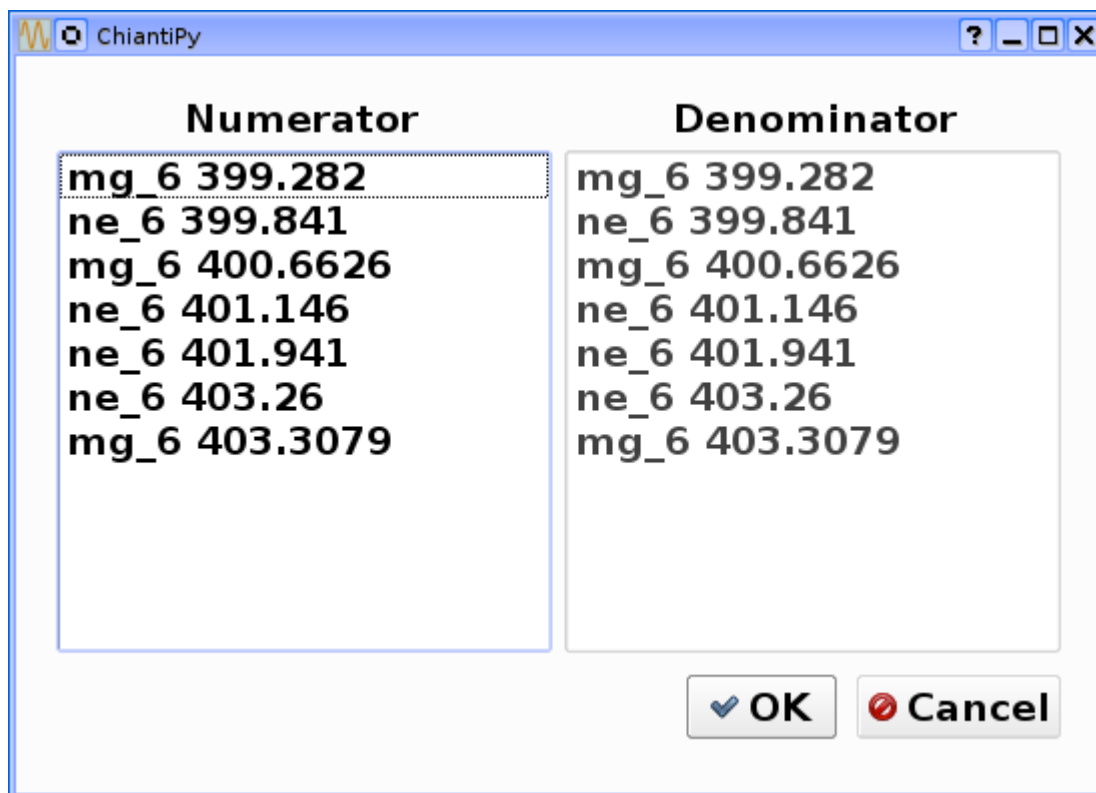
The multi-ion class **bunch** [new in v0.6] inherits a number of the same methods inherited by the ion class, for example *intensityList*, *intensityRatio*, and *intensityRatioSave*. As a short demonstration of its usefulness, Widing and Feldman (1989, ApJ, 344, 1046) used line ratios of Mg VI and Ne VI as diagnostics of elemental abundance variations in the solar atmosphere. For that to be accurate, it is necessary that the lines of the two ions have the same temperature response.

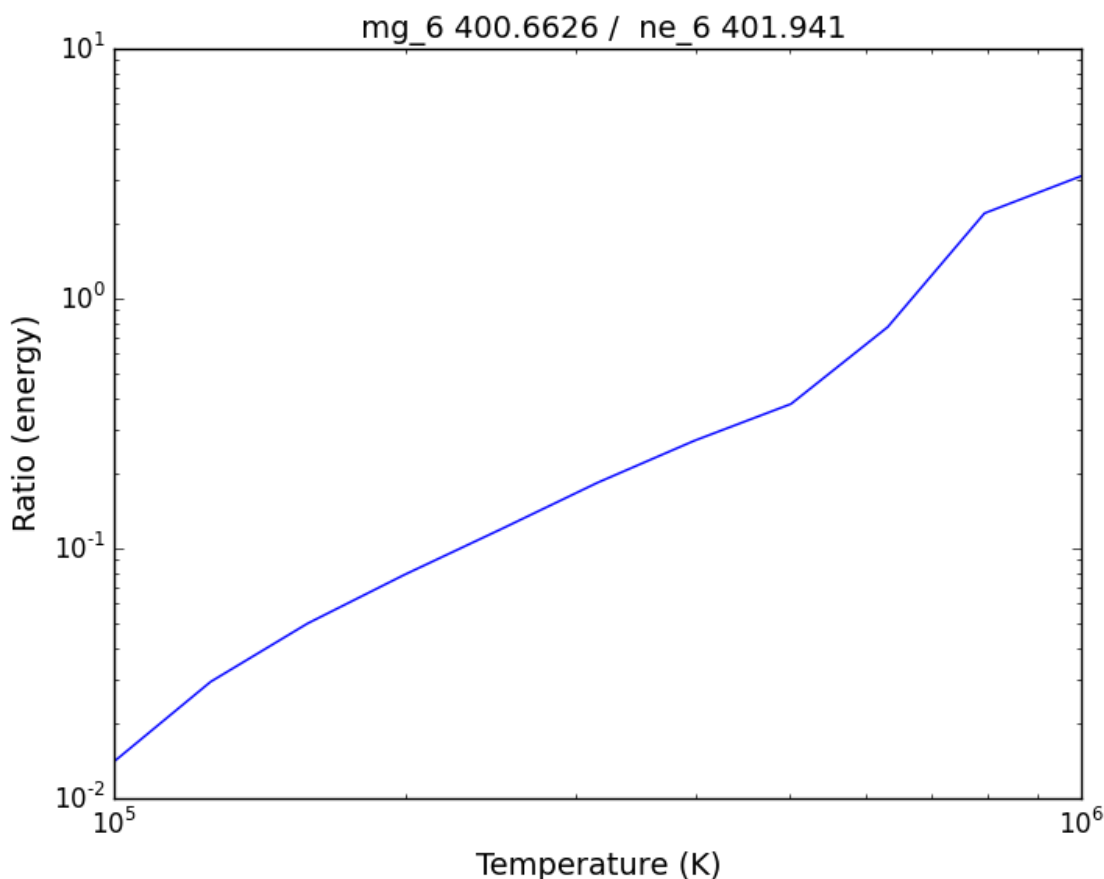
```
t = 10.** (5.0+0.1*np.arange(11))
bnch=ch.bunch(t,1.e+9,wvlRange=[300.,500.],ionList=['ne_6','mg_6'],abundance='unity')
bnch.intensityRatio(wvlRange=[395.,405.],top=7)
```

produces and initial plot of the selected lines, a selection widget and finally a plot of the ratio









there seems to be a significant temperature dependence to the ratio, even though both are formed near  $4.e+5$  K.

A new keyword argument **keepIons** has been added in v0.6 to the bunch and the 3 spectrum classes.

```
dwvl = 0.01
nwvl = (406.-394.)/dwvl
wvl = 394. + dwvl*np.arange(nwvl+1)
bnch2=ch.bunch(t, 1.e+9, wvlRange=[wvl.min(),wvl.max()], elementList=['ne','mg'],
    keepIons=1,em=1.e+27)
bnch2.convolve(wvl,filter=(chfilters.gaussian,5.*dwvl))
plt.plot(wvl, bnch2.Spectrum['intensity'][6],label='Total')
plt.title('Temperature = %10.2e for t[6]%(t[6]))
```

elapsed seconds = 11.000 elapsed seconds = 0.000e+00

```
for one in sorted(bnch2.IonInstances.keys()):
    print('%s'%(one))
```

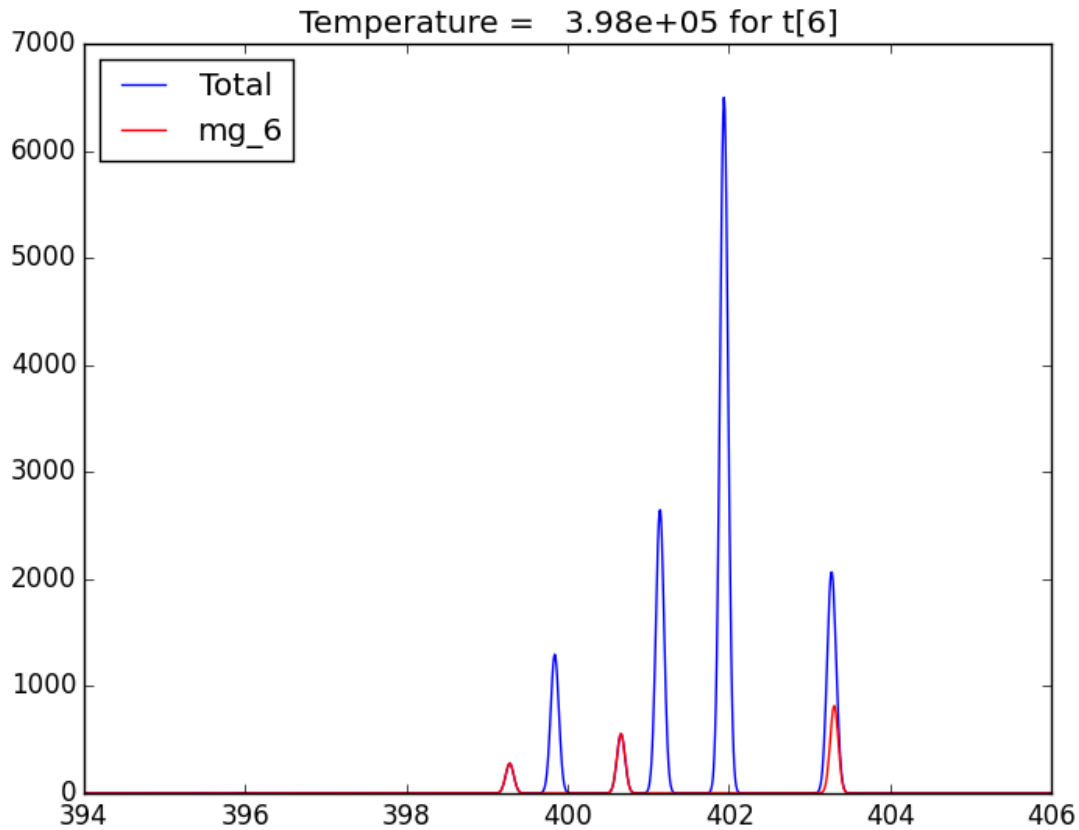
yields:

mg\_10 mg\_10d mg\_3 mg\_4 mg\_5 mg\_6 mg\_8 mg\_9 ne\_10 ne\_2 ne\_3 ne\_5 ne\_6 ne\_8

these IonInstances have all the properties of the Ion class for each of these ions

```
plt.plot(wvl,bnch2.IonInstances['mg_6'].Spectrum['intensity'][6], 'r',label='mg_6')
plt.legend(loc='upper left')
```

produces



## 2.10 Spectra of multiple ions and continuum

the spectrum for all ions in the CHIANTI database can also be calculated

The spectrum for a selection of all of the ions in the CHIANTI database can also be calculated. There are 3 spectral classes.

- **spectrum** - the single processor implementation that can be used anywhere
- **mspectrum** - uses the Python multiprocessing class and cannot be used in a IPython qtconsole or notebook
- **ipyspectrum** [new in v0.6] - uses the IPython parallel class and can be used in a IPython qtconsole or notebook

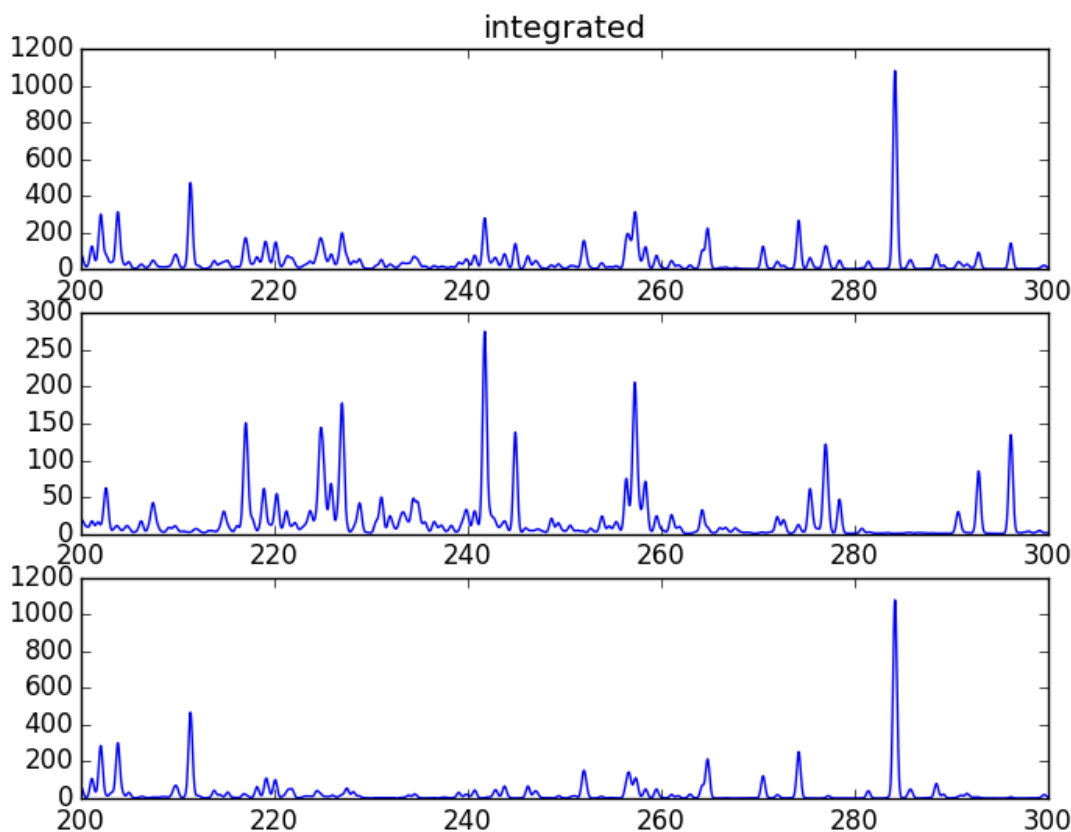
### 2.10.1 The single processor spectrum class

```

temperature = [1.e+6, 2.e+6]
density = 1.e+9
wvl = 200. + 0.05*arange(2001)
emeasure = [1.e+27, 1.e+27]
s = ch.spectrum(temperature, density, wvl, filter = (chfilters.gaussian, .2), em =
emeasure, doContinuum=0, minAbund=1.e-5)
subplot(311)
plot(wvl, s.Spectrum['integrated'])
subplot(312)
plot(wvl, s.Spectrum['intensity'][0])
subplot(313)
plot(wvl, s.Spectrum['intensity'][1])

```

produces



The integrated spectrum is formed by summing the spectra for all temperatures.

- For  $\text{minAbund}=1.e-6$ , the calculation takes 209 s on a 3.5 GHz processor.
- For  $\text{minAbund}=1.e-5$ , the calculation takes 122 s on a 3.5 GHz processor.

The filter is not applied to the continuum.

Calculations with the Spectrum module can be time consuming. One way to control the length of time the calculations take is to limit the number of ions with the `ionList` keyword and to avoid the continuum calculations by setting

the `doContinuum` keyword to 0 or False. Another way to control the length of time the calculations take is with the `minAbund` keyword. It sets the minimum elemental abundance that an element can have for its spectra to be calculated. The default value is set include all elements. Some usefull values of `minAbund` are:

- `minAbund = 1.e-4`, will include H, He, C, O, Ne
- `minAbund = 2.e-5` adds N, Mg, Si, S, Fe
- `minAbund = 1.e-6` adds Na, Al, Ar, Ca, Ni

## 2.10.2 The multiple processor `mspectrum` class

Another way to speed up calculations is to use the `mspectrum` class which uses multiple cores on your local computer. It requires the Python *multiprocessing* module which is available with Python versions 2.6 and later. `mspectrum` is called in the same way as `spectrum` but you can specify the number of cores with the `proc` keyword. The default is 3 but it will not use more cores than are available on your machine. For example,

```
s = ch.mspectrum(temperature, density ,wvl, em=emeasure, filter = (chfilters.gaussian,
↪.005), proc=4)
```

## 2.10.3 The multiple processor `ipyspectrum` class

next, we will use the `ipyspectrum` class. First, it is necessary to start up the cluster. In some shell

```
> ipcluster start -n=4
```

or, if you are using Python3

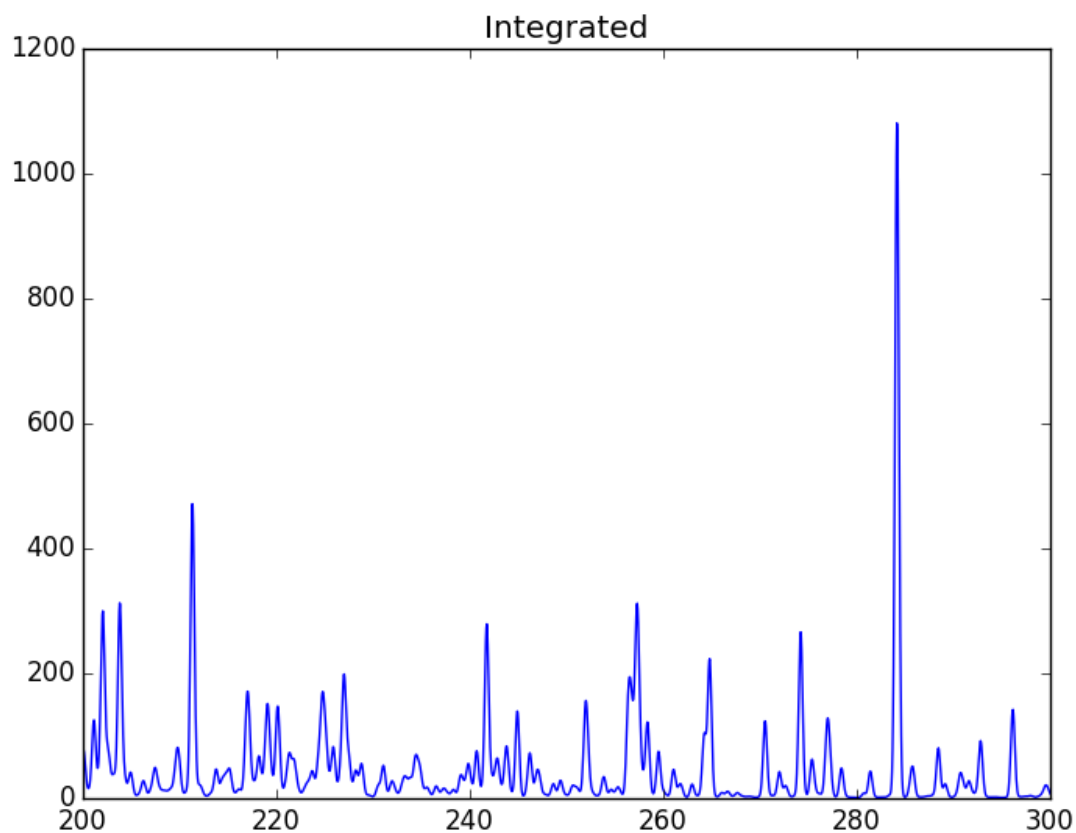
```
> ipcluster3 start -n=4
```

this will start 4 engines if you have 4 cores but it won't start more than you have

then in an IPython notebook or qtconsole

```
temp = [1.e+6, 2.e+6]
dens = 1.e+9
wvl = 200. + 0.05*np.arange(2001)
emeasure = [1.e+27 ,1.e+27]
s = ch.ipyspectrum(temp, dens, wvl, filter = (chfilters.gaussian,.2), em = emeasure,
↪doContinuum=1, minAbund=1.e-5, verbose=0)
plt.figure
plt.plot(wvl, s.Spectrum['integrated'])
```

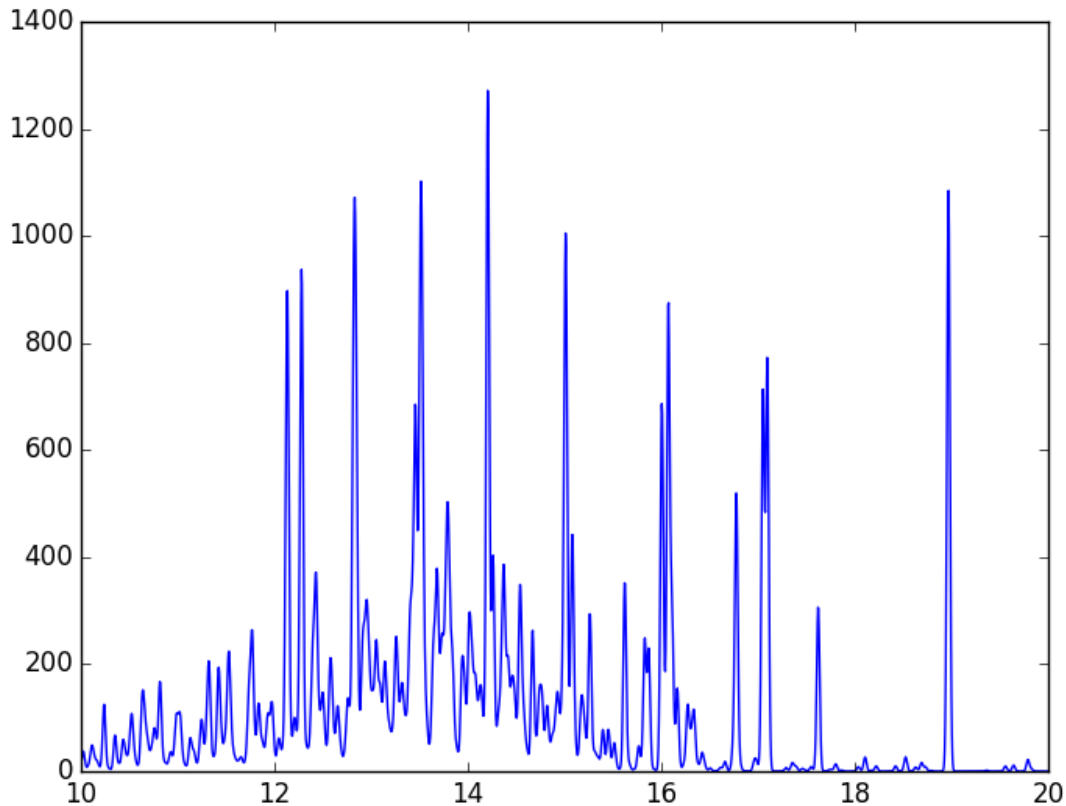
produces



spectrum, mspectrum and ipyspectrum can all be instantiated with the same arguments and keyword arguments. Most of the examples below use the ipyspectrum class for speed.

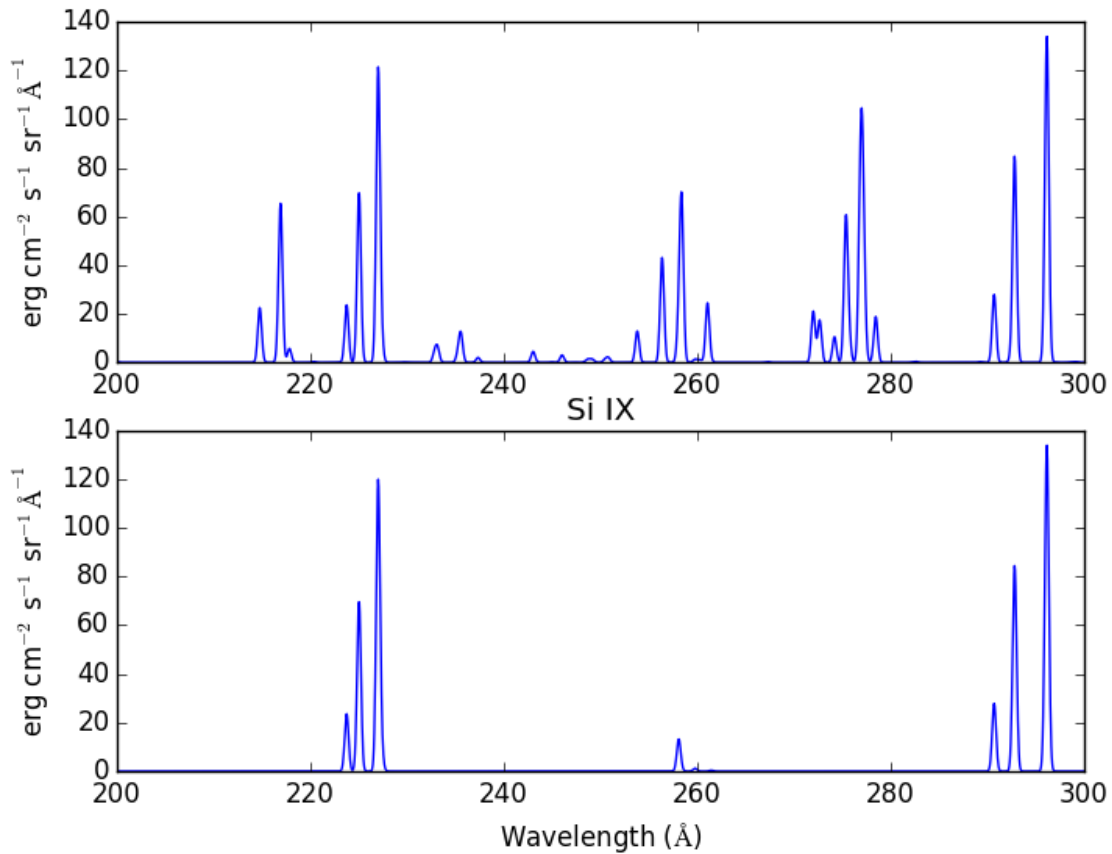
```
temperature = 1.e+7
wvl = 10. + 0.005*arange(2001)
s = ch.ipyspectrum(temperature, density, wvl, filter = (chfilters.gaussian,.015))
plot(wvl, s.Spectrum['intensity'])
```

produces



It is also possible to specify a selection of ions by means of the *ionList* keyword, for example, *ionList=['fe\_11','fe\_12','fe\_13']*

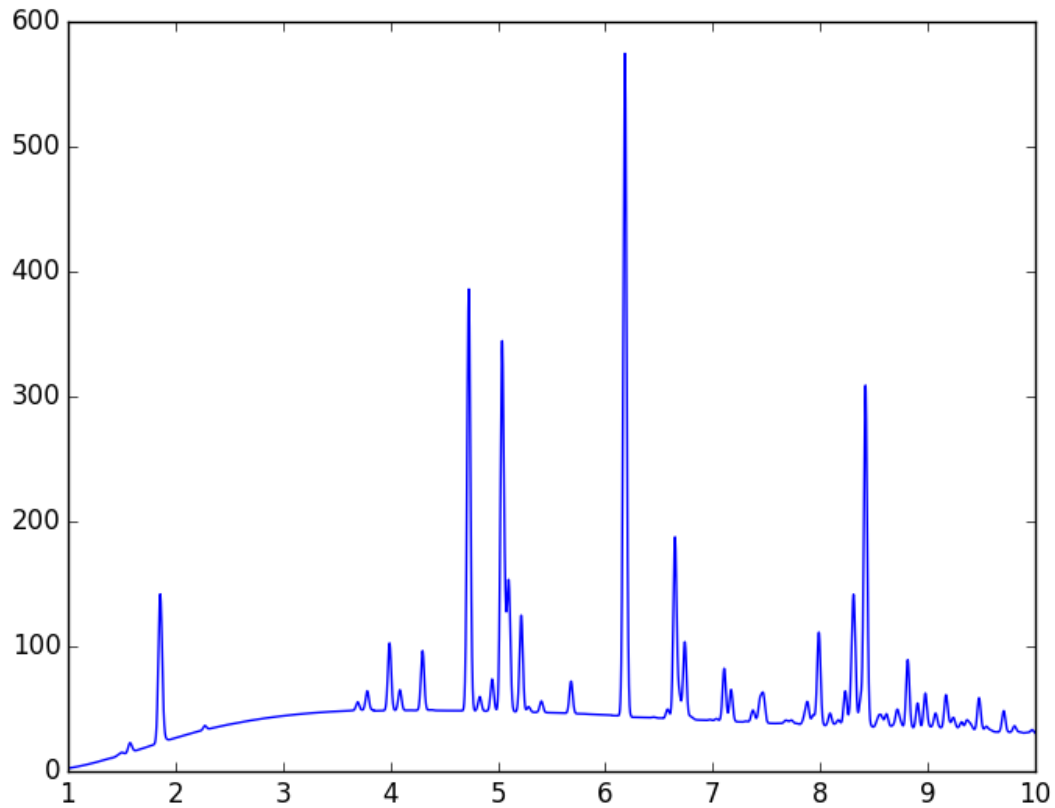
```
s2 = ch.ipyspectrum(temp, dens, wvl, filter = (chfilters.gaussian,.2), em = emeasure,
    ↪ doContinuum=0, keepIons=1, elementList=['si'], minAbund=1.e-4)
plt.subplot(211)
plt.plot(wvl,s2.Spectrum['intensity'][0])
plt.ylabel(r'erg cm$^{-2}$ s$^{-1}$ sr$^{-1}$ \AA$^{-1}$')
plt.subplot(212)
plt.plot(wvl,s2.IonInstances['si_9'].Spectrum['intensity'][0])
plt.ylabel(r'erg cm$^{-2}$ s$^{-1}$ sr$^{-1}$ \AA$^{-1}$')
plt.xlabel(r'Wavelength ($\AA$)')
plt.title('Si IX')
```



Because **keepIons** has been set, the ion instances of all of the ions are maintained in the `s2.IonInstances` dictionary. It has been possible to compare the spectrum of all of the ions with the spectrum of a single ion.

```
temp=2.e+7
dens=1.e+9
wvl = 1. + 0.002*np.arange(4501)
s3 = ch.ipyspectrum(temp, dens, wvl, filter = (chfilters.gaussian,.015),
    ↪doContinuum=1, em=1.e+27,minAbund=1.e-5,verbose=0)
plt.plot(wvl, s3.Spectrum['intensity'])
```

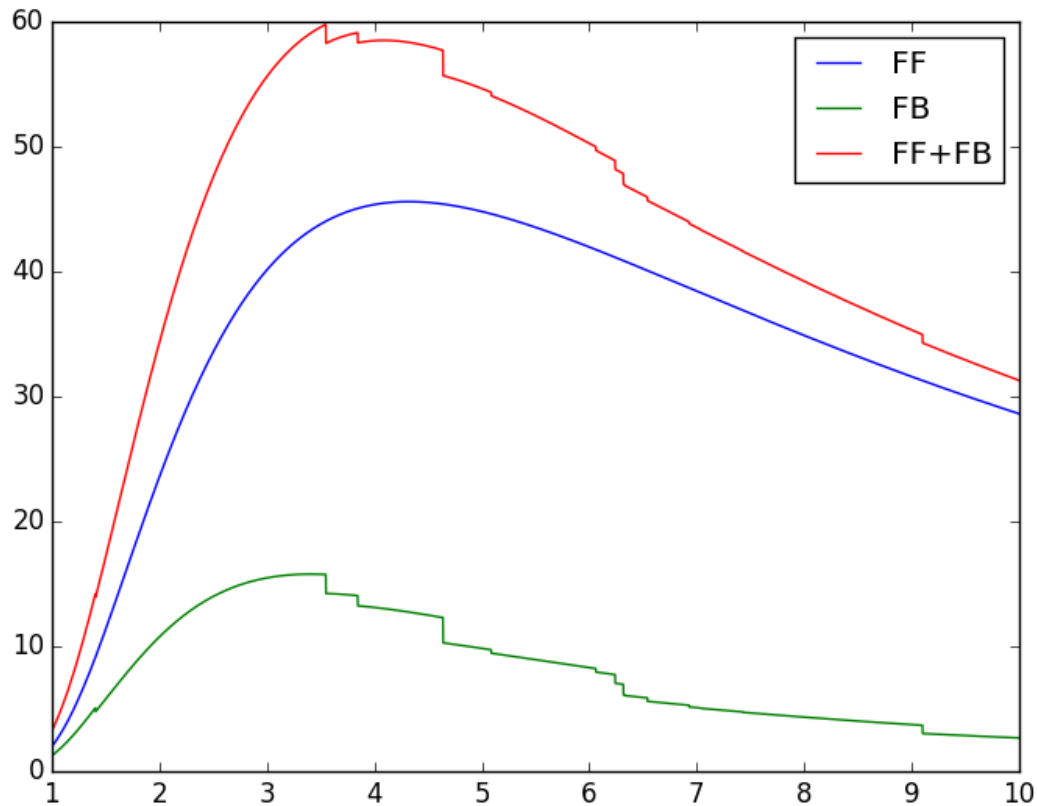




with `doContinuum=1`, the continuum can be plotted separately

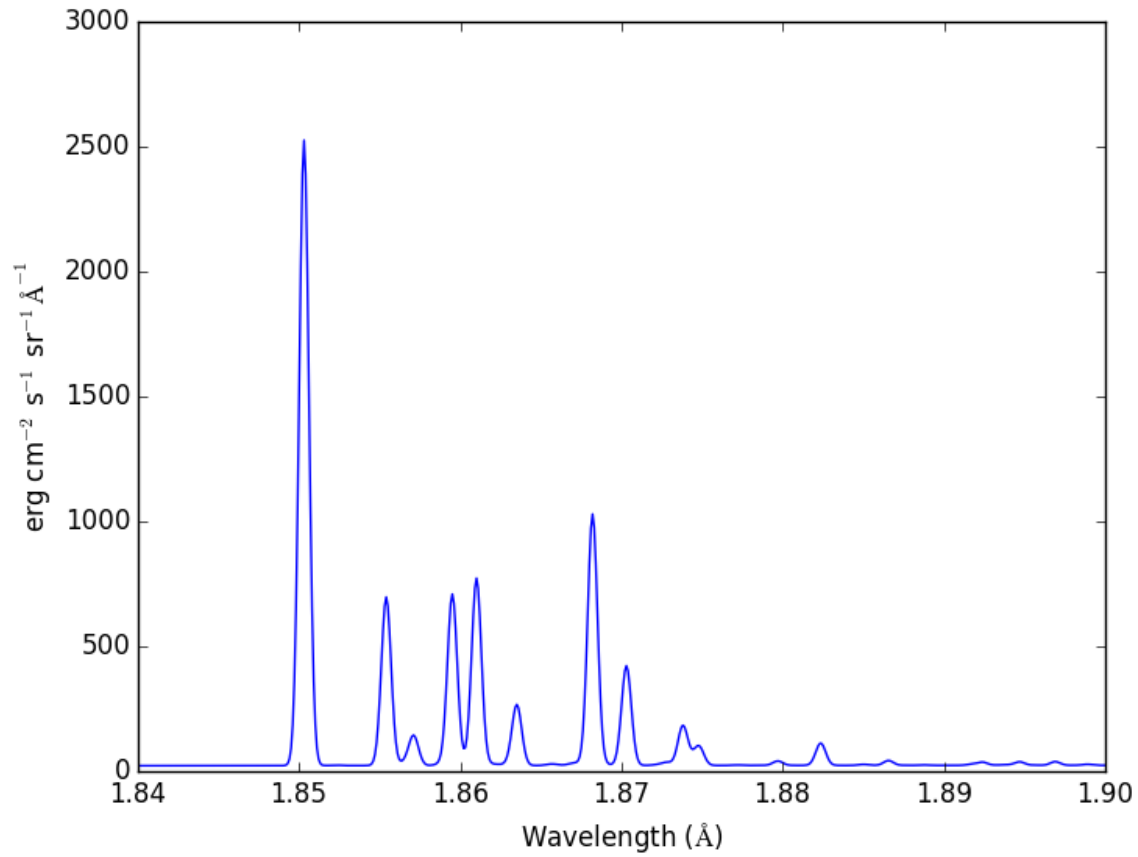
```
plot(wvl, s3.Spectrum['intensity']) plot(wvl,s.FreeFree['intensity'])
plot(wvl,s.FreeBound['intensity'])
plot(wvl,s.FreeBound['intensity']+s.FreeFree['intensity'])
```

produces



```
temperature = 2.e+7
density = 1.e+9
em=1.e+27
wvl = 1.84 + 0.0001*arange(601)
s4 = ch.ipyspectrum(temperature, density ,wvl, filter = (chfilters.gaussian,.0003),
doContinuum=1, minAbund=1.e-5, em=em, verbose=0)
```

produces

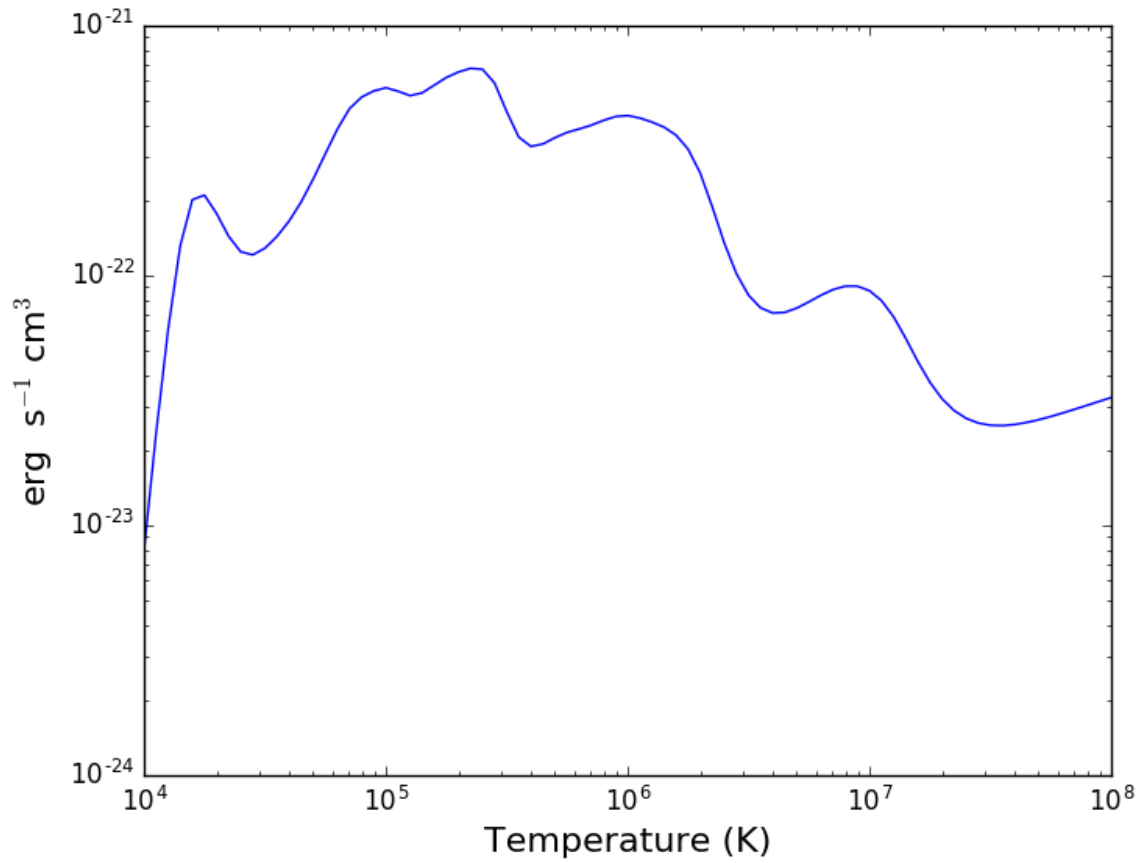


## 2.11 Radiative loss rate

the radiative loss rate can be calculated as a function of temperature and density:

```
temp = 10.**(4.+0.05*arange(81))
rl = ch.radLoss(temp, 1.e+4, minAbund=2.e-5)
rl.radLossPlot()
```

produces, in 446 s:



the radiative losses are kept in the `rl.RadLoss` dictionary

the **abundance** keyword argument can be set to the name of an available abundance file in `XUVTOP/abund`

if `abundance='abc'`, or some name that does not match an abundance name, a dialog will come up so that a abundance file can be selected

## 3.1 The ChiantiPy Approach

Python is a modern, object-oriented programming language. It provides a number of features that are useful to the programmer and the end user, such as, classes with methods (function-like) and attributes (data), and functions, among other things. ChiantiPy has been constructed so that the primary means to calculate the spectral properties of ions and groups of ions is by way of Python classes.

More detailed information can be found in the **API** Reference.

### 3.1.1 ChiantiPy Classes

There are 7 basic classes that are provided by ChiantiPy.

#### **ion**

this class is very useful in itself and is the basic unit employed by all of the other classes

#### **continuum**

for calculating the free-free (*bremstrahlung*), free-bound (radiative recombination) continuum as well as the radiative loss rates due to these processes.

#### **bunch**

allows the user to specify a *bunch* of ions and to calculate the radiative properties of the selected ions in a group. The ions can be specified by list of individual ions, list of elements, as well as by the minimum elemental abundance. The properties of each ion are available, as with members of the **ion** class. Among other things, the ratios of lines of different ions and elements can be calculated and then displayed.

#### **spectrum**

the **spectrum** class calculates the intensities of the line and continuum and then convolves the complete spectrum with a filter such as a Gaussian of specified width

there are actually 3 *spectrum* classes. Two of these all the use of multiple cpu cores to speed the calculation. The basic **spectrum** class does not do multi-processes and is therefore compatible with most Python environments

#### **mspectrum**

**mspectrum** duplicates the calculations of the **spectrum** class but it employs the Python multiprocessing package in order to use multiple cpu cores to calculate the spectrum. This class can be used in a basic Python shell, in a Python script, or in an IPython terminal. It can not be used in either the Jupyter qtconsole or the Jupyter notebook.

**ipynmspectrum**

this class employs the IPython `ipyparallel` module to provide access to multiple cpu cores. It can only be used in the Jupyter `qtconsole` and the Jupyter notebook.

**ioneq**

this class allows the user to load and plot the ionization equilibria of a specific element. It can read the ionization equilibrium files in the CHIANTI `$XUVTOP/ioneq` directory. Different ionization equilibria can be plotted against each other

it is also possible to calculate the ionization equilibria of an individual element using the ionization and recombination rates in the CHIANTI database. The results of this calculation can also be plotted against existing calculations in the CHIANTI `$XUVTOP/ioneq` directory.

### 3.1.2 ChiantiPy Classes, Methods and Attributes

Each of the ChiantiPy classes listed above has a number of methods for calculating various properties. The results of these calculations are stored as attributes of the class that has been instantiated (created). In Python, all objects, which includes everything in Python, have introspection so that all methods and attributes can be discovered and used. The IPython terminal and the Jupyter `qtconsole` both provide their own methods of easily displaying the methods and attributes.

Some methods in each class are more useful to the user than others. For example, below the `populate()` method is demonstrated below. However, it is generally not necessary for the user to use the `populate()` method. Methods that need the ion population will make use that the `Population` attribute is available and, if not, use the `populate()` method to create it.

Below, the methods that are most likely of interest to users are listed below. All of the available methods are presented and documented in the API section, a part of the ChiantiPy documentation.

**ion****popPlot()** method

plots the level populations of the *top* (most highly populated) levels as a function of temperature and/or density.

**gofnt()** method

an interactive method that plots the most intense lines of an ion in a give wavelength range (*wvlRange*) and allows the user to select a line or several lines, that will be summed, and then plots the *GofT* function for the selected lines and saves these values in the **Gofnt** dictionary as an attribute of the ion object.

**emiss()** method

calculates the spectral line emissivities of the ion and saves these in the **Emiss** dictionary as an attribute of the ion object.

**intensity()** method

calculates the *intensity* of the specified ion as a function of temperature and density. These properties are saved in the **Intensity** dictionary, available as an attribute of the ion.

**intensityList()** method

lists the spectral line intensities in a given wavelength range (*wvlRange*) in an interactive terminal or notebook

**intensityPlot()** method

plots the spectral line intensities in a given wavelength range (*wvlRange*) for the *top* most intensity lines.

#### **intensityRatio()** method

an interactive method that plots the most intense lines of an ion in a given wavelength range (*wvlRange*) and allows the user to select a pair of lines or a pair of lines to be summed and then plots the intensity ratio as a function of temperature and/or density. The ratio is saved in the **IntensityRatio** dictionary as an attribute of the ion.

#### **spectrum()** method

calculates the spectrum of the ion as a function of wavelength. The spectral line intensities are passed through a selectable filter to simulate the spectrometer line profile. The spectrum is saved in the **Spectrum** dictionary as an attribute of the ion.

#### **ionizRate()** method

calculates the ionization rate coefficient as a function of temperature. The rate coefficient is saved in the **IonizRate** dictionary as an attribute. Uses the methods **diRate()** and **eaRate()** to first calculate the direct and excitation-ionization (ea) rate coefficients and sums them.

#### **recombRate()** method

calculates the recombination rate coefficient as a function of temperature. The rate coefficient is saved in the **RecombRate** dictionary as an attribute. Uses the methods **rrRate()** and **drRate()** to first calculate the radiative recombination and dielectronic recombination rate coefficients and sums them.

### **ioneq**

#### **load()** method

reads a selected, existing ionization equilibrium calculation for a given element and saves it as a numpy array **Ioneq** as an attribute of the object.

#### **calculate()** method

calculates the ionization equilibrium of a selected element from the CHIANTI ionization and recombination rates for a specified temperature(s) and saves it as a numpy array **Ioneq** as an attribute of the object.

#### **plot()** method

plots the loaded or calculated ionization equilibrium. Various parameters can be specified to plot only those aspects that are desired. Can also plot an additional existing ionization equilibrium for comparison

### **bunch**

the **init** method calculates the spectral line intensities for the selection of ions saved in the **Intensity** dictionary as an attribute. It does not calculate the continuum.

beyond the **init** method, the **bunch** class inherits all of the following methods that are described under the **ion** class above

#### **intensityList()**

#### **intensityPlot()**

#### **intensityRatio()**

in addition, it inherits the following methods that are described under the **spectrum** class below

**convolve()**

**lineSpectrumPlot()**

**spectrumPlot()**

#### **spectrum**

the `init` method calculates the spectral line intensities and the continuum due to the free-free (*bremstrahlung*), free-bound (radiative recombination), and two-photon processes. The line intensities are convolved using the **convolve()** method (below). The sum is saved in the **Spectrum** dictionary as an attribute.

beyond the `init` method, the **spectrum** class also inherits the same methods as the **bunch** class including **intensityList()**, **intensityPlot**, and **intensityRatio**.

**convolve()**

convolves the line spectrum with specified filter from **ChiantiPy.tools.filters** using a specified width.

**lineSpectrumPlot()**

plots the convolved line spectrum as a function wavelength

**spectrumPlot()**

plots the spectrum calculated by the `init` method. The summed (integrated) spectrum can be plotted or the spectrum for a specific temperature can be plotted.

#### **mspectrum**

the **mspectrum** behaves in the same way as the **spectrum** class except that it invokes the Python multiprocessing module so that the calculations are made using a specified number of cpu cores. **mspectrum** can not be used in the Jupyter qtconsole or notebook.

#### **ipynmspectrum**

the **ipynmspectrum** behaves in the same way as the **spectrum** class except that it invokes the IPython ipyparallel module so that the calculations are made using a specified number of cpu cores. **ipynmspectrum** can only be used in the IPython terminal or the Jupyter qtconsole or notebook.

## 3.2 The ion class, basic properties

Bring up a Python session, or better yet, an IPython session

```
import ChiantiPy.core as ch
fe14 = ch.ion('fe_14')
```

The `fe14` object is instantiated with a number of methods and data. Methods start with lowercase letters and attributes start with uppercase letters. It is best not to simply import `ion` as there is a method with the same name in `matplotlib`. A few examples:

```
fe14.IonStr
>> 'fe_14'
fe14.Spectroscopic
>> 'Fe XIV'
```

CHIANTI and spectroscopic notation for the ion



```
fe14.Z
>> 26
fe14.Ion
>> 14
```

nuclear charge Z and the ionization stage (in spectroscopic notation) for the ion

```
fe14.Ip
>> 392.16196
```

this is the ionization potential in electron volts.

```
fe14.FIP
>> 7.9023801573028294
```

this is the first ionization potential (FIP) in electron volts - the ionization potential of the neutral (Fe I).

```
fe14.Abandance
>> 0.00012589265
fe14.AbandanceName
>> 'sun_photospheric_1998_grevesse'
```

this is the abundance of iron relative to hydrogen for the specified elemental abundance set. For the ion class, the abundance can be specified by the *abundance* keyword argument or the *abundanceName* keyword argument. In the case the abundance is taken from default abundance set. The specified defaults can be examined by

```
fe14.Defaults
>> {'abundfile': 'sun_photospheric_1998_grevesse', 'flux': 'energy', 'ioneqfile':
↳ 'chianti', 'wavelength': 'angstrom'}
```

the defaults can be specified by the user in the `~/chiantirc/chiantirc` file. One is included in the distribution but it must be placed in `~/chiantirc` for it to be read. If it is not found, a set of coded default values are used.

```
fe14.Elvlc.keys()
>> ['ecmth', 'term', 'ref', 'pretty', 'spd', 'ecm', 'j', 'l', 'erydth', 'conf', 'lvl
↳ ', 'spin', 'eryd', 'mult']
```

`fe14.Elvlc` is a dictionary that describes the energy levels of the Fe XIV ion. The key 'ecm' provides the energies, relative to the ground level, in inverse cm. The 'ref' key provides the references in the scientific literature where the data were provided.

```
fe14.Elvlc['ref']
>> [%filename: fe_14.elvlc,
    %observed energy levels: Churilov S.S., Levashov V.E., 1993, Physica Scripta 48,
↳ 425,
    %observed energy levels: Redfors A., Litzen U., 1989, J.Opt.Soc.Am.B 6, #8,
↳ 1447,
    %theoretical energy levels: Storey P.J., Mason H.E., Young P.R., 2000, A&ASS
↳ 141, 28,
    %comment,
    Only level 16 does not have an observed energy. I have placed in,
    the third energy column a recommended value for the energy value of,
    this level, based on the theoretical and observed splittings of the,
    4F levels. It is this energy value which is used to compute the,
    wavelengths of transitions involving level 16 given in the .wgfa,
    file.,
    %produced as part of the Arcetri/Cambridge/GMU/NRL 'CHIANTI' atomic data base
↳ collaboration,
```

(continues on next page)

(continued from previous page)

```
% ,  
% P.R.Young Feb 99']
```

If the `fe14 ion` object had be instantiated (created) with a temperature and an electron density, then many more attributes can be calculated. For example, if the `populate()` method is used, it creates a dictionary attribute *Population*. One thing to remember with Python is that capitalization matters.

```
import numpy as np  
t = 10.** (5.8+0.1*np.arange(11))  
dens = 1.e+9  
fe14 = ch.ion('fe_14')  
fe14.populate()  
fe14.Population.keys()  
>> ['ci', 'protonDensity', 'popmat', 'eDensity', 'rec', 'population', 'temperature']  
  
fe14.Population['population'].shape  
>> (21, 739)  
  
'%10.2e'%(fe14.Temperature[10])  
>> ' 2.00e+06'  
  
fe14.Population['population'][10,:5]  
>> array([ 8.71775703e-01, 1.27867444e-01, 4.91230626e-09, 4.29120495e-08, 1.35517895e-  
↪ 08])
```

gives the population of the first 5 of 739 levels of Fe XIV at a temperature of 2.00e+6

to be continued

## CHIANTIPY'S API DOCUMENTATION

### 4.1 ChiantiPy package

#### 4.1.1 Subpackages

ChiantiPy.Gui package

Subpackages

ChiantiPy.Gui.gui\_cl package

Submodules

ChiantiPy.Gui.gui\_cl.gui module

Command line selection dialogs.

**class** ChiantiPy.Gui.gui\_cl.gui.**choice2Dialog** (*items, label=None, parent=None*)

Bases: `object`

Make a single or multiple selection from a list of items and another single or multiple selection from the same list.

Useful for picking numerators and denominators.

expects the input of an array of items, will select one or more from both widgets the keywords label and parent are there for consistency with real gui dialogs

ChiantiPy.Gui.gui\_cl.gui.**chpicker** (*path, filter='\*.\*', label=None*)

Select a filename from using a command line dialog.

the label keyword is included for consistency but does nothing

**class** ChiantiPy.Gui.gui\_cl.gui.**selectorDialog** (*items, label=None, parent=None*)

Bases: `object`

Make a single or multiple selection from a list of items.

expects the input of an array of items, will select one or more the label and parent keywords are for consistency with other modules but do nothing

## Module contents

command-line selection dialogs

## ChiantiPy.Gui.gui\_qt5 package

### Submodules

#### ChiantiPy.Gui.gui\_qt5.gui module

PyQt5 widget selection dialogs

```
class ChiantiPy.Gui.gui_qt5.gui.choice2Dialog (items, label=None, parent=None, multi=True)
    Bases: PyQt5.QtWidgets.QDialog
    Make a single or multiple selection from a list of items and another single or multiple selection from the same list.
    Useful for picking numerators and denominators.
    expects the input of an array of items, will select one or more from both widgets.
    accept (self)
    reject (self)

class ChiantiPy.Gui.gui_qt5.gui.chpicker (dir, label)
    Bases: PyQt5.QtWidgets.QWidget
    dialog to select a single file name under the directory code largely taken from pythonspot.com
    initUI ()
    openFileNameDialog ()

class ChiantiPy.Gui.gui_qt5.gui.selectorDialog (items, label=None, parent=None, multi-Choice=True)
    Bases: PyQt5.QtWidgets.QDialog
    Make a single or multiple selection from a list of items.
    expects the input of an array of items, will select one or more
    accept (self)
    reject (self)
```

#### ChiantiPy.Gui.gui\_qt5.ui module

```
class ChiantiPy.Gui.gui_qt5.ui.Ui_choice2DialogForm
    Bases: object
    retranslateUi (choice2DialogForm)
    setupUi (choice2DialogForm)

class ChiantiPy.Gui.gui_qt5.ui.Ui_selectorDialogForm
    Bases: object
    retranslateUi (selectorDialogForm)
```

**setupUi** (*selectorDialogForm*)

## Module contents

PyQt5 selection dialog widgets

## Module contents

Select GUI package

## ChiantiPy.base package

### Submodules

### ChiantiPy.base.\_IonTrails module

Base classes used in the *ChiantiPy.core.ion* and *ChiantiPy.core.spectrum* classes. Mostly printing, plotting and saving routines.

**class** ChiantiPy.base.\_IonTrails.**ionTrails**

Bases: `object`

Base class for *ChiantiPy.core.ion* and *ChiantiPy.core.spectrum*

**argCheck** (*temperature=None, eDensity=None, pDensity='default', em=None, verbose=0*)  
to check the compatibility of the three arguments and put them into numpy arrays of `atleast_1d`

**intensityList** (*index=-1, wvlRange=None, wvlRanges=None, top=10, relative=0, outFile=0, right-Digits=4*)

List the line intensities. Checks to see if there is an existing Intensity attribute. If it exists, then those values are used. Otherwise, the *intensity* method is called.

This method prints an ASCII table with the following columns:

1. Ion: the CHIANTI style notation for the ion, e.g. 'c\_4' for C IV
2. lv11: the lower level of the transition in the CHIANTI .elvlc file
3. lv12: the upper level of the transition in the CHIANTI .elvlc file
4. lower: the notation, usually in LS coupling, of the lower fine structure level
5. upper: the notation, usually in LS coupling, of the upper fine structure level
6. Wvl(A): the wavelength of the transition in units as specified in the chiantirc file.
7. Intensity
8. A value: the Einstein coefficient for spontaneous emission from level 'j' to level 'i'
9. Obs: indicates whether the CHIANTI database considers this an observed line or one obtained from theoretical energy levels

Regarding the intensity column, if 'flux' in the chiantirc file is set to 'energy', the intensity is given by,

$$I = \Delta E_{ij} n_j A_{ij} \text{Ab} \frac{1}{N_e} \frac{N(X^{+m})}{N(X)} \text{EM},$$

in units of  $\text{ergs cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$ . If ‘flux’ is set to ‘photon’,

$$I = n_j A_{ij} \text{Ab} \frac{1}{N_e} \frac{N(X^{+m})}{N(X)} \text{EM},$$

where,

- $\Delta E_{ij}$  is the transition energy (ergs)
- $n_j$  is the fractions of ions in level  $j$
- $A_{ij}$  is the Einstein coefficient for spontaneous emission from level  $j$  to level  $i$  (in  $\text{s}^{-1}$ )
- Ab is the abundance of the specified element relative to hydrogen
- $N_e$  is the electron density (in  $\text{cm}^{-3}$ )
- $N(X^{+m})/N(X)$  is the fractional ionization of ion as a function of temperature
- EM is the emission measure integrated along the line-of-sight,  $\int dl N_e N_H$  ( $\text{cm}^{-5}$ ) where  $N_H$  is the density of hydrogen (neutral + ionized) ( $\text{cm}^{-3}$ )

Note that if *relative* is set, the line intensity is relative to the strongest line and so the output will be unitless.

#### Parameters

- **index** (*int*, optional) – Index the temperature or eDensity array to use. -1 (default) sets the specified value to the middle of the array
- **wvlRange** (*tuple*) – Wavelength range
- **wvlRanges** (*a tuple, list or array that contains at least 2*) – 2 element tuples, lists or arrays so that multiple wavelength ranges can be specified
- **top** (*int*) – Number of lines to plot, sorted by descending magnitude.
- **relative** (*int*) – specifies whether to normalize to strongest line default (relative = 0) specified that the intensities should be their calculated values
- **outFile** (*str*) – specifies the file that the intensities should be output to default(outFile = 0) intensities are output to the terminal
- **rightDigits** (*int*) – specifies the format for the wavelengths for the number of digits to right of the decimal place

**intensityPlot** (*index=-1, wvlRange=None, top=10, linLog='lin', relative=False, verbose=False, plotFile=0, em=0*)

Plot the line intensities. Uses *Intensity* if it already exists. If not, call the *intensity* method.

#### Parameters

- **index** (*integer*) – specified which value of the temperature array or eDensity array to use. default (index=-1) sets the specified value to the middle of the array
- **wvlRange** (*2 element tuple, list or array determines the wavelength range*)
- **top** (*integer*) – specifies to plot only the top strongest lines, default = 10
- **linLog** (*str*) – default(‘lin’) produces a plot where the intensity scale is linear if set to ‘log’, produces a plot where the intensity scale is logarithmic
- **normalize** (*= 1 specifies whether to normalize to strongest line*) – default (relative = 0) specified that the intensities should be their calculated values
- **plotFile** – default=0, the plot is not saved to a file otherwise, the plot is saved to the ‘plot-File’

- **em** (*emission measure*) – if an Intensity attribute needs be created, then the emission measure is applied

**intensityRatio** (*wvlRange=None, wvlRanges=None, top=10*)

Plot the intensity ratio of 2 lines or sums of lines. Shown as a function of density and/or temperature. For a single wavelength range, set *wvlRange* = [*wMin*, *wMax*] For multiple wavelength ranges, set *wvlRanges* = [[*wMin1*,*wMax1*],[*wMin2*,*wMax2*], ...] A plot of relative emissivities is shown and then a dialog appears for the user to choose a set of lines.

#### Parameters

- **wvlRange** (*array-like*) – Wavelength range, i.e. min and max
- **wvlRanges** (*a tuple, list or array that contains at least 2*) – 2 element tuples, lists or arrays so that multiple wavelength ranges can be specified
- **top** (*int*) – specifies to plot only the top strongest lines, default = 10

**intensityRatioSave** (*outFile=0*)

Save the intensity ratio to a file.

The intensity ratio as a function to temperature and eDensity is saved to an ascii file. Descriptive information is included at the top of the file.

**Parameters** **outFile** – default(0): the plot of the intensity ratio is not saved str/unicode: the plot is saved to the file names 'outFile'

## ChiantiPy.base.\_SpecTrails module

Base class used in several ChiantiPy objects

**class** ChiantiPy.base.\_SpecTrails.**specTrails** (*temperature, density*)

Bases: `object`

a collection of methods for use in spectrum calculations

**convolve** (*wavelength=0, filter=(*<function gaussianR>*, 1000.0), label=0, verbose=0*)

the first application of spectrum calculates the line intensities within the specified wavelength range and for set of ions specified

wavelength will not be used if applied to 'spectrum' objects

wavelength IS need for 'bunch' objects - in this case, the wavelength should not extend beyond the limits of the wvlRange used for the 'bunch' calculation

**ionGate** (*elementList=None, ionList=None, minAbund=None, doLines=1, doContinuum=1, doWvlTest=1, verbose=0*)

creates a list of ions for free-free, free-bound, and line intensity calculations if doing the radiative losses, accept all wavelength -> doWvlTest=0 the list is a dictionary self.TODO

**lineSpectrumPlot** (*index=0, integrated=0, saveFile=0, linLog='lin'*)

to plot the line spectrum as a function of wavelength

**spectrumPlot** (*index=-1, integrated=0, saveFile=0, linLog='lin'*)

to plot the spectrum as a function of wavelength

## Module contents

Base classes for ion- and spectrum-related objects.

## ChiantiPy.core package

### Subpackages

### ChiantiPy.core.tests package

### Submodules

#### ChiantiPy.core.tests.test\_Continuum module

#### ChiantiPy.core.tests.test\_Ion module

#### ChiantiPy.core.tests.test\_Ioneq module

Tests for ioneq class

```
ChiantiPy.core.tests.test_Ioneq.test_calculate_ioneq()  
ChiantiPy.core.tests.test_Ioneq.test_el_input()  
ChiantiPy.core.tests.test_Ioneq.test_el_z_inputs_same()  
ChiantiPy.core.tests.test_Ioneq.test_load_ioneq()  
ChiantiPy.core.tests.test_Ioneq.test_load_ioneq_alternate_file()  
ChiantiPy.core.tests.test_Ioneq.test_z_input()
```

#### ChiantiPy.core.tests.test\_Spectrum module

Tests for the spectrum and bunch classes

```
ChiantiPy.core.tests.test_Spectrum.test_bunch()  
ChiantiPy.core.tests.test_Spectrum.test_spectrum_array()  
ChiantiPy.core.tests.test_Spectrum.test_spectrum_scalar()
```

## Module contents

### Submodules

#### ChiantiPy.core.Continuum module

Continuum module

```
class ChiantiPy.core.Continuum.continuum(ionStr, temperature, abundance=None,  
                                         em=None, verbose=0)  
    Bases: ChiantiPy.base._IonTrails.IonTrails
```



The top level class for continuum calculations. Includes methods for the calculation of the free-free and free-bound continua.

### Parameters

- **ionStr** (*str*) – CHIANTI notation for the given ion, e.g. ‘fe\_12’ that corresponds to the Fe XII ion.
- **temperature** (*array-like*) – In units of Kelvin
- **abundance** (*float* or *str*, optional) – Elemental abundance relative to Hydrogen or name of CHIANTI abundance file, without the ‘.abund’ suffix, e.g. ‘sun\_photospheric\_1998\_grevesse’.
- **em** (*array-like*, *optional*) – Line-of-sight emission measure ( $\int dl n_e n_H$ ), in units of  $\text{cm}^{-5}$ , or the volumetric emission measure ( $\int dV n_e n_H$ ) in units of  $\text{cm}^{-3}$ .

### Examples

```
>>> import ChiantiPy.core as ch
>>> import numpy as np
>>> temperature = np.logspace(4, 9, 20)
>>> cont = ch.continuum('fe_15', temperature)
>>> wavelength = np.arange(1, 1000, 10)
>>> cont.freeFree(wavelength)
>>> cont.freeBound(wavelength, include_abundance=True, include_ioneq=False)
>>> cont.calculate_free_free_loss()
>>> cont.calculate_free_bound_loss()
```

### Notes

The methods for calculating the free-free and free-bound emission and losses return their result to an attribute. See the respective docstrings for more information.

### References

#### **calculate\_free\_bound\_loss** (*\*\*kwargs*)

Calculate the free-bound energy loss rate of an ion. The result is returned to the *free\_bound\_loss* attribute.

The free-bound loss rate can be calculated by integrating the free-bound emission over the wavelength. This is difficult using the expression in *calculate\_free\_bound\_emission* so we instead use the approach of<sup>108</sup> and<sup>106</sup>. Eq. 1a of<sup>106</sup> can be integrated over wavelength to get the free-bound loss rate,

$$\frac{dW}{dt dV} = C_{ff} \frac{k}{hc} T^{1/2} G_{fb},$$

in units of  $\text{erg cm}^3 \text{s}^{-1}$  where  $G_{fb}$  is the free-bound Gaunt factor as given by Eq. 15 of<sup>106</sup> (see *mewe\_gaunt\_factor* for more details) and  $C_{ff}$  is the numerical constant as given in Eq. 4 of<sup>108</sup> and can be written in terms of the fine structure constant  $\alpha$ ,

$$C_{ff} \frac{k}{hc} = \frac{8}{3} \left( \frac{\pi}{6} \right)^{1/2} \frac{h^2 \alpha^3}{\pi^2} \frac{k_B}{m_e^{3/2}} \approx 1.43 \times 10^{-27}$$

<sup>108</sup> Gronenschild, E.H.B.M. and Mewe, R., 1978, A&AS, 32, 283

<sup>106</sup> Mewe, R. et al., 1986, A&AS, 65, 511

**freeBound** (*wvl*, *verner=1*)

to calculate the free-bound (radiative recombination) continuum rate coefficient of an ion, where the ion is taken to be the target ion, including the elemental abundance and the ionization equilibrium population uses the Gaunt factors of Karzas, W.J, Latter, R, 1961, ApJS, 6, 167 for recombination to the ground level, the photoionization cross sections of Verner and Yakovlev, 1995, A&ASS, 109, 125 are used to develop the free-bound cross section includes the elemental abundance and the ionization fraction provides emissivity = ergs cm<sup>-2</sup> s<sup>-1</sup> str<sup>-1</sup> Angstrom<sup>-1</sup>

**freeBoundEmiss** (*wvl*, *verner=1*)

Calculates the free-bound (radiative recombination) continuum emissivity of an ion. Provides emissivity in units of ergs cm<sup>-2</sup> s<sup>-1</sup> str<sup>-1</sup> for an individual ion.

### Notes

- Uses the Gaunt factors of<sup>103</sup> for recombination to the ground level
- Uses the photoionization cross sections of<sup>2</sup> to develop the free-bound cross section
- Does not include the elemental abundance or ionization fraction
- The specified ion is the target ion

### References

**freeBoundLoss** ()

to calculate the free-bound (radiative recombination) energy loss rate coefficient of an ion, the ion is taken to be the target ion, including the elemental abundance and the ionization equilibrium population uses the Gaunt factors of<sup>103</sup> Karzas, W.J, Latter, R, 1961, ApJS, 6, 167 provides rate = ergs cm<sup>-2</sup> s<sup>-1</sup>

**freeBoundLossMewe** (*\*\*kwargs*)

Calculate the free-bound energy loss rate of an ion. The result is returned to the *free\_bound\_loss* attribute.

The free-bound loss rate can be calculated by integrating the free-bound emission over the wavelength. This is difficult using the expression in *calculate\_free\_bound\_emission* so we instead use the approach of<sup>108</sup> and<sup>106</sup>. Eq. 1a of<sup>106</sup> can be integrated over wavelength to get the free-bound loss rate,

$$\frac{dW}{dtdV} = C_{ff} \frac{k}{hc} T^{1/2} G_{fb},$$

in units of erg cm<sup>3</sup> s<sup>-1</sup> where  $G_{fb}$  is the free-bound Gaunt factor as given by Eq. 15 of<sup>106</sup> (see *mewe\_gaunt\_factor* for more details) and  $C_{ff}$  is the numerical constant as given in Eq. 4 of<sup>108</sup> and can be written in terms of the fine structure constant  $\alpha$ ,

$$C_{ff} \frac{k}{hc} = \frac{8}{3} \left(\frac{\pi}{6}\right)^{1/2} \frac{h^2 \alpha^3}{\pi^2} \frac{k_B}{m_e^{3/2}} \approx 1.43 \times 10^{-27}$$

**freeBoundWB** (*wavelength*, *includeAbundance=True*, *includeIoneq=True*, *useVerner=True*, *\*\*kwargs*)

Calculate the free-bound emission of an ion. The result is returned as a 2D array to the *free\_bound\_emission* attribute.

The total free-bound continuum emissivity is given by,

$$\frac{dW}{dtdVd\lambda} = \frac{1}{4\pi} \frac{2}{hk_B c^3 m_e \sqrt{2\pi k_B m_e}} \frac{E^5}{T^{3/2}} \sum_i \frac{\omega_i}{\omega_0} \sigma_i^{bf} \exp\left(-\frac{E - I_i}{k_B T}\right)$$

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<sup>103</sup> Karzas and Latter, 1961, ApJSS, 6, 167

<sup>2</sup> Verner & Yakovlev, 1995, A&AS, 109, 125

where  $E = hc/\lambda$  is the photon energy,  $\omega_i$  and  $\omega_0$  are the statistical weights of the  $i^{\text{th}}$  level of the recombined ion and the ground level of the recombining ion, respectively,  $\sigma_i^{bf}$  is the photoionization cross-section, and  $I_i$  is the ionization potential of level  $i$ . This expression comes from Eq. 12 of<sup>105</sup>. For more information about the free-bound continuum calculation, see [Peter Young's notes on free-bound continuum](#).

The photoionization cross-sections are calculated using the methods of<sup>102</sup> for the transitions to the ground state and<sup>103</sup> for all other transitions. See `verner_cross_section` and `karzas_cross_section` for more details.

The free-bound emission is in units of  $\text{erg cm}^3\text{s}^{-1}\text{\AA}^{-1}\text{str}^{-1}$ . If the emission measure has been set, the units will be multiplied by  $\text{cm}^{-5}$  or  $\text{cm}^{-3}$ , depending on whether it is the line-of-sight or volumetric emission measure, respectively.

#### Parameters

- **wavelength** (*array-like*) – In units of angstroms
- **include\_abundance** (*bool*, optional) – If True, include the ion abundance in the final output.
- **include\_ioneq** (*bool*, optional) – If True, include the ionization equilibrium in the final output
- **use\_verner** (*bool*, optional) – If True, cross-sections of ground-state transitions using<sup>2</sup>, i.e. `verner_cross_section`

**Raises** `ValueError` – If no .fbvl file is available for this ion

**freeFree** (*wavelength*, *include\_abundance=True*, *include\_ioneq=True*, *\*\*kwargs*)

Calculates the free-free emission for a single ion. The result is returned as a dict to the `FreeFree` attribute. The dict has the keywords *intensity*, *wvl*, *temperature*, *em*.

The free-free emission for the given ion is calculated according Eq. 5.14a of<sup>107</sup>, substituting  $\nu = c/\lambda$ , dividing by the solid angle, and writing the numerical constant in terms of the fine structure constant  $\alpha$ ,

$$\frac{dW}{dtdVd\lambda} = \frac{c}{3m_e} \left( \frac{\alpha h}{\pi} \right)^3 \left( \frac{2\pi}{3m_e k_B} \right)^{1/2} \frac{Z^2}{\lambda^2 T^{1/2}} \exp \left( -\frac{hc}{\lambda k_B T} \right) \bar{g}_{ff},$$

where  $Z$  is the nuclear charge,  $T$  is the electron temperature in K, and  $\bar{g}_{ff}$  is the velocity-averaged Gaunt factor. The Gaunt factor is estimated using the methods of<sup>104</sup> and<sup>101</sup>, depending on the temperature and energy regime. See `itoh_gaunt_factor` and `sutherland_gaunt_factor` for more details.

The free-free emission is in units of  $\text{erg cm}^3\text{s}^{-1}\text{\AA}^{-1}\text{str}^{-1}$ . If the emission measure has been set, the units will be multiplied by  $\text{cm}^{-5}$  or  $\text{cm}^{-3}$ , depending on whether it is the line-of-sight or volumetric emission measure, respectively.

#### Parameters

- **wavelength** (*array-like*) – In units of angstroms
- **include\_abundance** (*bool*, optional) – If True, include the ion abundance in the final output.
- **include\_ioneq** (*bool*, optional) – If True, include the ionization equilibrium in the final output

**freeFreeLoss** (*\*\*kwargs*)

Calculate the free-free energy loss rate of an ion. The result is returned to the `FreeFreeLoss` attribute.

<sup>105</sup> Young et al., 2003, ApJSS, 144, 135

<sup>102</sup> Verner & Yakovlev, 1995, A&AS, 109, 125

<sup>107</sup> Rybicki and Lightman, 1979, Radiative Processes in Astrophysics, (Wiley-VCH)

<sup>104</sup> Itoh, N. et al., 2000, ApJS, 128, 125

<sup>101</sup> Sutherland, R. S., 1998, MNRAS, 300, 321

The free-free radiative loss rate is given by Eq. 5.15a of<sup>107</sup>. Writing the numerical constant in terms of the fine structure constant  $\alpha$ ,

$$\frac{dW}{dtdV} = \frac{4\alpha^3 h^2}{3\pi^2 m_e} \left( \frac{2\pi k_B}{3m_e} \right)^{1/2} Z^2 T^{1/2} \bar{g}_B$$

where where  $Z$  is the nuclear charge,  $T$  is the electron temperature, and  $\bar{g}_B$  is the wavelength-averaged and velocity-averaged Gaunt factor. The Gaunt factor is calculated using the methods of<sup>103</sup>. Note that this expression for the loss rate is just the integral over wavelength of Eq. 5.14a of<sup>103</sup>, the free-free emission, and is expressed in units of  $\text{erg cm}^3 \text{s}^{-1}$ .

**free\_free\_loss** (*\*\*kwargs*)

Calculate the free-free energy loss rate of an ion. The result is returned to the *free\_free\_loss* attribute.

The free-free radiative loss rate is given by Eq. 5.15a of<sup>107</sup>. Writing the numerical constant in terms of the fine structure constant  $\alpha$ ,

$$\frac{dW}{dtdV} = \frac{4\alpha^3 h^2}{3\pi^2 m_e} \left( \frac{2\pi k_B}{3m_e} \right)^{1/2} Z^2 T^{1/2} \bar{g}_B$$

where where  $Z$  is the nuclear charge,  $T$  is the electron temperature, and  $\bar{g}_B$  is the wavelength-averaged and velocity-averaged Gaunt factor. The Gaunt factor is calculated using the methods of<sup>103</sup>. Note that this expression for the loss rate is just the integral over wavelength of Eq. 5.14a of<sup>107</sup>, the free-free emission, and is expressed in units of  $\text{erg cm}^3 \text{s}^{-1}$ .

**ioneqOne** ()

Provide the ionization equilibrium for the selected ion as a function of temperature. Similar to but not identical to `ion.ioneqOne()` - the ion class needs to be able to handle the ‘dielectronic’ ions returned in `self.IoneqOne`

**ioneq\_one** (*stage, \*\*kwargs*)

Calculate the equilibrium fractional ionization of the ion as a function of temperature.

Uses the *ChiantiPy.core.ioneq* module and does a first-order spline interpolation to the data. An ionization equilibrium file can be passed as a keyword argument, *ioneqfile*. This can be passed through as a keyword argument to any of the functions that uses the ionization equilibrium.

**Parameters** *stage* (*int*) – Ionization stage, e.g. 25 for Fe XXV

**itoh\_gaunt\_factor** (*wavelength*)

Calculates the free-free gaunt factors of<sup>104</sup>.

An analytic fitting formulae for the relativistic Gaunt factor is given by Eq. 4 of<sup>104</sup>,

$$g_Z = \sum_{i,j=0}^{10} a_{ij} t^i U^j$$

where,

$$t = \frac{1}{1.25} (\log_{10} T - 7.25), \quad U = \frac{1}{2.5} (\log_{10} u + 1.5),$$

$u = hc/\lambda k_B T$ , and  $a_{ij}$  are the fitting coefficients and are read in using *ChiantiPy.tools.io.itohRead* and are given in Table 4 of<sup>104</sup>. These values are valid for  $6 < \log_{10}(T) < 8.5$  and  $-4 < \log_{10}(u) < 1$ .

**See also:**

*ChiantiPy.tools.io.itohRead()* Read in Gaunt factor coefficients from<sup>104</sup>

**karzasCross** (*photon\_energy, ionization\_potential, n, l*)

Calculate the photoionization cross-sections using the Gaunt factors of<sup>103</sup>.

The free-bound photoionization cross-section is given by,

$$\sigma_i^{bf} = 1.077294 \times 8065.54 \times 10^{16} \left( \frac{I_i}{hc} \right)^2 \left( \frac{hc}{E} \right)^3 \frac{g_{bf}}{n_i},$$

where  $I_i$  is the ionization potential of the  $i^{\text{th}}$  level,  $E$  is the photon energy,  $g_{bf}$  is the Gaunt factor calculated according to<sup>103</sup>, and  $n_i$  is the principal quantum number of the  $i^{\text{th}}$  level.  $\sigma_i^{bf}$  is units of  $\text{cm}^2$ . This expression is given by Eq. 13 of<sup>105</sup>. For more information on the photoionization cross-sections, see [Peter Young's notes on free-bound continuum](#).

#### Parameters

- **photon\_energy** (*array-like*)
- **ionization\_potential** (*float*)
- **n** (*int*)
- **l** (*int*)

**klgfbInterp** (*wvl, n, l*)

A Python version of the CHIANTI IDL procedure `karzas_xs`.

Interpolates free-bound gaunt factor of Karzas and Latter, (1961, Astrophysical Journal Supplement Series, 6, 167) as a function of wavelength (*wvl*).

**mewe\_gaunt\_factor** (*\*\*kwargs*)

Calculate the Gaunt factor according to<sup>106</sup> for a single ion  $Z_z$ .

Using Eq. 9 of<sup>106</sup>, the free-bound Gaunt factor for a single ion can be written as,

$$G_{fb}^{Z,z} = \frac{E_H}{k_B T} \text{Ab}(Z) \frac{N(Z,z)}{N(Z)} f(Z, z, n)$$

where  $E_H$  is the ground-state potential of H,  $\text{Ab}(Z)$  is the elemental abundance,  $\frac{N(Z,z)}{N(Z)}$  is the fractional ionization, and  $f(Z, z, n)$  is given by Eq. 10 and is approximated by Eq 16 as,

$$f(Z, z, n) \approx f_2(Z, z, n_0) = 0.9 \frac{\zeta_0 z_0^4}{n_0^5} \exp \left( \frac{E_H z_0^2}{n_0^2 k_B T} \right) + 0.42 \frac{z^4}{n_0^{3/2}} \exp \left( \frac{E_H z^2}{(n_0 + 1)^2 k_B T} \right)$$

where  $n_0$  is the principal quantum number,  $z_0$  is the effective charge (see Eq. 7 of<sup>106</sup>), and  $\zeta_0$  is the number of vacancies in the 0th shell and is given in Table 1 of<sup>106</sup>. Here it is calculated in the same manner as in [fb\\_rad\\_loss.pro](#) of the CHIANTI IDL library. Note that in the expression for  $G_{fb}$ , we have not included the  $N_H/n_e$  factor.

**Raises `ValueError`** – If no `.fbvl` file is available for this ion

**sutherland\_gaunt\_factor** (*wavelength*)

Calculates the free-free gaunt factor calculations of<sup>101</sup>.

The Gaunt factors of<sup>101</sup> are read in using `ChiantiPy.tools.io.gffRead` as a function of  $u$  and  $\gamma^2$ . The data are interpolated to the appropriate wavelength and temperature values using `~scipy.ndimage.map_coordinates`.

**vernerCross** (*wvl*)

Calculates the photoionization cross-section using data from<sup>102</sup> for transitions to the ground state.

The photoionization cross-section can be expressed as  $\sigma_i^{fb} = F(E/E_0)$  where  $F$  is an analytic fitting formula given by Eq. 1 of<sup>102</sup>,

$$F(y) = ((y-1)^2 + y_w^2) y^{-Q} (1 + \sqrt{y/y_a})^{-P},$$

where  $E$  is the photon energy,  $n$  is the principal quantum number,  $l$  is the orbital quantum number,  $Q = 5.5 + l - 0.5P$ , and  $\sigma_0, E_0, y_w, y_a, P$  are fitting paramters. These can be read in using *ChiantiPy.tools.io.vernerRead*.

## ChiantiPy.core.Ion module

Ion class

```
class ChiantiPy.core.Ion.ion(ionStr, temperature=None, eDensity=None, pDensity='default',
                             radTemperature=None, rStar=None, abundance=None,
                             setup=True, em=None, verbose=0)
Bases: ChiantiPy.base._IoneqOne.ioneqOne, ChiantiPy.base._IonTrails.ionTrails,
ChiantiPy.base._SpecTrails.specTrails
```

The top level class for performing spectral calculations for an ion in the CHIANTI database.

### Parameters

- **ionStr** (*str*) – CHIANTI notation for the given ion, e.g. ‘fe\_12’ that corresponds to the *Fe XII* ion.
- **temperature** (*float*, *tuple*, *list*, or *~numpy.ndarray*, optional) – Temperature array (Kelvin)
- **eDensity** (*float*, *tuple*, *list*, or *~numpy.ndarray*, optional) – Electron density array ( $\text{cm}^{-3}$ )
- **pDensity** (*float*, *tuple*, *list* or *~numpy.ndarray*, optional) – Proton density ( $\text{cm}^{-3}$ )
- **radTemperature** (*float* or *~numpy.ndarray*, optional) – Radiation black-body temperature (in Kelvin)
- **rStar** (*float* or *~numpy.ndarray*, optional) – Distance from the center of the star (in stellar radii)
- **abundance** (*float* or *str*, optional) – Elemental abundance relative to Hydrogen or name of CHIANTI abundance file to use, without the ‘.abund’ suffix, e.g. ‘sun\_photospheric\_1998\_grevesse’.
- **setup** (*bool* or *str*, optional) – If True, run ion setup function. Otherwise, provide a limited number of attributes of the selected ion
- **em** (*float* or *~numpy.ndarray*, optional) – Emission Measure, for the line-of-sight emission measure ( $\int n_e n_H dl$ ) ( $\text{cm}^{-5}$ ), for the volumetric emission measure  $\int n_e n_H dV$  ( $\text{cm}^{-3}$ ).

### Variables

- **IonStr** (*str*) – Name of element plus ion, e.g. *fe\_12* for *Fe XII*
- **Z** (*int*) – the nuclear charge, 26 for *fe\_12*.
- **Ion** (*int*) – the ionization stage, 12 for *fe\_12*.
- **Dielectronic** (*bool*) – true if the ion is a ‘dielectronic’ ion where the levels are populated by dielectronic recombination.
- **Spectroscopic** (*str*) – the spectroscopic notation for the ion, such as *Fe XII* for *fe\_12*.
- **Filename** (*str*) – the complete name of the file *generic* filename in the CHIANTI database, such as *\$XUVTOP/fe/fe\_12/fe\_12*.
- **Ip** (*float*) – the ionization potential of the ion
- **FIP** (*float*) – the first ionization potential of the element

- **Defaults** (*dict*) – these are specified by the software unless a *chiantirc* file is found in '\$HOME/.chianti':

## Notes

The keyword arguments temperature, eDensity, radTemperature, rStar, em must all be either a float or have the same dimension as the rest if specified as lists, tuples or arrays.

The *Defaults* dict should have the following keys:

- *abundfile*, the elemental abundance file, unless specified in *chiantirc* this defaults to *sun\_photospheric\_1998\_grevesse*.
- *ioneqfile*, the ionization equilibrium file name. Unless specified in 'chiantirc' this is defaults to *chianti*. Other choices are available in \$XUVTOP/ioneq
- *wavelength*, the units of wavelength (Angstroms, nm, or kev), unless specified in the 'chiantirc' this is defaults to 'angstrom'.
- *flux*, specified whether the line intensities are give in energy or photon fluxes, unless specified in the 'chiantirc' this is defaults to *energy*.
- *gui*, specifies whether to use gui selection widgets (True) or to make selections on the command line (False). Unless specified in the 'chiantirc' this is defaults to *False*.

**boundBoundLoss** (*allLines=1*)

Calculate the summed radiative loss rate for all spectral lines of the specified ion.

### Parameters

- **allLines** (*bool*) – If True, include losses from both observed and unobserved lines. If False, only include losses from observed lines.
- **includes elemental abundance and ionization fraction.**

### Returns

- *creates the attribute*
- **BoundBoundLoss** (*dict* with the keys below.) – *rate* : the radiative loss rate ( $\text{erg cm}^{-3} \text{s}^{-1}$ ) per unit emission measure.  
*temperature* : (K).  
*eDensity* : electron density ( $\text{cm}^{-3}$ )

**diCross** (*energy=None, verbose=False*)

Calculate the direct ionization cross section ( $\text{cm}^2$ ) as a function of the incident electron energy in eV, puts values into attribute DiCross

### Parameters

- **energy** (*array-like*) – incident electron energy in eV
- **verbose** (*bool, int*) – with verbose set to True, printing is enabled

**Variables DiCross** (*dict*) – keys: energy, cross

**diRate** ()

Calculate the direct ionization rate coefficient as a function of temperature (K)

**drPopulate** (*popCorrect=1, verbose=0*)

Calculate level populations for specified ion. possible keyword arguments include temperature, eDensity, pDensity, radTemperature and rStar different from method populate() in that it includes the dielectronic

recombination from all levels specified by the .auto file - consequently, it also calculates the populations of the higher ionization stage

**drRate ()**

Provide the dielectronic recombination rate coefficient as a function of temperature (K).

**drRateLvl (verbose=0)**

to calculate the level resolved dielectronic rate from the higher ionization stage to the ion of interest rates are determined from autoionizing A-values the dictionary self.DrRateLvl contains rate = the dielectronic rate into an autoionizing level effRate = the dielectronic rate into an autoionizing level multiplied by the branching ratio for a stabilizing transition totalRate = the sum of all the effRates

**eaCross (energy=None, verbose=False)**

Provide the excitation-autoionization cross section.

Energy is given in eV.

**eaDescale ()**

Calculates the effective collision strengths (upsilon) for excitation-autoionization as a function of temperature.

**eaRate ()**

Calculate the excitation-autoionization rate coefficient.

**emiss (allLines=True)**

Calculate the emissivities for lines of the specified ion.

units: ergs s<sup>-1</sup> str<sup>-1</sup>

Does not include elemental abundance or ionization fraction

Wavelengths are sorted

set allLines = True to include unidentified lines

**emissList (index=-1, wvlRange=None, wvlRanges=None, top=10, relative=0, outFile=0)**

List the emissivities.

wvlRange, a 2 element tuple, list or array determines the wavelength range

Top specifies to plot only the top strongest lines, default = 10

normalize = 1 specifies whether to normalize to strongest line, default = 0

**emissPlot (index=-1, wvlRange=None, top=10, linLog='lin', relative=0, verbose=0, plotFile=0, saveFile=0)**

Plot the emissivities.

wvlRange, a 2 element tuple, list or array determines the wavelength range

Top specifies to plot only the top strongest lines, default = 10

linLog specifies a linear or log plot, want either lin or log, default = lin

normalize = 1 specifies whether to normalize to strongest line, default = 0

**emissRatio (wvlRange=None, wvlRanges=None, top=10)**

Plot the ratio of 2 lines or sums of lines. Shown as a function of density and/or temperature. For a single wavelength range, set wvlRange = [wMin, wMax] For multiple wavelength ranges, set wvlRanges = [[wMin1,wMax1],[wMin2,wMax2], ...] A plot of relative emissivities is shown and then a dialog appears for the user to choose a set of lines.

**gofnt (wvlRange=0, top=10, verbose=0, plot=True)**

Calculate the 'so-called' G(T) function.

Given as a function of both temperature and eDensity.



Only the top( set by 'top') brightest lines are plotted. the G(T) function is returned in a dictionary self.Gofnt

**intensity** (*allLines=1, verbose=0*)

Calculate the intensities for lines of the specified ion.

units:  $\text{ergs cm}^{-3} \text{ s}^{-1} \text{ str}^{-1}$

includes elemental abundance and ionization fraction.

the emission measure 'em' is included if specified

**intensityRatioInterpolate** (*data, scale='lin', plot=0, verbose=0*)

to take a set of data and interpolate against the IntensityRatio the scale can be one of 'lin'/'linear' [default], 'loglog', 'logx', 'logy',

**ionizCross** (*energy=None*)

Provides the total ionization cross section.

## Notes

uses *diCross* and *eaCross*.

**ionizRate** ()

Provides the total ionization rate.

Calls *diRate* and *eaRate*.

**p2eRatio** ()

Calculates the proton density to electron density ratio using Eq. 7 of<sup>1</sup>.

## Notes

Uses the abundance and ionization equilibrium.

## References

**popPlot** (*top=10, levels=[], scale=0, plotFile=0, outFile=0, pub=0, addTitle=None*)

Plots populations vs temperature or eDensity.

top specifies the number of the most highly populated levels to plot (the default)

or can set levels to an array such as a list to set the desired levels to plot

if scale is set, then the population, if plotted vs. density, is divided by density - only useful if plotting level populations vs density

if pub is set, the want publication plots (bw, lw=2).

**populate** (*popCorrect=1, verbose=0*)

Calculate level populations for specified ion. possible keyword arguments include temperature, eDensity, pDensity, radTemperature and rStar populate assumes that all of the population in the higher ionization stages exists only in the ground level use *drPopulate*() for cases where the population of various levels in the higher ionization stage figure into the calculation

**recombRate** ()

Provides the total recombination rate coefficient.

Calls *drRate* and *rrRate*

---

<sup>1</sup> Young, P. R. et al., 2003, ApJS, 144, 135

**rrRate** ()

Provide the radiative recombination rate coefficient as a function of temperature (K).

**rrlvlDescale** (*verbose=1*)

Interpolate and extrapolate rrlvl rates. Used in level population calculations.

**setup** (*alternate\_dir=None, verbose=False*)

Setup various CHIANTI files for the ion including .wgfa, .elvlc, .scups, .psplups, .reclvl, .cilvl, and others.

#### Parameters

- **alternate\_dir** (*str*) – directory cotaining the necessary files for a ChiantiPy ion; use to setup an ion with files not in the current CHIANTI directory
- **verbose** (*bool*)

#### Notes

If ion is initiated with *setup=False*, call this method to do the setup at a later point.

**setupIonrec** (*alternate\_dir=None, verbose=False*)

Setup method for ion recombination and ionization rates.

#### Notes

Allows a bare-bones ion object to be setup up with just the ionization and recombination rates. For ions without a complete set of files - one that is not in the MasterList.

**spectrum** (*wavelength, filter=( $<function\ gaussianR>$ , 1000.0), label=0, allLines=1*)

Calculates the line emission spectrum for the specified ion.

Convolve the results of intensity to make them look like an observed spectrum the default filter is the gaussianR filter with a resolving power of 1000. Other choices include `chianti.filters.box` and `chianti.filters.gaussian`. When using the box filter, the width should equal the wavelength interval to keep the units of the continuum and line spectrum the same.

includes ionization equilibrium and elemental abundances

can be called multiple times to use different filters and widths uses label to keep the separate applications of spectrum sorted by the label for example, do `.spectrum( ... label = 'test1')` and `.spectrum( ... label = 'test2')` then will get `self.Spectrum.keys() = test1, test2` and `self.Spectrum['test1'] = {'intensity':aspectrum, 'wvl':wavelength, 'filter':useFilter.__name__, 'filterWidth':useFactor}`

#### Notes

`scipy.ndimage.filters` also includes a range of filters.

**twoPhoton** (*wvl, verbose=False*)

to calculate the two-photon continuum - only for hydrogen- and helium-like ions includes the elemental abundance and the ionization equilibrium includes the emission measure if specified

**twoPhotonEmiss** (*wvl*)

To calculate the two-photon continuum rate coefficient - only for hydrogen- and helium-like ions

**twoPhotonLoss** ()

to calculate the two-photon energy loss rate - only for hydrogen- and helium-like ions includes the elemental abundance and the ionization equilibrium does not include the emission measure

**upsilonDescale** (*prot=0*)

Provides the temperatures and effective collision strengths (upsilons) set prot for proton rates otherwise, ce will be set for electron collision rates uses the new format “scups” files

## ChiantiPy.core.Ioneq module

Ionization equilibrium class

**class** ChiantiPy.core.Ioneq.ioneq (*el\_or\_z*)

Bases: `object`

Calculation ion fractions as a function of temperature assuming ionization equilibrium.

**Parameters** *el\_or\_z* (*int* or *str*) – Atomic number or symbol

---

**Note:** When either loading or calculating a set of ion fractions, the temperature and ion fractions are returned to the *Temperature* and *Ioneq* attributes, respectively.

---

**calculate** (*temperature*)

Calculate ion fractions for given temperature array using the total ionization and recombination rates.

**load** (*ioneqName=None*)

Read temperature and ion fractions from a CHIANTI “.ioneq” file.

**plot** (*stages=0, tRange=0, yr=0, oplot=False, label=1, title=1, bw=False, semilogx=0, verbose=0*)

Plots the ionization equilibria.

`self.plot(tRange=None, yr=None, oplot=False)` *stages* = sequence of ions to be plotted, neutral == 1, fully stripped == Z+1 *tRange* = temperature range, *yr* = ion fraction range

for overplotting: *oplot* = “ioneqfilename” such as ‘mazzotta’ or if *oplot*=True or *oplot*=1 and a widget will come up so that a file can be selected. *bw*, if True, the plot is made in black and white

**plotRatio** (*stageN, stageD, tRange=0, yr=0, label=1, title=1, bw=0, semilogx=1, verbose=0*)

Plots the ratio of the ionization equilibria of two stages of a given element

`self.plotRatio(stageN, stageD)` *stages* = sequence of ions to be plotted, neutral == 1, fully stripped == Z+1 *stageN* = numerator *stageD* = denominator *tRange* = temperature range, *yr* = ion fraction range

## ChiantiPy.core.IpyMspectrum module

ChiantiPy.core.IpyMspectrum.doAll (*inpt*)

to process ff, fb and line inputs

**class** ChiantiPy.core.IpyMspectrum.ipymspectrum (*temperature, eDensity, wavelength, filter=(<function gaussianR>, 1000.0), label=None, elementList=None, ionList=None, minAbund=None, keepIons=0, doLines=1, doContinuum=1, allLines=1, em=None, abundance=None, verbose=0, timeout=0.1*)

Bases: `ChiantiPy.base._IonTrails.ionTrails, ChiantiPy.base._SpecTrails.specTrails`

this is the multiprocessing version of spectrum for using inside an IPython Qtconsole or notebook.

be for creating an instance, it is necessary to type something like the following into a console

> ipcluster start -n=3

this is the way to invoke things under the IPython 6 notation

Calculate the emission spectrum as a function of temperature and density.

temperature and density can be arrays but, unless the size of either is one (1), the two must have the same size

the returned spectrum will be convolved with a filter of the specified width on the specified wavelength array

the default filter is gaussianR with a resolving power of 100. Other filters, such as gaussian, box and lorentz, are available in ChiantiPy.filters. When using the box filter, the width should equal the wavelength interval to keep the units of the continuum and line spectrum the same.

A selection of elements can be made with elementList a list containing the names of elements that are desired to be included, e.g., ['fe','ni']

A selection of ions can be made with ionList containing the names of the desired lines in Chianti notation, i.e. C VI = c\_6

Both elementList and ionList can not be specified at the same time

a minimum abundance can be specified so that the calculation can be speeded up by excluding elements with a low abundance. With solar photospheric abundances -

setting minAbund = 1.e-4 will include H, He, C, O, Ne setting minAbund = 2.e-5 adds N, Mg, Si, S, Fe setting minAbund = 1.e-6 adds Na, Al, Ar, Ca, Ni

Setting em will multiply the spectrum at each temperature by the value of em.

em [for emission measure], can be a float or an array of the same length as the temperature/density. allLines = 1 will include lines with either theoretical or observed wavelengths. allLines=0 will include only those lines with observed wavelengths

proc = the number of processors to use timeout - a small but non-zero value seems to be necessary

## ChiantiPy.core.Mspectrum module

```
class ChiantiPy.core.Mspectrum.mspectrum(temperature, eDensity, wavelength, filter=(<function gaussianR>, 1000.0), label=0, elementList=None, ionList=None, minAbund=None, keepIons=0, abundance=None, doLines=1, doContinuum=1, allLines=1, em=None, proc=3, verbose=0, timeout=0.1)

Bases: ChiantiPy.base._IonTrails.ionTrails, ChiantiPy.base._SpecTrails.specTrails
```

this is the multiprocessing version of spectrum set proc to the desired number of processors, default=3

Calculate the emission spectrum as a function of temperature and density.

temperature and density can be arrays but, unless the size of either is one (1), the two must have the same size

the returned spectrum will be convolved with a filter of the specified width on the specified wavelength array

the default filter is gaussianR with a resolving power of 100. Other filters, such as gaussian, box and lorentz, are available in chianti.filters. When using the box filter, the width should equal the wavelength interval to keep the units of the continuum and line spectrum the same.

A selection of elements can be made with elementList a list containing the names of elements that are desired to be included, e.g., ['fe','ni']

A selection of ions can be made with `ionList` containing the names of the desired lines in Chianti notation, i.e. C VI = c\_6

Both `elementList` and `ionList` can not be specified at the same time

a minimum abundance can be specified so that the calculation can be speeded up by excluding elements with a low abundance. With solar photospheric abundances -

setting `minAbund = 1.e-4` will include H, He, C, O, Ne setting `minAbund = 2.e-5` adds N, Mg, Si, S, Fe setting `minAbund = 1.e-6` adds Na, Al, Ar, Ca, Ni

Setting `em` will multiply the spectrum at each temperature by the value of `em`.

`em` [for emission measure], can be a float or an array of the same length as the temperature/density. `allLines = 1` will include lines with either theoretical or observed wavelengths. `allLines=0` will include only those lines with observed wavelengths

`proc` = the number of processors to use `timeout` - a small but non-zero value seems to be necessary

## ChiantiPy.core.RadLoss module

```
class ChiantiPy.core.RadLoss.radLoss(temperature, eDensity, elementList=0, ionList=0,
                                     minAbund=0, doContinuum=1, abundance=None,
                                     verbose=0, allLines=1, keepIons=0)
```

Bases: `ChiantiPy.base._IonTrails.ionTrails`, `ChiantiPy.base._SpecTrails.specTrails`

Calculate the emission spectrum as a function of temperature and density.

includes elemental abundances or ionization equilibria

temperature and density can be arrays but, unless the size of either is one (1), the two must have the same size

the returned spectrum will be convolved with a filter of the specified width on the specified wavelength array

the default filter is `gaussianR` with a resolving power of 1000. Other filters, such as `gaussian`, `box` and `lorentz`, are available in `ChiantiPy.filters`. When using the `box` filter, the width should equal the wavelength interval to keep the units of the continuum and line spectrum the same.

A selection of ions can be made with `ionList` containing the names of the desired lines in Chianti notation, i.e. C VI = c\_6

a minimum abundance can be specified so that the calculation can be speeded up by excluding elements with a low abundance. With solar photospheric abundances -

setting `minAbund = 1.e-4` will include H, He, C, O, Ne setting `minAbund = 2.e-5` adds N, Mg, Si, S, Fe setting `minAbund = 1.e-6` adds Na, Al, Ar, Ca, Ni

Setting `em` will multiply the spectrum at each temperature by the value of `em`.

`em` [for emission measure], can be a float or an array of the same length as the temperature/density.

**abundance:** to select a particular set of abundances, set **abundance** to the name of a CHIANTI abundance file, without the '.abund' suffix, e.g. 'sun\_photospheric\_1998\_grevesse' If set to a blank (''), a gui selection menu will popup and allow the selection of an set of abundances

```
radLossPlot (title=0)
```

to plot the radiative losses vs temperature

## ChiantiPy.core.Spectrum module

**class** ChiantiPy.core.Spectrum.**bunch**(*temperature, eDensity, wvlRange, elementList=None, ionList=None, minAbund=None, keepIons=0, em=None, abundance=None, verbose=0, allLines=1*)  
Bases: ChiantiPy.base.\_IonTrails.ionTrails, ChiantiPy.base.\_SpecTrails.specTrails

Calculate the emission line spectrum as a function of temperature and density.

‘bunch’ is very similar to ‘spectrum’ except that continuum is not calculated and the spectrum is not convolved over a filter. However, this can be done with the inherited convolve method

one of the convenient things is that all of the instantiated ion classes, determined through such keywords as ‘elementList’, ‘ionList’, and ‘minAbund’ are kept in a dictionary self.IonInstances where self.IonInstances[‘mg\_7’] is the class instance of ChiantiPy.core.ion for ‘mg\_7’. All its methods and attributes are available.

includes elemental abundances and ionization equilibria

the set of abundances, a file in \$XUVTOP/abundance, can be set with the keyword argument ‘abundanceName’

temperature and density can be arrays but, unless the size of either is one (1), the two must have the same size

Inherited methods include ‘intensityList’, ‘intensityRatio’ (between lines of different ions), and ‘intensityRatioSave’ and ‘convolve’.

A selection of elements can be make with elementList a list containing the names of elements that are desired to be included, e.g., [‘fe’, ‘ni’]

A selection of ions can be make with ionList containing the names of the desired lines in Chianti notation, i.e. C VI = c\_6

Both elementList and ionList can not be specified at the same time

a minimum abundance can be specified so that the calculation can be speeded up by excluding elements with a low abundance. With solar photospheric abundances -

setting minAbund = 1.e-4 will include H, He, C, O, Ne setting minAbund = 2.e-5 adds N, Mg, Si, S, Fe setting minAbund = 1.e-6 adds Na, Al, Ar, Ca, Ni

At least one of elementList, ionList, or minAbund must be set in order for ‘bunch’ to include any ions.

Setting em will multiply the spectrum at each temperature by the value of em.

em [for emission measure], can be a float or an array of the same length as the temperature/density

**class** ChiantiPy.core.Spectrum.**spectrum**(*temperature, eDensity, wavelength, filter=(<function gaussianR>, 1000.0), label=None, elementList=None, ionList=None, minAbund=None, doLines=1, doContinuum=1, em=None, keepIons=0, abundance=None, verbose=0, allLines=1*)  
Bases: ChiantiPy.base.\_IonTrails.ionTrails, ChiantiPy.base.\_SpecTrails.specTrails

Calculate the emission spectrum as a function of temperature and density.

one of the convenient things is that all of the instantiated ion classes, determined through such keywords as ‘elementList’, ‘ionList’, and ‘minAbund’ are kept in a dictionary self.IonInstances where self.IonInstances[‘mg\_7’] is the class instance of ChiantiPy.core.ion for ‘mg\_7’. All its methods and attributes are available.

includes elemental abundances and ionization equilibria

the set of abundances, a file in \$XUVTOP/abundance, can be set with the keyword argument ‘abundanceName’

temperature and density can be arrays but, unless the size of either is unity (1), the two must have the same size

the returned spectrum will be convolved with a filter of the specified width on the specified wavelength array  
the default filter is gaussianR with a resolving power of 1000. Other filters, such as gaussian, box and lorentz, are available in ChiantiPy.tools.filters. When using the box filter, the width should equal the wavelength interval to keep the units of the continuum and line spectrum the same.

Inherited methods include 'intensityList', 'intensityRatio' (between lines of different ions), 'intensityRatioSave' and 'convolve'

A selection of elements can be make with elementList a list containing the names of elements that are desired to be included, e.g., ['fe','ni']

A selection of ions can be make with ionList containing the names of the desired lines in CHIANTI notation, i.e. C VI = c\_6

Both elementList and ionList can not be specified at the same time

a minimum abundance can be specified so that the calculation can be speeded up by excluding elements with a low abundance. The default of minAbund is 1.e-6

It is necessary to specify at least an elementList, an ionList, or a minAbund to select any ions for a spectrum calculation

With solar photospheric abundances -

setting minAbund = 1.e-4 will include H, He, C, O, Ne setting minAbund = 2.e-5 adds N, Mg, Si, S, Fe setting minAbund = 1.e-6 adds Na, Al, Ar, Ca, Ni

Setting doLines = 0 will skip the calculation of spectral lines. Setting doContinuum =0 will skip the continuum calculation.

Setting em [for emission measure] will multiply the spectrum at each temperature by the value of em.

em [for emission measure] can be a float or an array of the same length as the temperature/density

keepIons: set this to keep the ion instances that have been calculated in a dictionary self.IonInstances with the keywords being the CHIANTI-style ion names

abundance: to select a particular set of abundances, set abundance to the name of a CHIANTI abundance file, without the '.abund' suffix, e.g. 'sun\_photospheric\_1998\_grevesse'

If set to a blank (''), a gui selection menu will popup and allow the selection of an set of abundances

## Module contents

chianti.core - contains the main classes for ChiantiPy users.

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## ChiantiPy.fortranformat package

### Submodules

#### ChiantiPy.fortranformat.FortranRecordReader module

```
class ChiantiPy.fortranformat.FortranRecordReader.FortranRecordReader (format)  
    Bases: object
```

Generate a reader object for FORTRAN format strings

Typical use case ...

```
>>> header_line = FortranRecordReader('(A15, A15, A15)')
>>> header_line.read('          x          y          z')
['          x', '          y', '          z']
>>> line = FortranRecordReader('(3F15.3)')
>>> line.read('          1.000          0.000          0.500')
[1.0, 0.0, 0.5]
>>> line.read('          1.100          0.100          0.600')
[1.1, 0.1, 0.6]
```

Note: it is best to create a new object for each format, changing the format causes the parser to reevaluate the format string which is costly in terms of performance

**property format**

**get\_format** ()

**match** (*record*)

**read** (*record*)

Pass a string representing a FORTRAN record to obtain the relevant values

**set\_format** (*format*)

## ChiantiPy.fortranformat.FortranRecordWriter module

**class** ChiantiPy.fortranformat.FortranRecordWriter.**FortranRecordWriter** (*format*)  
Bases: `object`

Generate a writer object for FORTRAN format strings

Typical use case ...

```
>>> header_line = FortranRecordWriter('(A15, A15, A15)')
>>> header_line.write(['x', 'y', 'z'])
'          x          y          z'
>>> line = FortranRecordWriter('(3F15.3)')
>>> line.write([1.0, 0.0, 0.5])
'          1.000          0.000          0.500'
>>> line.write([1.1, 0.1, 0.6])
'          1.100          0.100          0.600'
```

Note: it is best to create a new object for each format, changing the format causes the parser to reevaluate the format string which is costly in terms of performance

**property format**

**get\_format** ()

**set\_format** (*format*)

**write** (*values*)

Pass a list of values corresponding to the FORTRAN format specified to generate a string



## ChiantiPy.fortranformat.config module

ChiantiPy.fortranformat.config.**reset**()

## Module contents

### ChiantiPyGui package

### Submodules

## ChiantiPy.model.Maker module

classes and methods to analyze observations of astrophysical spectra

ChiantiPy.model.Maker.**doDemGofntQ**(*inQueue, outQueue*)  
helper for multiprocessing with maker.mgofnt()

ChiantiPy.model.Maker.**makeMatchPkl**(*specData, temp, dens, wghtFactor=0.25, abundanceName=None, minAbund=1e-06, useMgofnt=1, verbose=0*)

input a data dictionary and instantiate a dem class, and run mgofnt and then make a pickle file to use multiprocessing, this needs to be run in an ipython console

Parameters

**specData** [dict] the observed line intensities, wavelengths ...

**class** ChiantiPy.model.Maker.**maker**(*specData, wghtFactor=0.0, ionList=False, allLines=True, abundanceName=None, minAbund=1e-06, verbose=False*)  
Bases: *ChiantiPy.base.\_IonTrails.ionTrails*

a class matching observed lines to lines in the CHIANTI database

**argCheck**(*temperature=None, eDensity=None, pDensity='default', verbose=0*)  
to check the compatibility of the three arguments and put them into numpy arrays of atleast\_1d

**diffPrintChi**(*dir='.', filename='diffPrintChi.txt'*)  
calculates the weighted and straight differences between observed and predicted prints the values saves the as a dictionary self.Diff to be used together with a prior brute-force chi-squared minimization approach

**diffPrintMc**(*dir='.', filename='diffPrintMc.txt'*)  
calculates the weighted and straight differences between observed and predicted prints the values saves to a file from a PyMC run

**emFitPlot**()  
to plot the emission measures derived from search over temperature

**emNtSetIndices**(*indices, verbose=0*)  
to set the indices of the N temperature EM distribution

**emPlot**(*vs='T'*)  
to plot line intensities divided by gofnt

**emSet**(*value*)  
sets the EM values for a N temperature EM distribution

**emSetIndices**(*indices, verbose=0*)  
to set the indices of the N temperature/density EM distribution

**findMinMaxIndices** (*verbose=0*)

to find the minimum and maximum indices where all match['intensitySum'] are greater than 0

**fit1t** (*initialValue, maxfev=0*)

calls leastsq to fit the 1t (single temperature) model

**fitFunc1t** (*em*)

the fitting function for the isothermal model to be called by leastsq

**Parameters** **em** (*number*) – the log10 value of the emission measure

**Returns** **weighted chisquared**

**Return type** **float**

**fitFuncNt** (*value*)

the fitting function for the 1 (single temperature) temperature model to be called by leastsq

**fitNt** (*initialValue, maxfev=0*)

calls leastsq to fit the 2d model

**getChisq** ()

return chisq

**getNormalizedChisq** ()

return normalized chisq

**getWeightedDiff** ()

to calculate the weighted difference of each of the intensities returns a 1D array

**gofnt** (*temperature, density, verbose=1*)

calculate the gofnt function for each of the matched lines do each ion only once

**mgofnt** (*temperature, density, proc=6, timeout=0.1, verbose=0*)

calculate the gofnt function for each of the matched lines this is the multiprocessing version do each ion only once

**predict** ()

to predict the intensities of the observed lines from an emission measure the emission measure is already specified as self.Em which is an np array

**predictPrint** (*minContribution=0.1, outfile=0, verbose=0*)

to predict the intensities of the observed lines from an emission measure the emission measure is already specified as self.Em which is an np array

**predictPrint1d** (*minContribution=0.1, outfile=0, verbose=0*)

to predict the intensities of the observed lines from an emission measure the emission measure is already specified as self.Em which is an np array

**search1dSpace** (*initialEm, indxlimits=None, verbose=0, log=0, maxfev=0*)

to conduct a brute force search over electron density for an isothermal-space and find the best fit to the em and density indxlimits give the range of indices to fit over can use self.MinIndex and self.MaxIndex+1  
initialEm = log value of the emission measure to begin the searching

**search1tEmSpace** (*verbose=0*)

to find the value of chisq as a function of Em with T = best-fit

## Module contents

classes and methods for modeling observed spectra.

## ChiantiPy.tests package

### Submodules

#### ChiantiPy.tests.setup\_package module

`ChiantiPy.tests.setup_package.get_package_data()`

## Module contents

This packages contains affiliated package tests.

## ChiantiPy.tools package

### Submodules

#### ChiantiPy.tools.archival module

Functions for reading pre-v8 CHIANTI files

`ChiantiPy.tools.archival.elvlcRead(ions, filename=0, verbose=0, useTh=0)`  
 read a chianti energy level file and returns {"lvl":lvl,"conf":conf,"term":term,"spin":spin,"l":l,"spd":spd,"j":j,"mult":mult,"ecm":ecm,"eryd":eryd,"ecmth":ecmth,"erydth":erydth,"ref":ref,"pretty":pretty, 'ionS':ions} if a energy value for ecm or eryd is zero(=unknown), the theoretical values (ecmth and erydth) are inserted

`ChiantiPy.tools.archival.elvlcWrite(info, outfile=None, addLvl=0)`  
 Write Chianti data to .elvlc file.

#### Parameters

- **info** (*dict*) – Information about the Chianti data to write. Should contain the following keys: ionS, the Chianti style name of the ion such as c\_4 conf, an integer denoting the configuration - not too essential term, a string showing the configuration spin, an integer of the spin of the state in LS coupling l, an integer of the angular momentum quantum number spd, an string for the alphabetic symbol of the angular momemtum, S, P, D, etc j, a floating point number, the total angular momentum ecm, the observed energy in inverse cm, if unknown, the value is 0. eryd, the observed energy in Rydbergs, if unknown, the value is 0. ecmth, the calculated energy from the scattering calculation, in inverse cm erydth, the calculated energy from the scattering calculation in Rydbergs ref, the references in the literature to the data in the input info
- **outfile** (*str*) – Output filename. ionS+'.elvlc' (in current directory) if None
- **addLvl** (*int*) – Add a constant value to the index of all levels

## Notes

For use with files created before elvlc format change in November 2012

See also:

`ChiantiPy.tools.io.elvlcWrite()` Write .elvlc file using the new format.

`ChiantiPy.tools.archival.wgfaRead(ions, filename=None, elvlcname=-1, total=False, verbose=False)`

Read CHIANTI data from a .wgfa file.

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*) – Custom filename, will override that specified by *ions*
- **elvlcname** (*str*) – If specified, the lsj term labels are returned in the *pretty1* and *pretty2* keys of *Wgfa*
- **total** (*bool*) – Return the level 2 avalue data in *Wgfa*
- **verbose** (*bool*)

**Returns** *Wgfa* – Information read from the .wgfa file. The dictionary structure is {"lv11", "lv12", "wvl", "gf", "avalue", "ref", "ionS", "filename"}

**Return type** *dict*

## Notes

This is text-wise not different than the v8 version except that it uses the archival `elvlcRead` in `~ChiantiPy.tools.archival` though this has now been commented out. Can this routine be removed? Should the `elvlcRead` routine be uncommented?

See also:

`ChiantiPy.tools.io.wgfaRead()` Read .wgfa file with the new format.

## ChiantiPy.tools.constants module

A set of physical constants, in (mostly) cgs unit system. Most are from<sup>11</sup>.

## References

### Notes

Many of these can be replaced by the `~astropy.constants` module. Elemental symbols can be removed in favor of the `periodictable` module. Spectroscopic roman numerals can be removed in favor of `roman` module. The Gauss-Laguerre weights can be calculated by `~numpy.polynomial.laguerre.laggauss`.

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<sup>11</sup> NIST Reference on Constants, Units and Uncertainty ([link](#))

## ChiantiPy.tools.data module

Module for collecting various top-level Chianti data

Descriptions for *keywordArgs*:

- *temperature* : temperature (in K)
- *eDensity* : electron density (in  $\text{cm}^{-3}$ )
- *hDensity* : hydrogen density (in  $\text{cm}^{-3}$ )
- *pDensity* : proton density (in  $\text{cm}^{-3}$ )
- *radTemperature* : radiation temperature of central source (in K)
- *rStar* : distance of the plasma from the source (in units of the source's radius)
- *distance* : distance from the central source

### Parameters

- **XUVTOP** (*str*) – the root of the CHIANTI database
- **Defaults** (*dict*) – default values used by ChiantiPy, can also be set by an optional HOME/.chianti/chiantirc file
- **Ip** (*np.ndarray*) – the ionization potentials for all ionization stages up to Zn, in eV
- **MasterList** (*list*) – the CHIANTI style names of all ions in the CHIANTI database
- **IoneqAll** (*dict*) – a dict containing the ionization equilibrium values for the default ionization file
- **ChiantiVersion** (*str*) – the version of the CHIANTI database
- **AbundanceDefault** (*dict*) – the elemental abundances in the default abundance file
- **AbundanceList** (*list*) – the names of all abundance files included in the CHIANTI database
- **GrndLevels** (*list*) – the number of levels that should be considered in an ionization calculation

## ChiantiPy.tools.filters module

Line profile filters for creating synthetic spectra.

ChiantiPy.tools.filters.**boxcar** (*wvl, wvl0, factor=None*)

Box-car filter

### Parameters

- **wvl** (*~numpy.ndarray*) – Wavelength array
- **wvl0** (*~numpy.float64*) – Wavelength filter should be centered on.
- **factor** (*~numpy.float64*) – Full width of the box-car filter

ChiantiPy.tools.filters.**gaussian** (*wvl, wvl0, factor=1.0*)

A gaussian filter

### Parameters

- **wvl** (*~numpy.ndarray*) – Wavelength array
- **wvl0** (*~numpy.float64*) – Wavelength filter should be centered on.
- **factor** (*~numpy.float64*) – Gaussian width

- **integrated value is unity**

`ChiantiPy.tools.filters.gaussianR(wvl, wvl0, factor=1000.0)`

A gaussian filter where the gaussian width is given by `wvl0/factor`.

#### Parameters

- **wvl** (`~numpy.ndarray`) – Wavelength array
- **wvl0** (`~numpy.float64`) – Wavelength filter should be centered on.
- **factor** (`~numpy.float64`) – Resolving power

`ChiantiPy.tools.filters.lorentz(wvl, wvl0, factor=1.0)`

Lorentz profile filter with the exception that all factors are in wavelength units rather than frequency as the lorentz profile is usually defined.

#### Parameters

- **wvl** (`~numpy.ndarray`) – Wavelength array
- **wvl0** (`~numpy.float64`) – Wavelength filter should be centered on.
- **factor** (`~numpy.float64`) – Value of the so-called constant gamma
- **integrated value is unity**
- **the FWHM is 2\*gamma**
- **.. math::** –  $L = \frac{1}{\pi \gamma} \frac{\gamma^2}{(\lambda - \lambda_0)^2 + \gamma^2}$

`ChiantiPy.tools.filters.moffat(wvl, wvl0, factor=2.5)`

Moffat profile with parameters suited to Chandra Letg spectra

#### Parameters

- **wvl** (`~numpy.ndarray`) – Wavelength array
- **wvl0** (`~numpy.float64`) – Wavelength the filter is centered on.
- **factor** (`~numpy.float64`) – Resolving power (TODO: correct description)

`ChiantiPy.tools.filters.voigt(wvl, wvl0, factor=(0.5, 1.0))`

pseudo-Voigt filter the sum of a Gaussian and a Lorentzian

#### Parameters

- **wvl** (`~numpy.ndarray`) – Wavelength array
- **wvl0** (`~numpy.float64`) – Wavelength the filter is centered on.
- **factor** (`array-type`) – contains the following 2 parameters
- **A** (`~numpy.float64`) – relative size of gaussian and lorentz components must be between 0. and 1. but this is not currently checked
- **sigma** (`~numpy.float64`) – the gaussian width of the gaussian profile (the standard deviation) also creates the lorentz component with the same fwhm

## ChiantiPy.tools.io module

Reading and writing functions

`ChiantiPy.tools.io.abundanceRead(abundancename="")`

Read abundance file *abundancename* and return the abundance values relative to hydrogen

`ChiantiPy.tools.io.autoRead(ions, filename=None, total=True, verbose=False)`

Read CHIANTI autoionization rates from a .auto file.

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*) – Custom filename, will override that specified by *ions*
- **elvlcname** (*str*) – If specified, the lsj term labels are returned in the ‘pretty1’ and ‘pretty2’ keys of ‘Wgfa’ dict
- **total** (*bool*) – Return the summed level 2 autoionization rates in ‘Auto’
- **verbose** (*bool*)

**Returns** **Auto** – Information read from the .wgfa file. The dictionary structure is {“lv1”, “lv2”, “avalue”, “pretty1”, “pretty2”, “ref”, “ionS”, “filename”}

**Return type** *dict*

`ChiantiPy.tools.io.autoWrite(info, outfile=None, minBranch=None)`

Write data to a CHIANTI .wgfa file

### Parameters

- **info** (*dict*) – Should contain the following: ionS, the Chianti style name of the ion such as c\_4 for C IV, lv1, the lower level, the ground level is 1, lv2, the upper level, wvl, the wavelength (in Angstroms), avalue, the autoionization rate, pretty1, descriptive text of the lower level (optional), pretty2, descriptive text of the upper level (optional), ref, reference text, a list of strings
- **outfile** (*str*)
- **minBranch** (*~numpy.float64*) – The transition must have a branching ratio greater than the specified to be written to the file

`ChiantiPy.tools.io.cireclvlRead(ions, filename=None, filetype='civl')`

Read Chianti cilvl, reclvl, or rrlvl files and return data

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*
- **filetype** (*str*) – {‘cilvl’, ‘reclvl’, ‘rrlvl’} Type of the file to read

`ChiantiPy.tools.io.convertName(name)`

Convert ion name string to Z and Ion and other interesting info

**Parameters** **name** (*str*) – a generic name of an ion in the CHIANTI database, such as fe\_14 for Fe XIV

## Notes

A duplicate of the routine in *ChiantiPy.tools.util* but needed by masterList Info TODO: Put in separate module to avoid multiple copies

`ChiantiPy.tools.io.defaultsRead(verbose=False)`

Read in configuration from .chiantirc file or set defaults if one is not found.

`ChiantiPy.tools.io.diRead(ions,filename=None)`

Read chianti direct ionization .params files and return data.

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*

`ChiantiPy.tools.io.drRead(ions,filename=None)`

Read CHIANTI dielectronic recombination .drparams files if filename is set, then reads that file

`ChiantiPy.tools.io.eaRead(ions,filename=None)`

Read a CHIANTI excitation-autoionization file and calculate the EA ionization rate data derived from splupsRead.

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*

`ChiantiPy.tools.io.elvlcRead(ions, filename=None, getExtended=False, verbose=False, useTh=True)`

Reads the new format elvlc files.

### Parameters

- **ions** (*str*) – Ion, e.g. ‘c\_5’ for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*
- **getExtended** (*bool*)
- **verbose** (*bool*)
- **useTh** (*bool*) – If True, the theoretical values (ecmth and erydth) are inserted when an energy value for ecm or eryd is zero(=unknown)

`ChiantiPy.tools.io.elvlcWrite(info, outfile=None, round=0, addLvl=0, includeRyd=False, includeEv=False)`

Write Chianti data to .elvlc file.

### Parameters

- **info** (*dict*) – Information about the Chianti data to write. Should contain the following keys: ionS, the Chianti style name of the ion such as c\_4 term, a string showing the configuration spin, an integer of the spin of the state in LS coupling l, an integer of the angular momentum quantum number spd, a string for the alphabetic symbol of the angular momentum, S, P, D, etc j, a floating point number, the total angular momentum ecm, the observed energy in inverse cm, if unknown, the value is 0. eryd, the observed energy in Rydbergs, if unknown, the value is 0. ecmth, the calculated energy from the scattering calculation, in inverse cm erydth, the calculated energy from the scattering calculation in Rydbergs ref, the references in the literature to the data in the input info
- **outfile** (*str*) – Output filename. ionS + ‘.elvlc’ (in current directory) if None



- **round** (*int*) – input to ‘np.round’ to round input values to maintain the correct number of significant figures
- **addLvl** (*int*) – Add a constant value to the index of all levels
- **includeRyd** (*bool*) – If True, write the Rydberg energies in the extended area, delimited by a comma
- **includeEv** (*bool*) – If True, write the energies in eV in the extended area, delimited by a comma

## Notes

For use with files created after elvlc format change in November 2012

See also:

`ChiantiPy.tools.archival.elvlcWrite()` Write .elvlc file using the old format.

`ChiantiPy.tools.io.fblvlRead(ions, filename=None, verbose=False)`

Read a Chianti energy level file for calculating the free-bound continuum

`ChiantiPy.tools.io.gffRead()`

Read the free-free gaunt factors of<sup>1</sup>.

## References

## Notes

This function reads the file and reverses the values of g2 and u

`ChiantiPy.tools.io.gffintRead()`

Read the integrated free-free gaunt factors of<sup>1</sup>.

`ChiantiPy.tools.io.grndLevelsRead()`

to read the grndLevels.dat file give the number of ground levels to sum over in populate and drPopulate

`ChiantiPy.tools.io.ioneqRead(ioneqName="", minIoneq=1e-20, verbose=False)`

Reads an ioneq file ionization equilibrium values less than minIoneq are returns as zeros :returns: {'ioneq-name', 'ioneqAll', 'ioneqTemperature', 'ioneqRef'} – Ionization equilibrium values and the reference to the literature :rtype: *dict*

`ChiantiPy.tools.io.ipRead(verbose=False)`

Reads the ionization potential file

**Returns** **ip** – Ionization potential (in eV)

**Return type** array-like

`ChiantiPy.tools.io.itohRead()`

Read in the free-free gaunt factors of<sup>2</sup>.

<sup>1</sup> Sutherland, R. S., 1998, MNRAS, 300, 321

<sup>2</sup> Itoh, N. et al., 2000, ApJS, 128, 125

## References

`ChiantiPy.tools.io.klgfbRead()`

Read CHIANTI files containing the free-bound gaunt factors for n=1-6 from<sup>13</sup>.

**Returns** {'pe', 'klgfb'} – Photon energy and the free-bound gaunt factors

**Return type** *dict*

## References

`ChiantiPy.tools.io.masterListInfo(force=False, verbose=False)`

Get information about ions in the CHIANTI masterlist.

**Returns** `masterListInfo` – {'wmin', 'wmax', 'tmin', 'tmax'} Minimum and maximum wavelengths in the wgfa file. Minimum and maximum temperatures for which the ionization balance is nonzero.

**Return type** *dict*

## Notes

This function speeds up multi-ion spectral calculations. The information is stored in a pickled file 'masterlist\_ions.pkl'. If the file is not found, one will be created.

`ChiantiPy.tools.io.masterListRead()`

Read a CHIANTI masterlist file.

**Returns** `masterlist` – All ions in Chianti database

**Return type** *list*

`ChiantiPy.tools.io.photoxRead(ions)`

Read CHIANTI photoionization .photox files

**Returns** {'lvl1', 'lvl2', 'energy', 'cross', 'ref'} – Energy (in Rydbergs) and cross section (in  $\text{cm}^{-2}$ )

**Return type** *dict*

## Notes

The photox files are not in any released version of the CHIANTI database.

`ChiantiPy.tools.io.rLossRead()`

to read the Mao 2017 rr loss parameters<sup>12</sup>

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<sup>13</sup> Karzas and Latter, 1961, ApJSS, 6, 167

<sup>12</sup> Mao J., Kaastra J., Badnell N.R., 2017 Astron. Astrophys. 599, A10

## References

`ChiantiPy.tools.io.rrRead(ions,filename=None)`

Read CHIANTI radiative recombination .rrparams files

**Returns** {'rrtype','params','ref'}

**Return type** *dict*

`ChiantiPy.tools.io.scupsRead(ions,filename=None,verbose=False)`

Read the new format v8 scups file containing the scaled temperature and upsilons from<sup>8</sup>.

### Parameters

- **ions** (*str*) – Ion, e.g. 'c\_5' for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*
- **verbose** (*bool*)

`ChiantiPy.tools.io.splomRead(ions,ea=False,filename=None)`

Read chianti .splom files

### Parameters

- **ions** (*str*) – Ion, e.g. 'c\_5' for C V
- **ea** (*bool*) – Read .easplom file
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*

**Returns** {'lvl1', 'lvl2', 'ttype', 'gf', 'deryd', 'c', 'splom', 'ref'}

**Return type** *dict*

## Notes

Still needed for ionization cross sections

`ChiantiPy.tools.io.splupsRead(ions,filename=None,filetype='splups')`

Read a CHIANTI .splups file

### Parameters

- **ions** (*str*) – Ion, e.g. 'c\_5' for C V
- **filename** (*str*, optional) – Custom filename, will override that specified by *ions*
- **filetype** (*str*, optional) – {*psplups*, *cisplups*, *splups*} Type of file to read

**Returns** {'lvl1', 'lvl2', 'ttype', 'gf', 'de', 'cups', 'bsplups', 'ref'}

**Return type** *dict*

`ChiantiPy.tools.io.trRead(ionS)`

read the files containing total recombination rates .trparams

`ChiantiPy.tools.io.twophotonHRead()`

Read the two-photon Einstein A values and distribution function for the H sequence.

**Returns** {'y0', 'z0', 'avalue', 'asum', 'psi0'}

**Return type** *dict*

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<sup>8</sup> Burgess, A. and Tully, J. A., 1992, A&A, 254, 436

`ChiantiPy.tools.io.twophotonHeRead()`

Read the two-photon Einstein A values and distribution function for the He sequence.

**Returns** {'y0', 'avalue', 'psi0'}

**Return type** *dict*

`ChiantiPy.tools.io.vernerRead()`

Reads the photoionization cross-section data from<sup>6</sup>.

**Returns** {'pqn','l','eth','e0','sig0','ya','p', 'yw'} – *pqn* is the principal quantum number, *l* is the subshell orbital quantum number, *eth* (in eV) is the subshell ionization threshold energy; *sig0*, *ya*, *p*, and *yw* are all fit parameters used in calculating the total photoionization cross-section.

**Return type** *dict*

## References

`ChiantiPy.tools.io.versionRead()`

Read the version number of the CHIANTI database

`ChiantiPy.tools.io.wgfaRead(ions,filename=None,elvlcname=0,total=False,verbose=False)`

Read CHIANTI data from a .wgfa file.

### Parameters

- **ions** (*str*) – Ion, e.g. 'c\_5' for C V
- **filename** (*str*) – Custom filename, will override that specified by *ions*
- **elvlcname** (*str*) – If specified, the lsj term labels are returned in the 'pretty1' and 'pretty2' keys of 'Wgfa' dict
- **total** (*bool*) – Return the summed level 2 avalue data in 'Wgfa'
- **verbose** (*bool*)

**Returns** **Wgfa** – Information read from the .wgfa file. The dictionary structure is {'lv11','lv12','wvl','gf','avalue','ref','ionS','filename'}

**Return type** *dict*

See also:

`ChiantiPy.tools.archival.wgfaRead()` Read .wgfa file with the old format.

`ChiantiPy.tools.io.wgfaWrite(info, outfile=None, minBranch=1e-05, rightDigits=4, maxLv11=None)`

Write data to a CHIANTI .wgfa file

### Parameters

- **info** (*dict*) – Should contain the following keys: ionS, the Chianti style name of the ion such as c\_4 for C IV, lv11, the lower level, the ground level is 1, lv12, the upper level, wvl, the wavelength (in Angstroms), gf, the weighted oscillator strength, avalue, the A value, pretty1, descriptive text of the lower level (optional), pretty2, descriptive text of the upper level (optional), ref, reference text, a list of strings
- **outfile** (*str*)

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<sup>6</sup> Verner & Yakovlev, 1995, A&AS, 109, 125

- **minBranch** (*~numpy.float64*) – The transition must have a branching ratio greater than the specified minBranch to be written to the file

`ChiantiPy.tools.io.zion2name(z, ion, dielectronic=False)`

Convert Z and ion to generic name, e.g. 26, 13 -> fe\_13

#### Parameters

- **z** (*int*) – the nuclear charge, for example 26 for Fe XIV
- **ion** (*int*) – the ion stage, for example, 14 for Fe XIV

#### Notes

A duplicate of the routine in *ChiantiPy.tools.util* but needed by masterList Info TODO: Put in separate module to avoid multiple copies

### ChiantiPy.tools.mputil module

Functions needed for standard Python multiprocessing module mspectrum

`ChiantiPy.tools.mputil.doFbQ(inQ, outQ)`

Multiprocessing helper for *ChiantiPy.core.continuum.freeBound*

#### Parameters

- **inQ** (*~multiprocessing.Queue*) – Ion free-bound emission jobs queued up by multiprocessing module
- **outQ** (*~multiprocessing.Queue*) – Finished free-bound emission jobs

`ChiantiPy.tools.mputil.doFfQ(inQ, outQ)`

Multiprocessing helper for *ChiantiPy.core.continuum.freeFree*

#### Parameters

- **inQ** (*~multiprocessing.Queue*) – Ion free-free emission jobs queued up by multiprocessing module
- **outQ** (*~multiprocessing.Queue*) – Finished free-free emission jobs

`ChiantiPy.tools.mputil.doIonQ(inQueue, outQueue)`

Multiprocessing helper for *ChiantiPy.core.ion* and *ChiantiPy.core.ion.twoPhoton*

#### Parameters

- **inQueue** (*~multiprocessing.Queue*) – Jobs queued up by multiprocessing module
- **outQueue** (*~multiprocessing.Queue*) – Finished jobs

## ChiantiPy.tools.sources module

Blackbody temperature calculations

**class** ChiantiPy.tools.sources.**blackStar** (*temperature, radius*)

Bases: `object`

Calculate blackbody radiation

### Parameters

- **temperature** (*~numpy.ndarray*) – Temperature in Kelvin
- **radius** (*~numpy.ndarray*) – Stellar radius in cm

### Variables

- **Temperature** (*~numpy.ndarray*) – Temperature in Kelvin
- **Radius** (*~numpy.ndarray*) – Stellar radius in cm
- **Incident** (*~numpy.ndarray*) – Blackbody photon distribution

**incident** (*distance, energy*)

Calculate photon distribution times the visible cross-sectional area.

### Parameters

- **distance** (*~numpy.ndarray*) – Distance to the stellar object in cm
- **energy** (*~numpy.ndarray*) – Energy range in erg

## Notes

This function returns the photon distribution instead of the distribution times the cross-sectional area. Is this correct? Why is the incident photon distribution calculated at all?

ChiantiPy.tools.sources.**blackbody** (*temperature, variable, hnu=1*)

Calculate the blackbody photon distribution as a function of energy (*hnu* = 1) or as a function of wavelength (*hnu* = 0) in units of photons cm<sup>-2</sup> s<sup>-1</sup> str<sup>-1</sup> erg<sup>-1</sup>

### Parameters

- **temperature** (*~numpy.float64*) – Temperature at which to calculate the blackbody photon distribution
- **variable** (*~numpy.ndarray*) – Either energy (in erg) or wavelength (in angstrom)
- **hnu** (*int*) – If 1, calculate distribution as a function of energy. Otherwise, calculate it as a function of wavelength

**Returns** {'photons', 'temperature', 'energy'} or {'photons', 'temperature', 'wvl'}

**Return type** *dict*

## ChiantiPy.tools.util module

Utility functions

### Notes

Some of these functions can be replaced by roman numeral and periodic table lookup libraries. some functions using `os.walk` can be replaced by `os.path`

`ChiantiPy.tools.util.between` (*array*, *limits*)

Find the indices of *array* corresponding to values in the range given by *limits*

`ChiantiPy.tools.util.convertName` (*name*)

Convert ion name string (e.g. 'fe\_13') to atomic number and ion number.

**Parameters** *name* (*str*) – the CHIANTI style name for an ion, such as fe\_14

**Returns** {'Z', 'Ion', 'Dielectronic', 'Element', 'higher', 'lower'} – *higher* and *lower* are the Chianti-style names for the higher and lower ionization stages, respectively.

**Return type** *dict*

`ChiantiPy.tools.util.descale_bt` (*bte*, *btomega*, *f*, *evl*)

Apply excitation descaling of<sup>3</sup> to energy and collision strength

**Parameters**

- *bte* (*array-like*) – Scaled energy
- *btomega* (*array-like*) – Scaled collision strength
- *f* (*array-like*)
- *evl* (*array-like*)

**Returns** [*energy*,*omega*] – Descaled energy and collision strength

**Return type** *list*

### Notes

Need more details here. Not clear which equations are being used.

See also:

`scale_bt()` Apply scaling to excitation energy and cross-section

### References

`ChiantiPy.tools.util.descale_bti` (*bte*, *btx*, *f*, *evl*)

Apply ionization descaling of<sup>9</sup> to energy and cross-sections of<sup>7</sup>.

**Parameters**

- *bte* (*array-like*) – Scaled energy
- *btx* (*array-like*) – Scaled cross-section

<sup>3</sup> Burgess, A. and Tully, J. A., 1992, A&A, 254, 436

<sup>9</sup> Burgess, A. and Tully, J. A., 1992, A&A, 254, 436

<sup>7</sup> Dere, K. P., 2007, A&A, 466, 771,

- **f** (*float*) – Scaling parameter
- **ev1** (*float*) – ionization potential - units determine the output energy units

**Returns** [energy,cross] – Descaled energy and cross-section

**Return type** *list*

## Notes

This is the scaling used and discussed in the Dere (2007) calculation<sup>7</sup> of cross sections. It was derived from similar scalings provided by reference [2]

**See also:**

`scale_bti()` Descale ionization energy and cross-section

## References

`ChiantiPy.tools.util.dilute(radius)`

Calculate the dilution factor.

**Parameters** **radius** (*array-like*) – Distance from the center of a star in units of the stellar radius.

## Notes

If *radius* is less than 1.0, returns 0.

`ChiantiPy.tools.util.e12z(els)`

Convert elemental symbol to atomic number

**Parameters** **els** (*str*) – the abbreviated element name

**Returns** **z** – the atomic number or nuclear charge of the element

**Return type** *int*

`ChiantiPy.tools.util.ion2filename(ions)`

Convert ion name string to generic directory-file name. `convertName` has probably made this redundant

**Parameters** **ions** (*str*) – the CHIANTI style name for an ion, such as `fe_14`

**Returns** **fname** – the full file name of the ion directory in the CHIANTI database assumes a top directory from the environmental variable `XUVTOP`

**Return type** *str*

`ChiantiPy.tools.util.listFiles(dir)`

Walks the path and subdirectories to return a list of files.

**Parameters** **dir** (*str*) – the top directory to search subdirectories are also searched

**Returns** **listname** – a list of files in `dir` and subdirectories

**Return type** *list*



## Notes

This can be replaced by functions in *os.path*, as if 3.4, pathlib is probably better. It is not clear that this function is used anywhere in ChiantiPy

`ChiantiPy.tools.util.listRootNames(dir)`

Walks the path and subdirectories to return a list of file root names.

## Notes

This can be replaced by functions in *os.path*, as if 3.4, pathlib is probably better. Only seems to be used by

`ChiantiPy.tools.util.qrp(z, u)`

Calculate  $Q'_R(Z, u)$ , where  $u = \epsilon/I$  is the impact electron energy in threshold units, from Eq. 2.12 of<sup>4</sup>.

### Parameters

- **z** (*int*) – Atomic number
- **u** (*array-like*) – Impact electron energy in threshold units.

**Returns** **q** – 1s ionization cross-section,  $Q'_R(Z, u)$

**Return type** array-like

## Notes

Used for calculations of direct ionization cross-sections of the H and He sequences in *ChiantiPy.tools.io.twophotonHRead* and *ChiantiPy.tools.io.twophotonHeRead*, respectively.

## References

`ChiantiPy.tools.util.scale_bt(evin, omega, f, ev1)`

Apply excitation scaling of<sup>8</sup> to energy and collision strength.

### Parameters

- **evin** (*array-like*)
- **omega** (*array-like*)
- **f** (*array-like*)
- **ev1** (*array-like*)

**Returns** [**bte**, **btomega**] – Scaled energy and collision strength

**Return type** list

---

<sup>4</sup> Fontes, C. et al., 1999, PhRvA, 59, 1329

## Notes

Need more details here. Not clear which equations are being used.

See also:

`descale_bt()` Descalor excitation energy and cross-section

`ChiantiPy.tools.util.scale_bt_rate(inDict, ip, f=1.7)`

Apply ionization descaling of<sup>7</sup>, a Burgess-Tully type scaling to ionization rates and temperatures. The result of the scaling is to return a scaled temperature between 0 and 1 and a slowly varying scaled rate as a function of scaled temperature. In addition, the scaled rates vary slowly along an iso-electronic sequence.

### Parameters

- **inDict** (*dict*) – the input dictionary should have the following key pairs: *temperature*, array-like and *rate*, array-like
- **ip** (*float*) – the ionization potential in eV.
- **f** (*float* (optional)) – the scaling parameter, 1.7 generally works well

## Notes

*btTemperature* and *btRate* keys are added to *inDict*

`ChiantiPy.tools.util.scale_bti(evin, crossin, f, ev1)`

Apply ionization scaling of<sup>7</sup>,<sup>[8]</sup>, to energy and cross-section.

### Parameters

- **evin** (*float*) – Energy - same units as *ev1*
- **crossin** (*array-like*) – Cross-section
- **f** (*float - the scale factor*)
- **ev1** (*float*) – the ionization potential units - the same as *evin*

**Returns** [**bte**,**btX**] – Scaled energy and cross-section

**Return type** *list*

## Notes

This is the scaling used and discussed in the Dere (2007) calculation [1] of cross sections. It was derived from similar scalings derived in reference [2]

See also:

`descale_bti()` Descalor ionization energy and cross-section

## References

`ChiantiPy.tools.util.scale_classical` (*inDict*, *ip*)

Apply the classical scaling to the input data

### Parameters

- **inDict** (*dictionary*) –  
the input dictionary should have the following key pairs energy and cross or temperature and rate
- **energy** (*array-like*) – energy values of the cross-section
- **cross** (*array-like*) – a cross-section
- **temperature** (*array-like*)
- **rate** (*array-like*)
- **ip** (*float*) – the ionization potential. Typically in eV.
- **Returns** – the following keys are added to inDict
- **\_\_\_\_\_**
- **{‘csEnergy’, ‘csCross’, ‘ip’} or {‘csTemperature’, ‘csRate’, ‘ip’}**

`ChiantiPy.tools.util.spectroscopic2name` (*el*, *roman*)

Convert from spectroscopic notation, e.g. Fe XI to ‘fe\_11’.

### Parameters

- **el** (*str*) – Elemental symbol, e.g. Fe for iron
- **roman** (*str*) – Roman numeral spectroscopic symbol

`ChiantiPy.tools.util.splomDescale` (*splom*, *energy*)

Calculate the collision strength for excitation-autoionization as a function of energy.

### Parameters

- **energy** (*array-like*) – In eV
- **splom** (*dict*) – Structure returned by `ChiantiPy.tools.io.splomRead`

**Returns** **omega** – Collision strength

**Return type** *array-like*

`ChiantiPy.tools.util.z2element` (*z*)

Convert atomic number *z* to its elemental symbol.

**Parameters** **z** (*int*) – The atomic number/nuclear charge

**Returns** **element** – the abbreviated element name

**Return type** *str*

`ChiantiPy.tools.util.zion2dir` (*z*, *ion*, *dielectronic=False*, *xuvtop=""*)

Convert atomic number and ion number to CHIANTI database directory.

### Parameters

- **z** (*int*) – The atomic number/nuclear charge
- **ion** (*int*) – the ionization stage, 1 for the neutral, 2 for the first ionization stage, ...
- **dielectronic** (*bool*, optional)

- **xuvtop** (*str*, optional) – Set different CHIANTI database than the default

**Returns fname** – the CHIANTI directory where the file for the ion specified by *z* and *ion* are found

**Return type** *str*

`ChiantiPy.tools.util.zion2filename(z, ion, dielectronic=False, xuvtop="")`

Convert nuclear charge/atomic number and ion number to CHIANTI database filename.

**Parameters**

- **z** (*int*) – The atomic number/nuclear charge
- **ion** (*int*) – the ionization stage, 1 for neutrals, 2 for singly ionized
- **dielectronic** (*bool*, optional) – whether the ion is the simple dielectronic model
- **xuvtop** (*str*, optional) – the top directory of the CHIANTI database to be used

**Returns fname** – CHIANTI database filename

**Return type** *str*

`ChiantiPy.tools.util.zion2localFilename(z, ion, dielectronic=False)`

Convert atomic number and ion number to generic file name with current directory at top.

**Parameters**

- **z** (*int*) – The atomic number/nuclear charge
- **ion** (*int*) – the ionization stage, 1 for neutrals, 2 for singly ionized
- **dielectronic** (*bool*, optional)

`ChiantiPy.tools.util.zion2name(z, ion, dielectronic=False)`

Convert atomic number and ion number to generic name, e.g. (26,13) to 'fe\_13'

**Parameters**

- **z** (*int*) – The atomic number/nuclear charge
- **ion** (*int*) – the ionization stage, 1 for the neutral, 2 for the first ionization stage, ...
- **dielectronic** (*bool*, optional)

**Returns thisone** – the CHIANTI style ion name, such as 'fe\_13'

**Return type** *str*

`ChiantiPy.tools.util.zion2spectroscopic(z, ion, dielectronic=False)`

Convert atomic number and ion number to spectroscopic notation string

**Parameters**

- **z** (*int*) – The atomic number/nuclear charge
- **ion** (*int*) – the ionization stage, 1 for neutrals, 2 for singly ionized
- **dielectronic** (*bool*, optional)

**Returns spect** – the spectroscopic notation for the ion, such as Fe XIV

**Return type** *str*

## Module contents

Basic tools and utilities used in ChiantiPy

### 4.1.2 Submodules

### 4.1.3 ChiantiPy.version module

the current version of the ChiantiPy package

### 4.1.4 Module contents

ChiantiPy - CHIANTI Python package Calculates various aspects of emission lines and continua from the CHIANTI atomic database for astrophysical spectroscopy.



## INDICES AND TABLES

- `genindex`
- `modindex`
- `search`





## 6.1 Setting default values

Several parameter values can be set to define the way ChiantiPy behaves. The following parameter values can be set in your `~username/.chianti/chiantirc` file. These are:

**wavelength** this parameter has not been fully implemented. The default is angstrom. In the future, nanometers (nm) and kilo volts (kev) will be implemented.

**flux** acceptable values are *energy* and *photon* and these govern emissivities and intensities. the default value is *energy*.

**abundfile** the name of the abundance file. Acceptable values are any of the file names in XUVTOP/abundance, such as *cosmic\_1973\_allen*. The default value is *sun\_photospheric\_1998\_grevesse* which includes the abundances of Grevesse and Sauval, 1998, Space Science Reviews, 85, 161.

**ioneqfile** the name of the ionization equilibrium file. Acceptable values are any of the file names in XUVTOP/ioneq such as *arnaud\_raymond*, *arnaud\_rothenflug*, or *chianti*. The default value is *chianti* which includes the ionization equilibrium calculations of Dere, et al., 2009, Astronomy and Astrophysics, 498, 915 and are considered to be based on the best ionization and recombination rates currently available.

## 6.2 Setting *minAbund* in spectrum calculations

When calculation spectra with *spectrum* or *mspectrum*, it is often useful to set the “minAbund” keyword which governs the minimum abundance of any element included in the calculation. Below is a list of elemental abundances for the elements through zinc and the elements that will be included by several value of “minAbund”.

| Element | Abundance | minAbund |       |       |
|---------|-----------|----------|-------|-------|
|         |           | 1.e-6    | 1.e-5 | 1.e-4 |
| H       | 1.00e+00  | •        | •     | •     |
| He      | 8.51e-02  | •        | •     | •     |
| Li      | 1.26e-11  |          |       |       |
| Be      | 2.51e-11  |          |       |       |
| B       | 3.55e-10  |          |       |       |
| C       | 3.31e-04  | •        | •     | •     |

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Table 1 – continued from previous page

| Element | Abundance | minAbund |       |       |
|---------|-----------|----------|-------|-------|
|         |           | 1.e-6    | 1.e-5 | 1.e-4 |
| N       | 8.32e-05  | •        | •     |       |
| O       | 6.76e-04  | •        | •     | •     |
| F       | 3.63e-08  |          |       |       |
| Ne      | 1.20e-04  | •        | •     | •     |
| Na      | 2.14e-06  | •        |       |       |
| Mg      | 3.80e-05  | •        | •     |       |
| Al      | 2.95e-06  | •        |       |       |
| Si      | 3.55e-05  | •        | •     |       |
| P       | 2.82e-07  |          |       |       |
| S       | 2.14e-05  | •        | •     |       |
| Cl      | 3.16e-07  |          |       |       |
| Ar      | 2.51e-06  | •        |       |       |
| K       | 1.32e-07  |          |       |       |
| Ca      | 2.29e-06  | •        |       |       |
| Sc      | 1.48e-09  |          |       |       |
| Ti      | 1.05e-07  |          |       |       |
| V       | 1.00e-08  |          |       |       |
| Cr      | 4.68e-07  |          |       |       |
| Mn      | 2.45e-07  |          |       |       |
| Fe      | 3.16e-05  | •        | •     |       |
| Co      | 8.32e-08  |          |       |       |
| Ni      | 1.78e-06  | •        |       |       |
| Cu      | 1.62e-08  |          |       |       |
| Zn      | 3.98e-08  |          |       |       |

It should be noted that CHIANTI does not include a complete set of data for every ion of every element in this list.

## INTRO FOR CHIANTI IDL USERS

ChiantiPy was not developed to be a clone of the CHIANTI IDL code. The IDL code largely consists of functions that can be used to calculate or plot a variety of properties. Structures are often used to carry the results of one function to be used by another.

ChiantiPy is object oriented. For example, the building block of ChiantiPy is the **ion** class. It carries with it all the methods that are needed as well as various calculated properties that are kept as attributes. By following the Quick Start guide, you will become familiar with how ChiantiPy works. Because one can always inquire, for example with *dir*, as to the methods and attributes of a an object, such as an ion, it is easy to remember what you might want to calculate next. For example, you have created an object *myion*. It is possible to then invoke

```
myion.popPlot()
```

to plot the level populations of *myion*. If you have not already calculated the levels populations, the **ion** class knows to calculate the level populations (*myion.populate()*) and save them for later use as the dictionary attribute *myion.Population* and then plot the specified level populations. The level populations are then available as a 2 dimensional *numpy* array

```
myion.Population['population']
```

Python and IPython provide many tools for examining any give object (everything in Python is an object of some sort)

```
mg4 = ch.ion('mg_4', setup=0)
```

By setting the keyword *setup* to 0 or *False* the complete setup of the ion is not performed, but a certain amount of information is retrieved. Since neither a temperature or electron density were specified, none of these attributes know anything related to these two quantities. The default value for setup is True so that the setup is performed with the specified temperatures and electron densities. In general, this is the way you want to construct the *mg\_4 ion*.

```
for attr in vars(mg4):
    print(attr)

IonStr
FIP
Z
AbundanceName
Ip
Ion
IoneqAll
PDensity
Abundance
RadTemperature
FileName
RStar
```

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```
ProtonDensityRatio
Dielectronic
Defaults
IoneqName
Spectroscopic
```

The Python function `vars` retrieves the attributes of the `mg4` object.

```
print('%s'%(mg4.IoneqName))

chianti

print('the abundance file name is %s'%(mg4.AbandanceName))

the abundance file name is sun_coronal_1992_feldman_ext

print('the abundance of %s is %10.2e'%(mg4.Spectroscopic,mg4.Abandance))

the abundance of Mg IV is 1.41e-04
```

One can get a more complete description of the various attributes and methods of the `mg4` object

```
for one in dir(mg4):
    print(one)

Abundance
AbundanceName
Defaults
Dielectronic
FIP
FileName
Ion
IonStr
IoneqAll
IoneqName
Ip
PDensity
ProtonDensityRatio
RStar
RadTemperature
Spectroscopic
Z
__class__
__delattr__
__dict__
__dir__
__doc__
__eq__
__format__
__ge__
__getattr__
__gt__
__hash__
__init__
__le__
__lt__
__module__
```

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```

__ne__
__new__
__reduce__
__reduce_ex__
__repr__
__setattr__
__sizeof__
__str__
__subclasshook__
__weakref__
boundBoundLoss
circularDescale
convolve
diCross
diRate
drRate
drRateLvl
eaCross
eaDescale
eaRate
emiss
emissList
emissPlot
emissRatio
gofnt
intensity
intensityList
intensityPlot
intensityRatio
intensityRatioInterpolate
intensityRatioSave
ionGate
ioneqOne
ionizCross
ionizRate
lineSpectrumPlot
p2eRatio
popPlot
populate
recombRate
rrRate
setup
setupIonrec
spectrum
spectrumPlot
twoPhoton
twoPhotonEmiss
twoPhotonLoss
upsilonDescale
upsilonDescaleSplups

```

First, at the top of the list are the attributes that were listed by the *vars* function. Then come a number of methods starting with ‘\_\_’. These are generally not used and called *private* methods although nothing in Python is really private. In IDL, just about everything is private. After the *private* methods is a list of the methods provided by the *mg4 ion* class object. These all start with a lower case letter to separate them from the attributes that start with an upper case letter (this is a ChiantiPy convention).

In the IPython and jupyter-qtconsole, typing

```
mg4.diCross(
```

and then hitting the tab key will bring up the doc-string for the diCross method, also found in the API reference in the documentation. And then

```
mg4.diCross()
```

calculates the direct ionization cross section of Mg IV for a set of energies above the ionization potential  $I_p$ . The direct ionization cross sections are then provided in the mg4.DiCross dictionary.

Experience using the CHIANTI IDL package will provide the user with a background with what ChiantiPy can do. However, the way to accomplish them are much easier but must be learned. The best way to start is with the Quick Start guide and a book about Python. Book suggestions are *Learning Python* by Mark Lutz and the handy *Python Pocket Reference*, also by Mark Lutz. The first one is always in reach and copies of the latter is on all of my computer desks.

## RESOURCES FOR CHIANTIPY USERS

### 8.1 Forums

The CHIANTI Google group

Github





## CHANGELOG

### 9.1 Changes from 0.9.4 to 0.9.5

this is a bug-fix release.

a bug in the inherited method `base._IntensityRatio()` had a problem if lines were selected from different ions

### 9.2 Changes from 0.9.3 to 0.9.4

this is a bug-fix release.

changes in version 0.9.2 continued to give problems with ions that included autoionization rates

### 9.3 Changes from 0.9.3 to 0.9.3

this is a bug-fix release.

changes in version 0.9.2 gave problems with ions that included autoionization rates

### 9.4 Changes from 0.9.2 to 0.9.3

this is a bug-fix release.

changes in version 0.9.2 led to an error where `ion.recombRate` did not work. This has been fixed

### 9.5 Changes from 0.9.1 to 0.9.2

this is a bug-fix release.

changes in version 0.9.1 lead to an error where a bare ion has not recombination rate. This has been fixed

## 9.6 Changes from 0.9.0 to 0.9.1

this is a bug-fix release.

in cases when it is not possible to calculate the free-bound continuum for some ion, mspectrum did not handle this correctly and crashed

also, the ion zn\_31 (Zn XXXI) is a bare ion and has no ionization potential (IP) and looking it up caused indexing errors.

## 9.7 Changes from 0.8.7 to 0.9.0

a new module model.maker has been added

```
import ChiantiPy.model as mdl
mymodel = mdl.maker(...)
```

a serious bug in ch.freeBound was fixed - the use of a single temperature was problematic

## 9.8 Changes from 0.8.6 to 0.8.7

continued code cleanup

## 9.9 Changes from 0.8.5 to 0.8.6

another bug-fix release

added argCheck method to make sure that sizes of temperature, density and emission measure were compatible

## 9.10 Changes from 0.8.4 to 0.8.5

## 9.11 This is a major bug-fix release.

Errors in calculating the proton rates were corrected.

All temperatures and densities are now numpy arrays

## 9.12 Changes from 0.8.3 to 0.8.4

## 9.13 This is a major bug-fix release.

Another significant bug was fixed in the important ion.populate method.

## 9.14 Changes from 0.7.1 to 0.8.3

## 9.15 This is a major bug-fix release.

a small but mighty bug was found in the important `ion.populate` method.

## 9.16 Version 0.8.x files are necessary to use with the new CHIANTI Version 9.0 database

Changes have been made to take into account the new way that CHIANTI is handling dielectronic recombination and autoionization

The release is also available on [PyPI](<https://pypi.org/project/ChiantiPy/>)

Documentation is available on [github.io](<https://chianti-atomic.github.io/>)

and on [ReadTheDocs](<https://chiantipy.readthedocs.io/en/latest/?badge=latest>)

## 9.17 changes from 0.7.1 to 0.8.0

ChiantiPy is now only compliant with Python 3. Development is currently with Python 3.6

The use of the PyQt4 and WxWidgets packages have been dropped and PyQt5 is now used

The documentation is now available on [github.io](https://github.io) and [ReadTheDocs](https://readthedocs.io)

## 9.18 changes from 0.7.0 to 0.7.1

version 0.7.0 included some changes in the ChiantiPy naming conventions, largely in the continuum class. These are being reverted to the original ChiantiPy naming conventions.

the `ion.freeBoundxxx` methods have been fixed and this also fixes the problem with the `RadLoss` class.

a pseudo-voigt filter has been added to `tools.filters`

the keyword argument `wvlRange` has been removed from the `ion.emiss` and `ion.intensity` methods

the keyword argument for the Emission Measure, `em`, has been removed from the `ion.intensity` and similar methods. It is now necessary to specify the emission when the object is instantiated.

a set of PyQt5 dialogs have been developed by **ktritz** and are now included

this is the last release that will use the PyQt4 widgets as an option.

the method **ioneqOne** is used by both the `Ion` and `Continuum` class. It has been moved to a single `_IoneqOne.py` file in the **base** directory

## 9.19 changes from 0.6.5 to 0.7.0

The primary change is that code development has been moved to [Github](#).

Also, in order to be more compliant with other astrophysical packages on Github ([Astropy](#) and [SunPy](#)) the directory layout has been changed and renamed.

The core routines are now imported as

```
import ChiantiPy.core as ch
```

this give access to `ch.ion`, `ch.spectrum`, etc.

In terms of bug-fixes, the calculation of excitation-autoionization cross-sections and rates have been corrected in the `eaCross()` and `eaRate()` methods

Current development is with Python 3.4

## 9.20 changes from 0.6.0 to 0.6.5

`matplotlib.pyplot` is now imported for plotting

IPython version 4 / Jupyter is now listed as a prerequisite. However, v0.6.4 can be made compatible with IPython 2 or 3 with a simple edit.

An error in calculating the proton excitation rates was fixed.

The code has been edited to make it compatible with Python 3 and has been tested against Python 3.3

## 9.21 changes from 0.5.3 to 0.6.0

This is a major release.

First, ChiantiPy 0.6.0 is compatible with the most recently released CHIANTI database version 8.0. It also fixes some major bugs in the previous version. Documentation has been improved and a IPython notebook **QuickStart.ipynb**, that largely follows the ‘Quick Start’ documentation pages, has also been included.

There are two new multi-ion classes: **bunch** and **ipyspectrum**. **bunch** allows the user to calculate line intensities for a specified set of elements or individual ions as a function of temperature or density. One advantage of **bunch** is the ability to calculate the intensity ratio of lines of two different ions as a function of temperature or density.

**ipyspectrum** is much like the existing **spectrum** and **mspectrum** classes. **mspectrum** allows the use of the Python **multiprocessing** module to speed up spectral calculations. The **ipyspectrum** class uses the IPython **parallel** module so that multiprocessing spectral calculations can be performed in the IPython QtConsole and Notebook.

A new method **intensityList** has been developed to allow the user to list the most intense lines within a given wavelength range. This new methods, together with previously existing **intensityRatio** and **intensityRatioSave** are all now inherited by the **ion** classes and the multi-ion classes.

The **ion** and multi-ion classes now accept the keyword argument **abundanceName** that allow the user to specify the set of elemental abundances rather than just the default abundance file.

Additional we have replaced the FortranFormat module of Scientific Python by Konrad Hinsén with the **fortran-format** module of Brendan Arnold at <http://bitbucket.org/brendanarnold/py-fortranformat>. I have slightly modified `fortranformat` to make it Python 3 compliant.

For the future, I plan to make ChiantiPy compliant with both Python 2.7 and the current version of Python 3 (now 3.4), improve the documentation and move the projec to github, in no particular order.

ChiantiPy is now released under a new license, the OSI approved ISCL license. From [Wikipedia](#) *The ISCL license is a permissive free software license written by the Internet Software Consortium (ISC). It is functionally equivalent to the simplified BSD and MIT/Expat licenses, ...*



## **REPORTING BUGS**

All bugs are kept track of on the [GitHub issue tracker](#). If you run into any unexpected behavior or have a question please [send an email](#) or add an issue directly to the issue tracker.





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